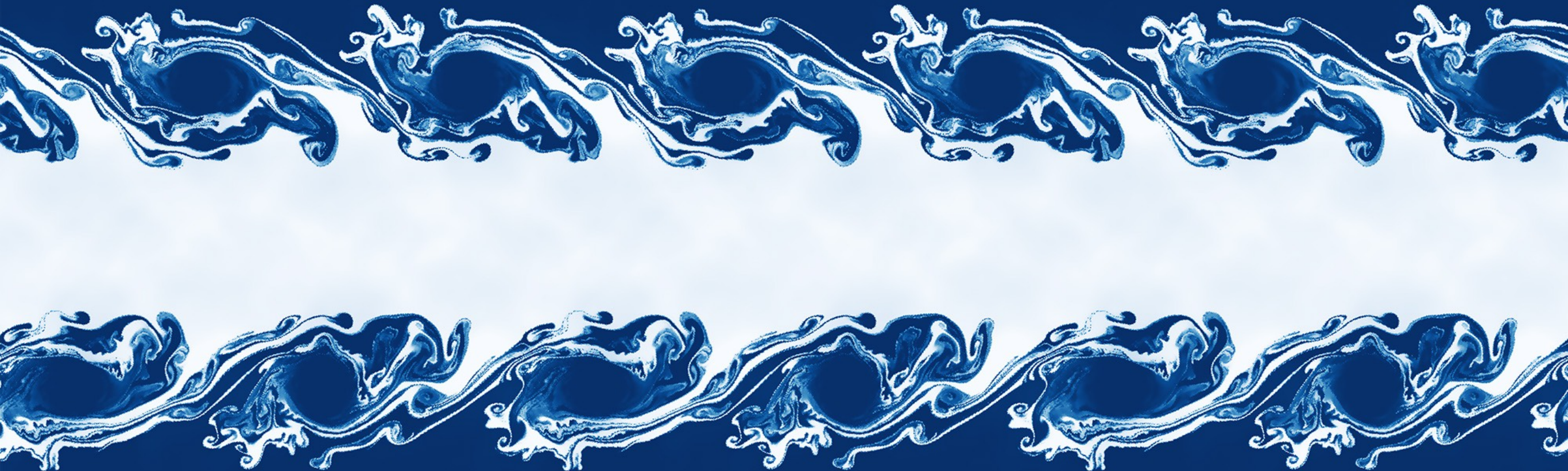


ASTR 670: Interstellar medium and gas dynamics

Prof. Benedikt Diemer

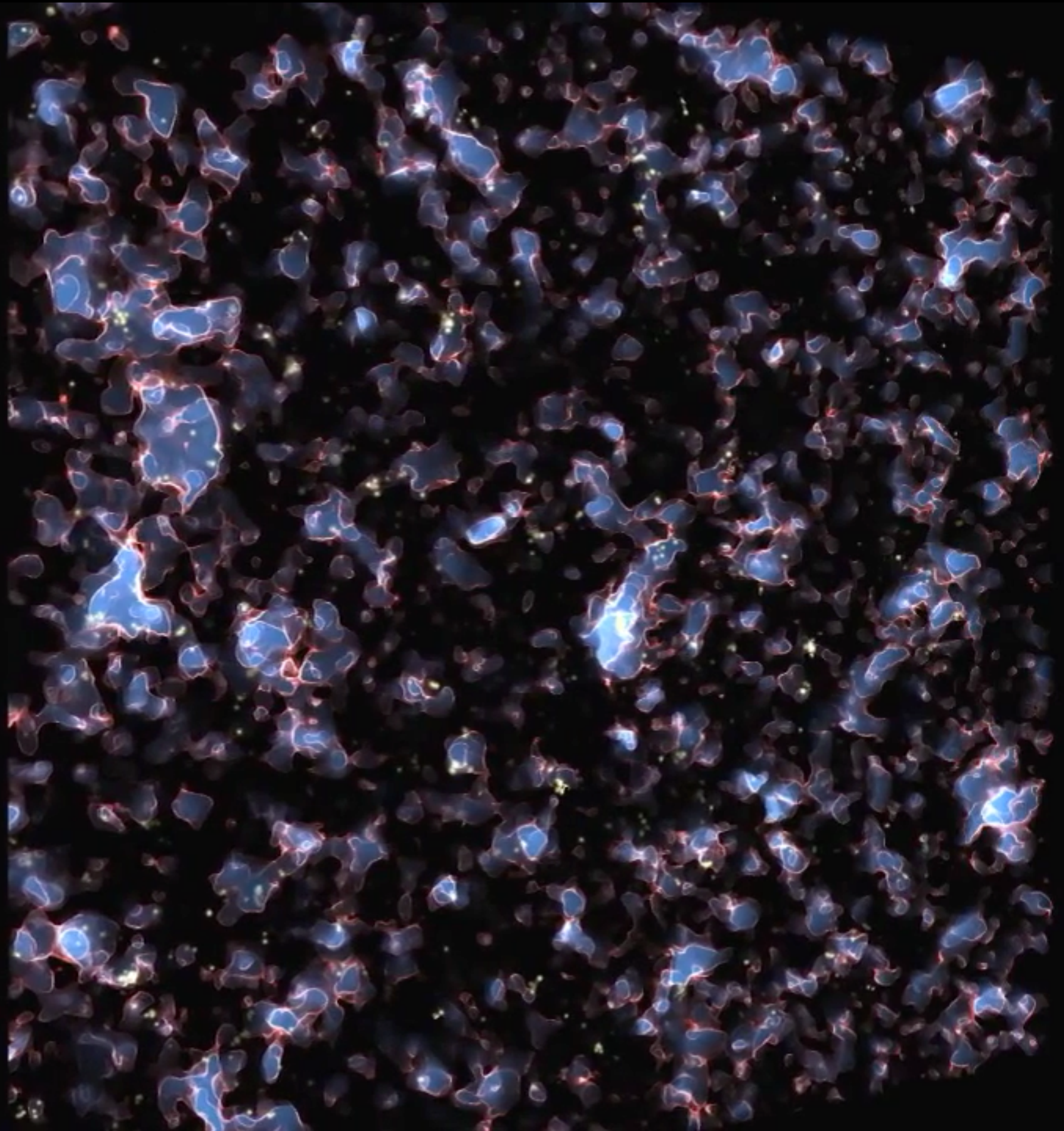


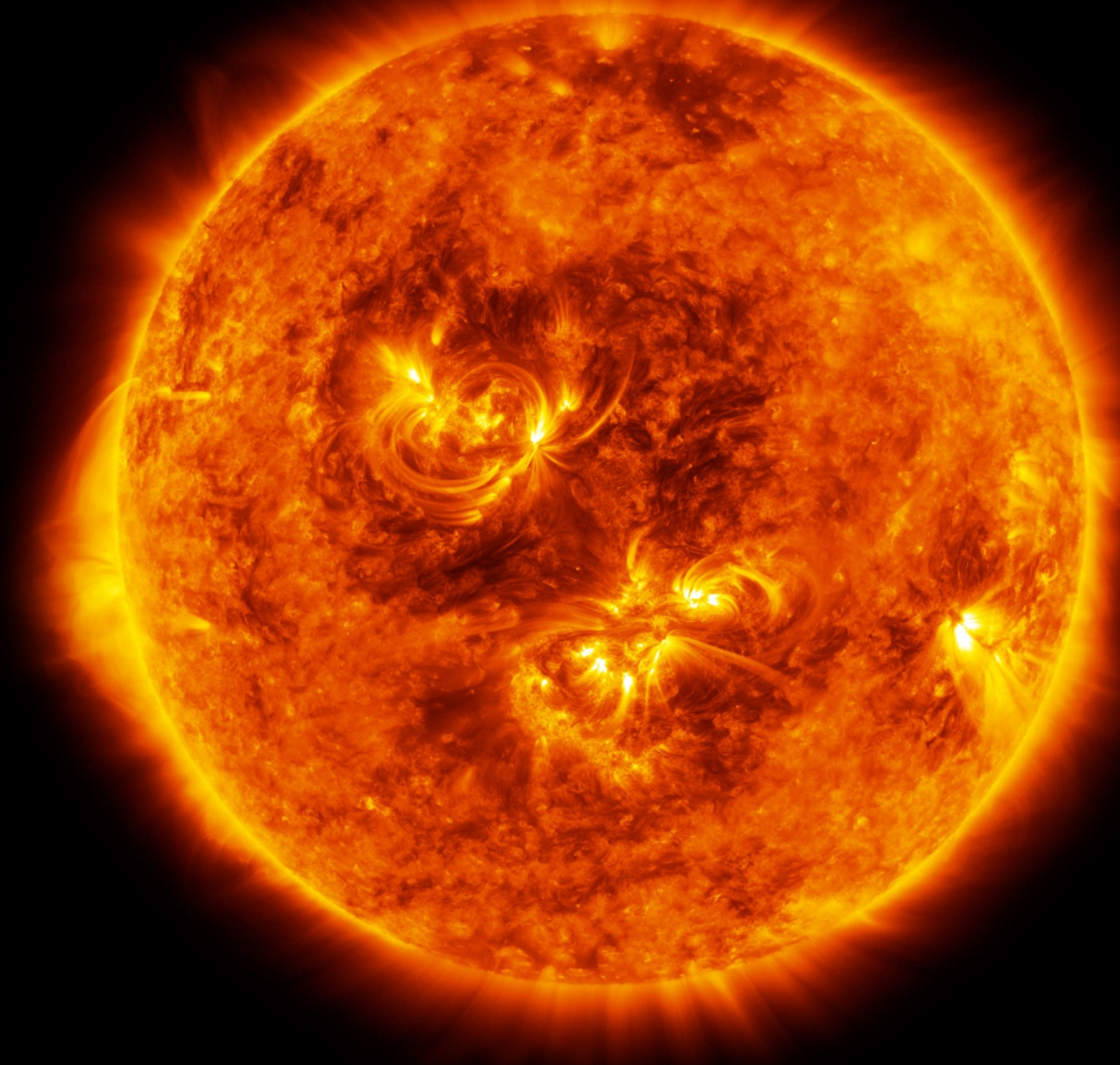
Chapter 10 • Magnetohydrodynamics

“The larger one's ignorance,
the stronger the magnetic field”

Lodewijk Woltjer

Reionization





MHD conservation equations

Ideal MHD equations in conservation-law form

$$\frac{\partial \rho}{\partial t} + \nabla \cdot (\rho \mathbf{u}) = 0$$

$$\frac{\partial(\rho \mathbf{u})}{\partial t} + \nabla \cdot \left(\rho \mathbf{u} \otimes \mathbf{u} + \mathbf{I} \left[P + \frac{\mathbf{B}^2}{2} \right] - \mathbf{B} \otimes \mathbf{B} \right) = -\rho \nabla \Phi$$

$$\frac{\partial E}{\partial t} + \nabla \cdot \left(\left[E + P + \frac{\mathbf{B}^2}{2} \right] \mathbf{u} - \mathbf{u} \cdot \mathbf{B} \otimes \mathbf{B} \right) = \rho \frac{\partial \Phi}{\partial t} + \mathcal{H} - \mathcal{C}$$

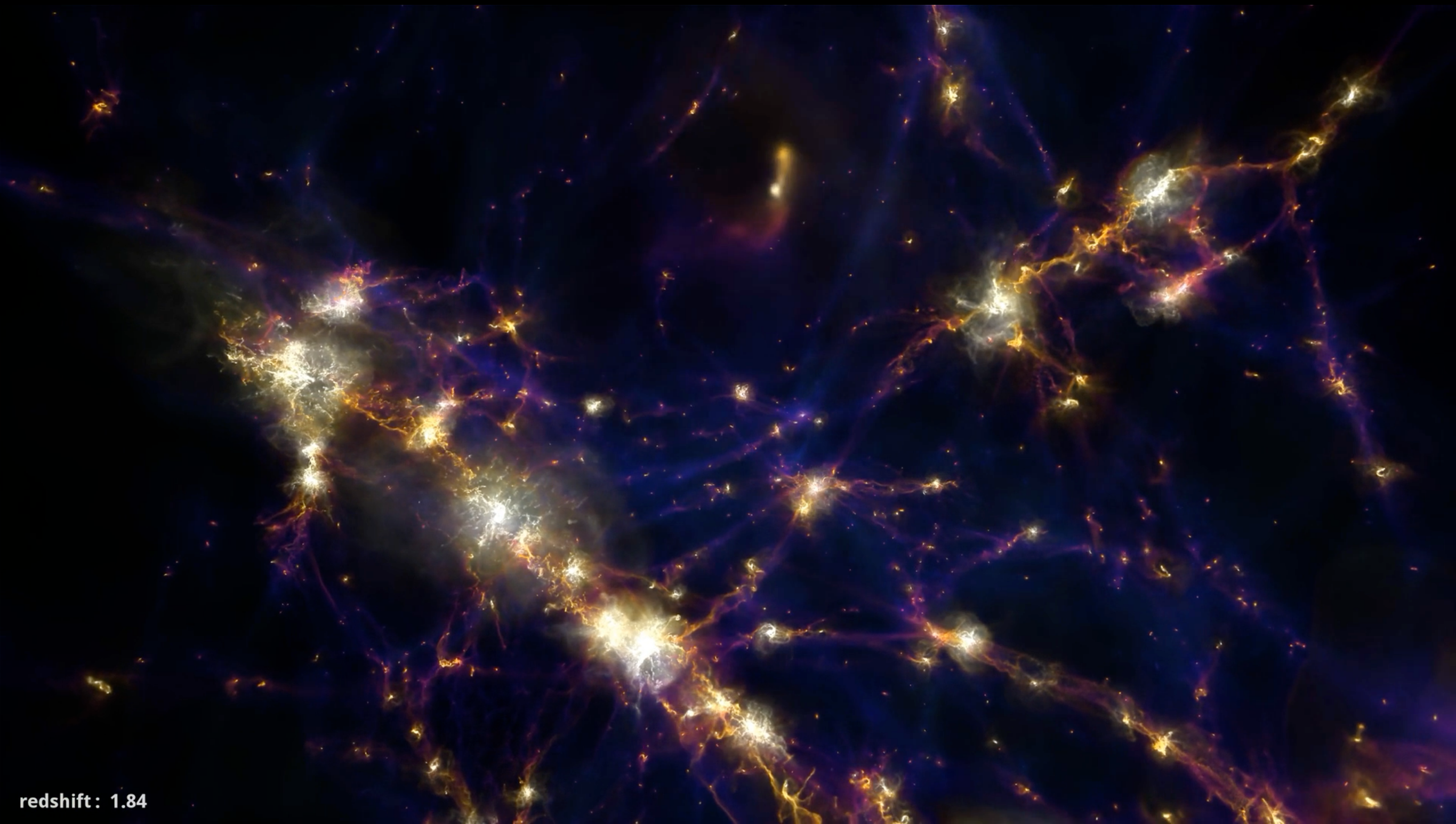
$$\frac{\partial \mathbf{B}}{\partial t} + \nabla \cdot (\mathbf{u} \otimes \mathbf{B} - \mathbf{B} \otimes \mathbf{u}) = 0$$

$$E \equiv \rho \left(\frac{\mathbf{u}^2}{2} + \varepsilon + \frac{\mathbf{B}^2}{2} + \Phi \right)$$

The Riemann problem (1D Euler)

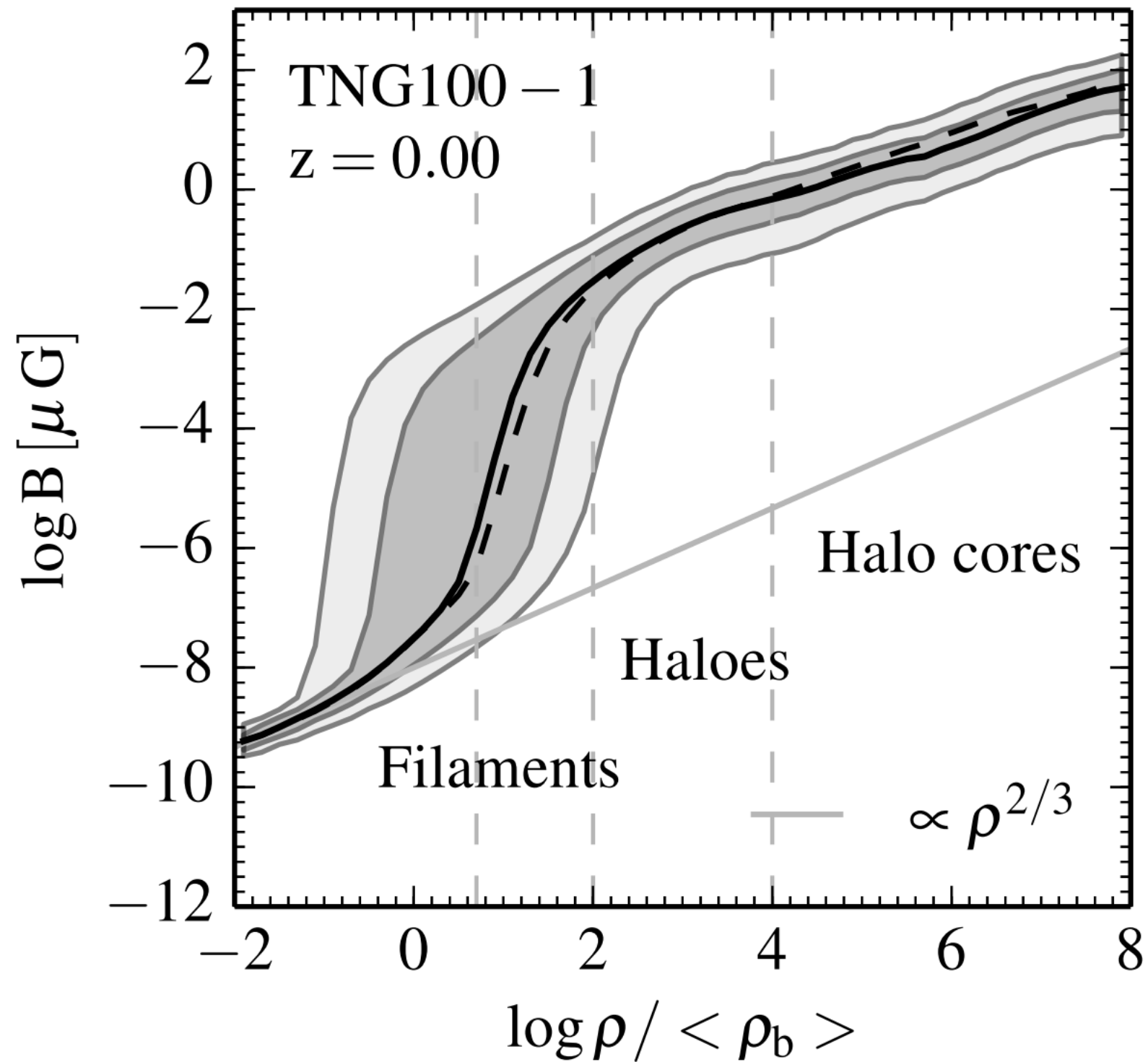
System	n (cm ⁻³)	T (K)	B (G)	β
Magnetosphere of the Earth	10^4	10^4	0.3	4×10^{-6}
Coronal loop in the Sun	10^{10}	10^6	300	0.0004
Solar wind	10	10^5	6×10^{-5}	1
Molecular cloud	100	20	5×10^{-6}	≈ 0.1
Warm HI in interstellar medium	1	10^4	10^{-5}	≈ 0.1
Galaxy cluster	10^{-3}	10^7	$10^{-6} - 10^{-5}$	≈ 1

Magnetic fields in Illustris TNG100

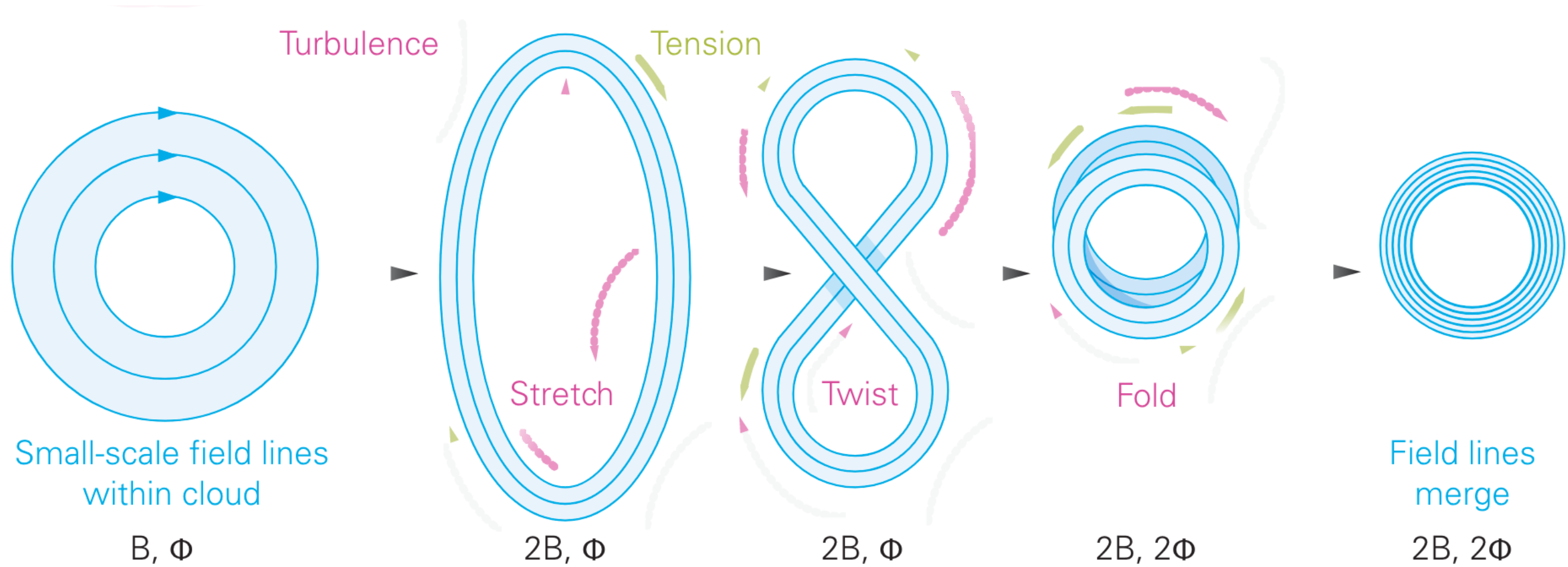


redshift : 1.84

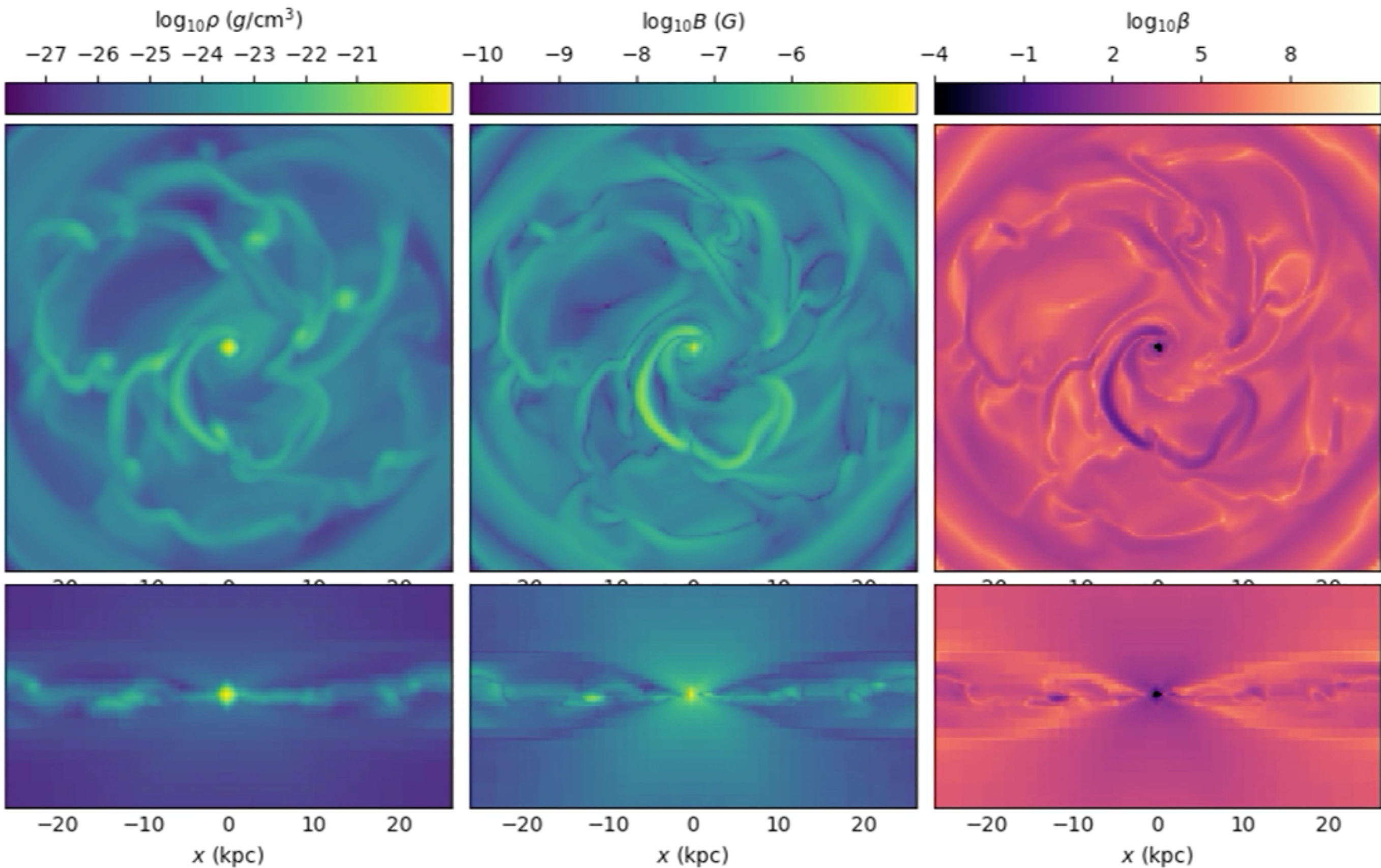
The Riemann problem (1D Euler)



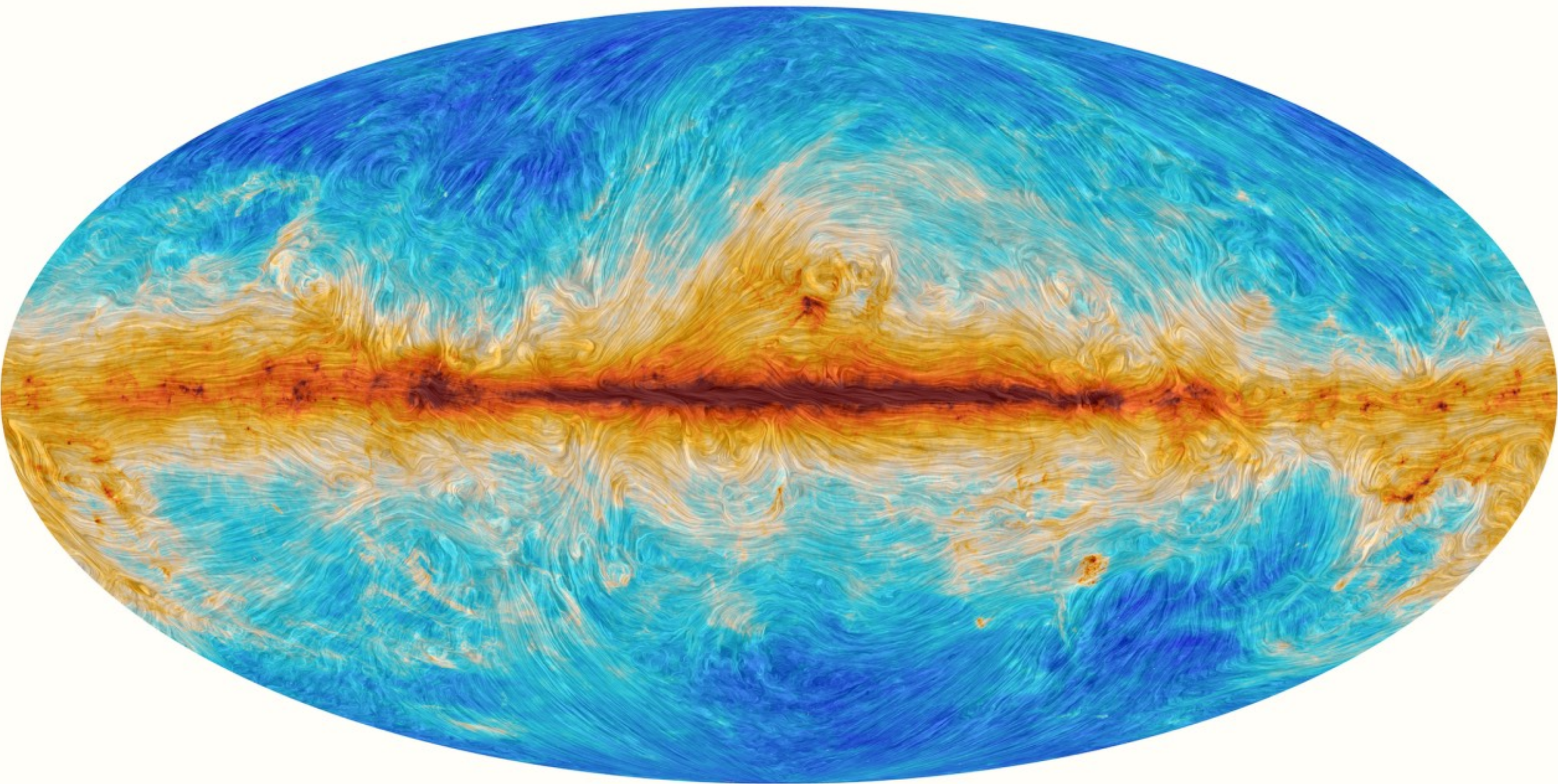
Galactic dynamo



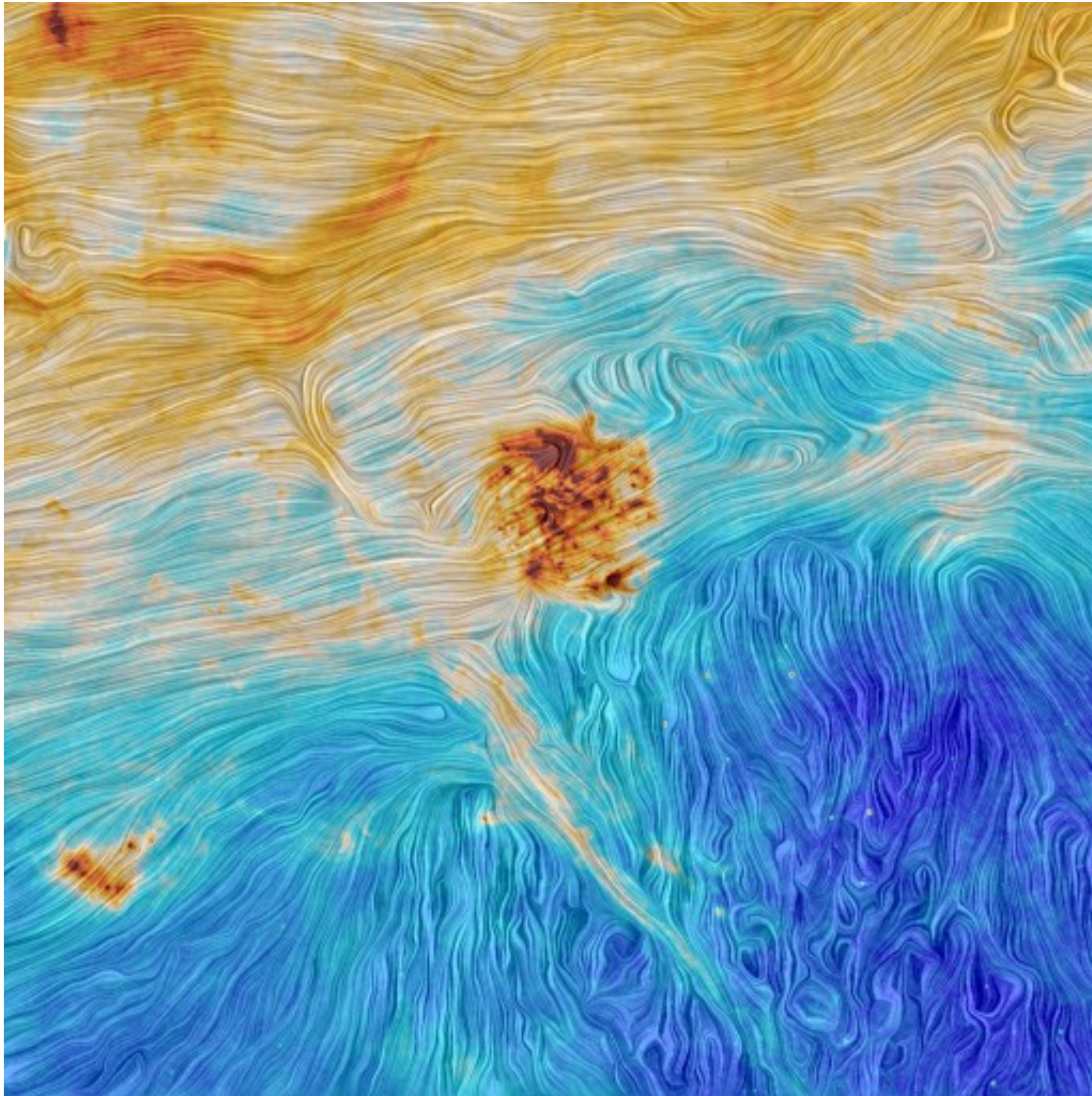
Galactic dynamo



The galactic magnetic field



The galactic magnetic field



Reading

- Recommended: CC §13.1-4
- Additional: Shu §21, §22