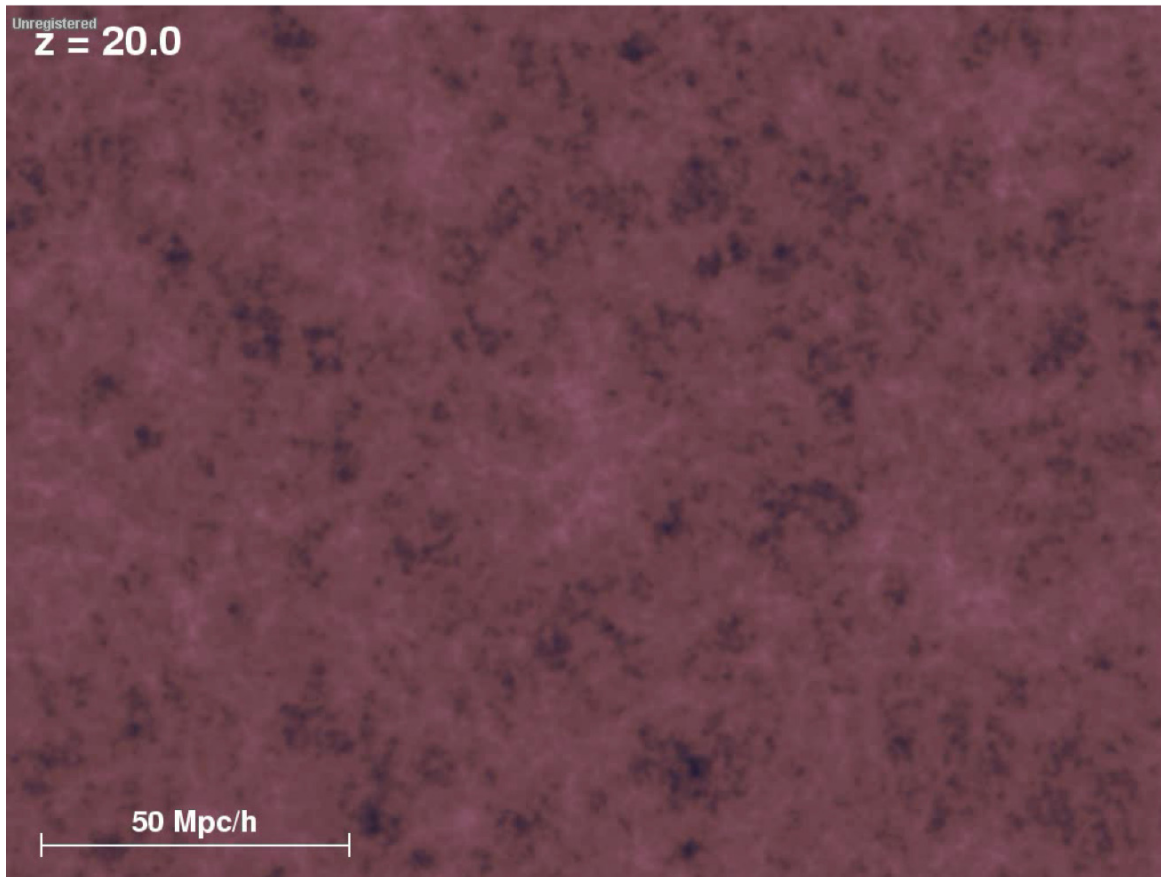


Gas and Star Formation in Galaxies: the Nearby and the Distant Universe

Prof. Alberto Bolatto

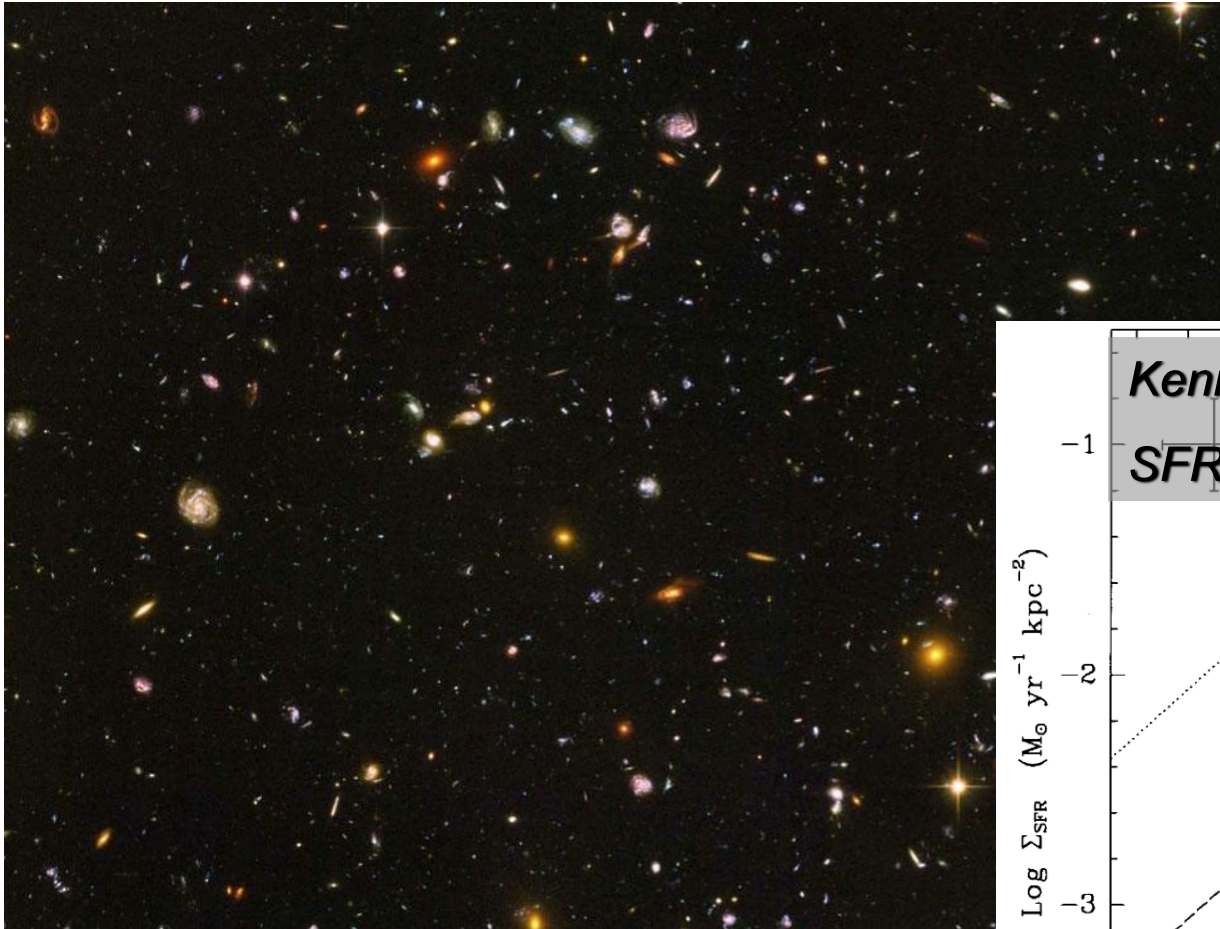
Motivation: understanding structure formation



*V. Springel et al.,
Millenium Simulation*

Dark matter (gravity)

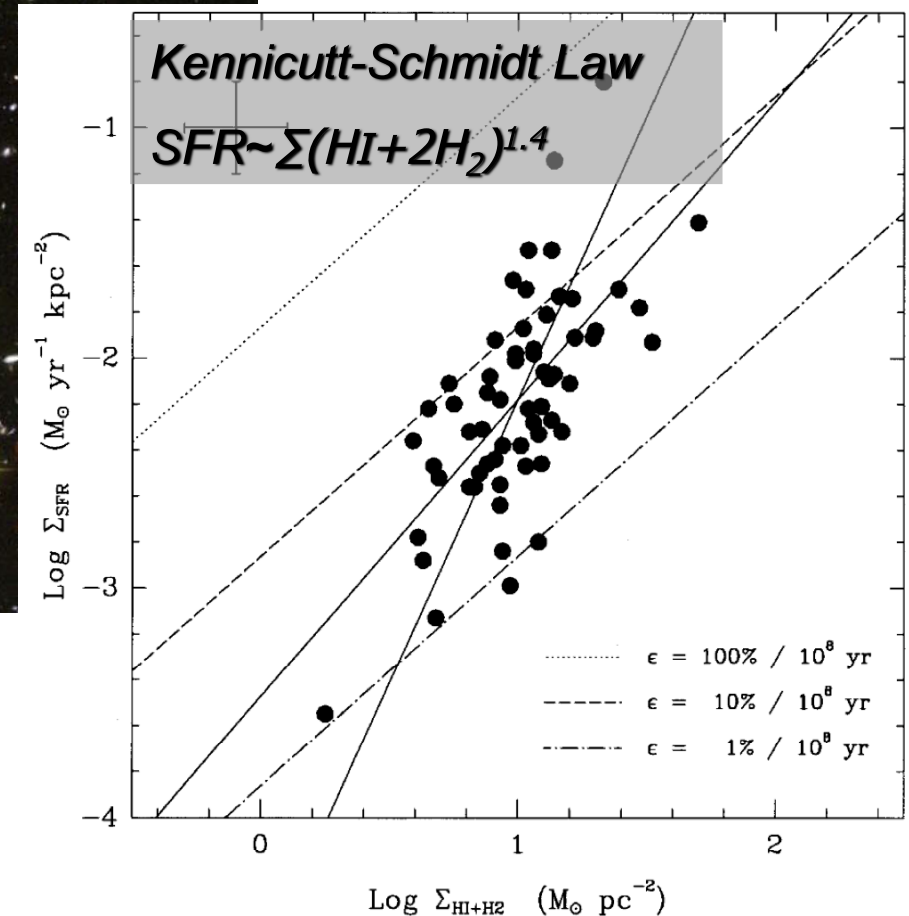
Structure formation in the Universe: Light



Hubble UDF

Kennicutt (1989, 1998)

*Dark matter (gravity) $\rightarrow$ HI $\rightarrow$ H₂
 $\rightarrow$ stars $\rightarrow$ observable structure
(astrophysics)*

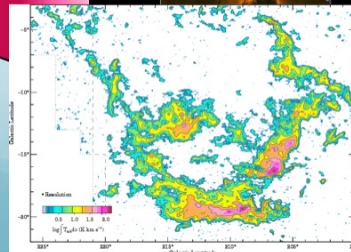
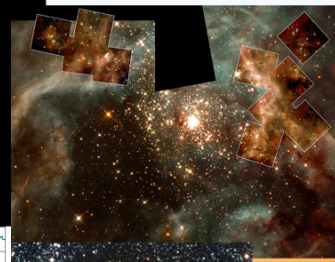
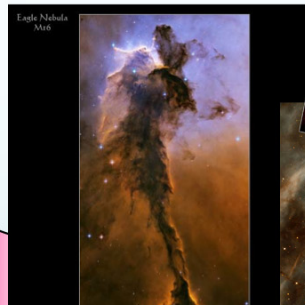
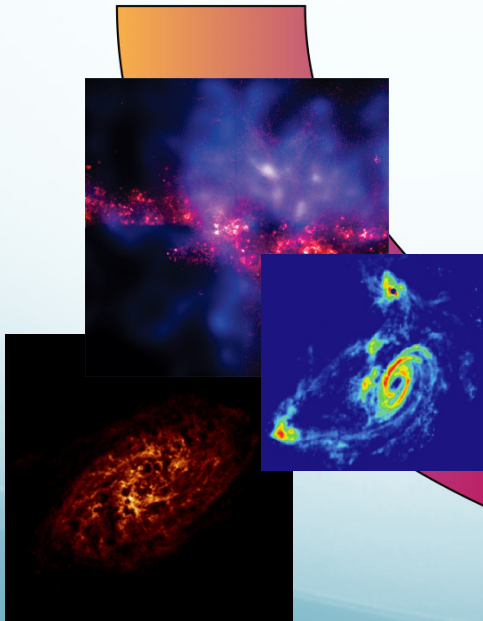
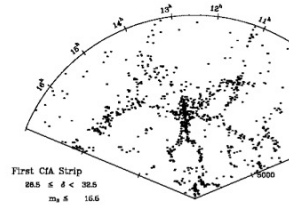
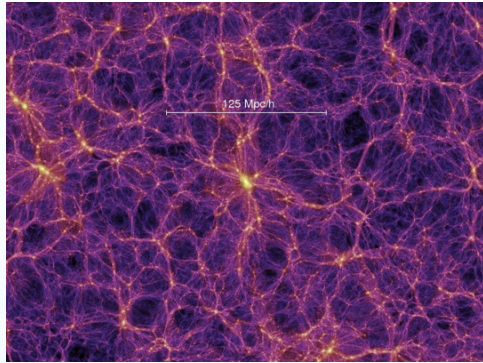


The evolution of structure in the universe

Stars form in Giant Molecular Clouds (GMCs)

Molecular transitions are necessary to radiate the heat of contraction

HI is the dominant gas reservoir in galaxies, but star formation occurs only in H_2



Total gas or molecular H_2

S³MC+SMC-SAGE

Spitzer 24, 70, 160 μm
imaging

Bolatto et al. (2007), Gordon et al.
in prep.

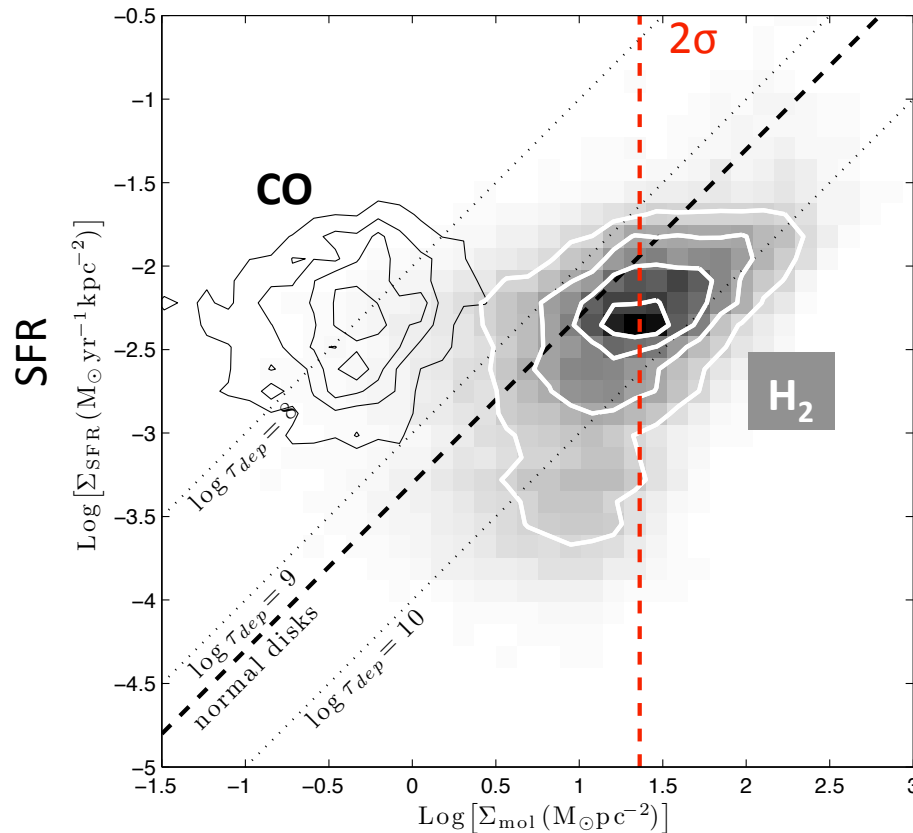
- Better
- We use H_2 by bypass the CO and Xco using dust as the H_2 tracer (Israel 1997, Dame et al. 2001)

Leroy, Bolatto, et al. (2007, 2009, 2010); Bolatto, Leroy, et al. in prep.

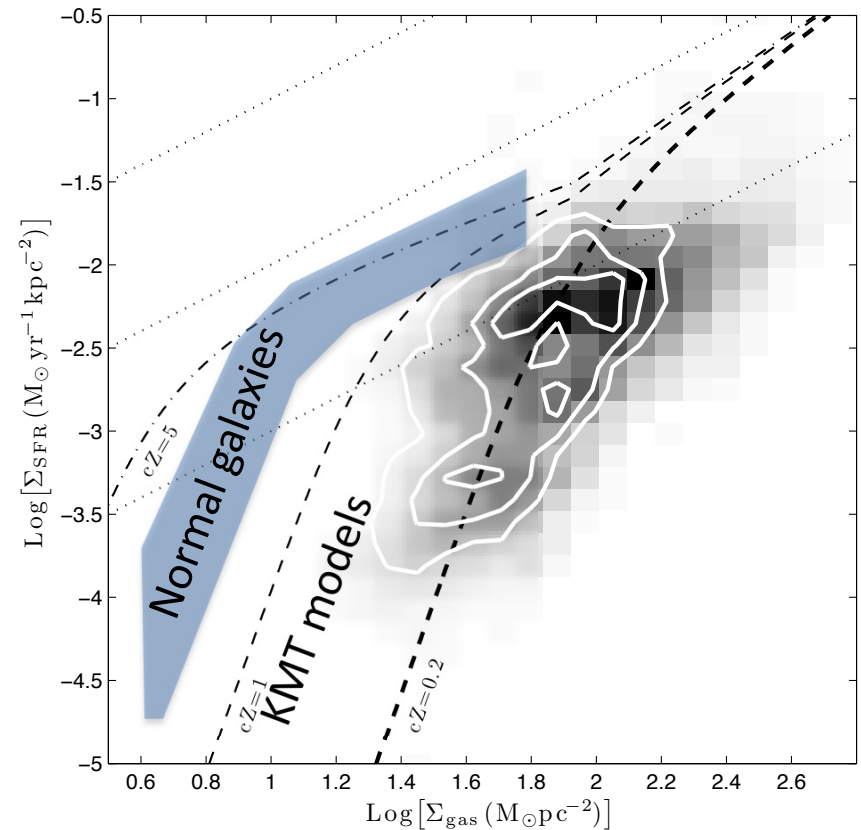


Relations for H₂ and total gas

Molecular Gas

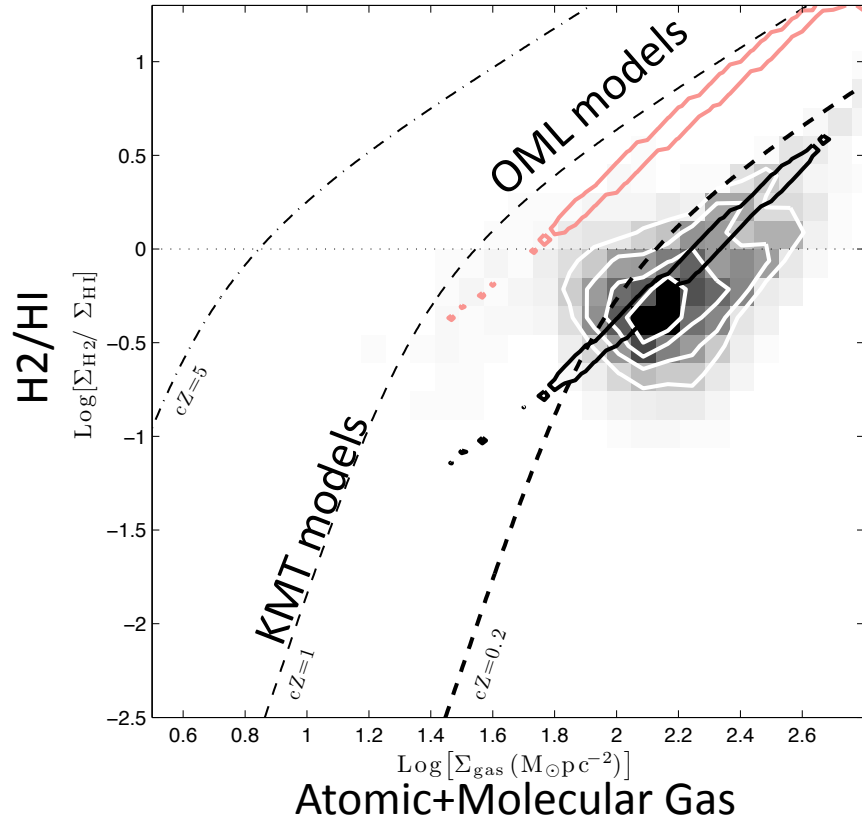


Atomic+Molecular Gas



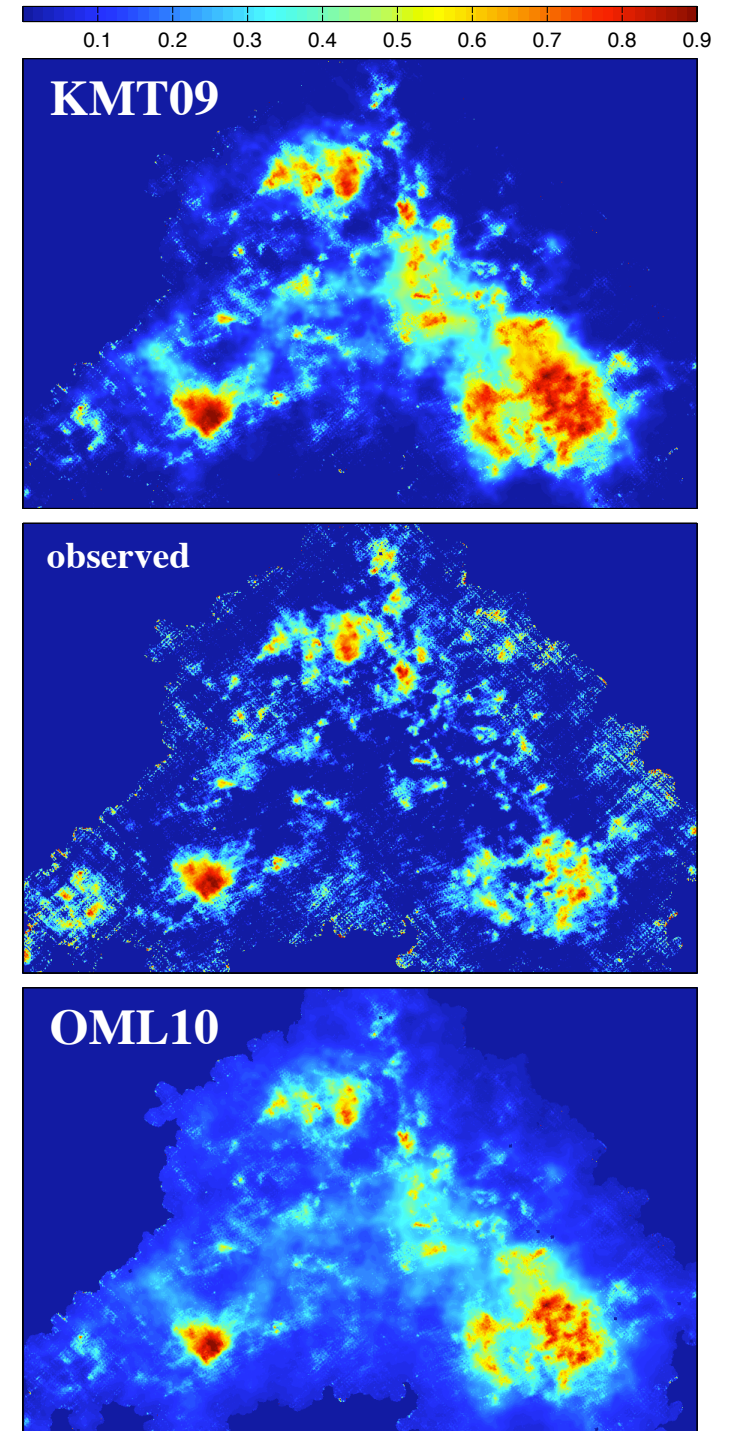
- The SMC is “normal” in H₂ vs. SFR
- The SMC is vastly underperforming in total gas vs. SFR (explains Wolfe & Chen 2006)
- KMT09 models where $\text{HI} \rightarrow \text{H}_2 \rightarrow \text{SFR}$ fit the data reasonably well

Molecular fraction and pressure

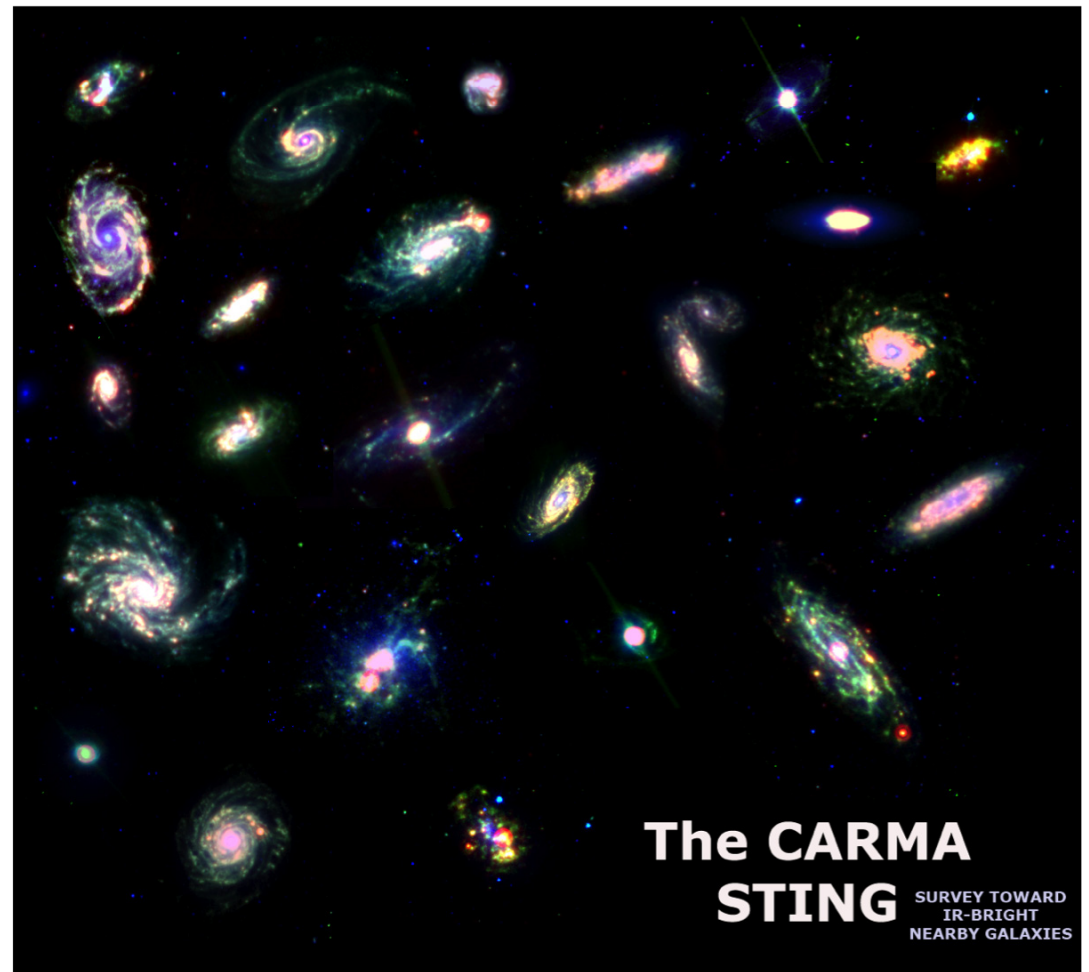
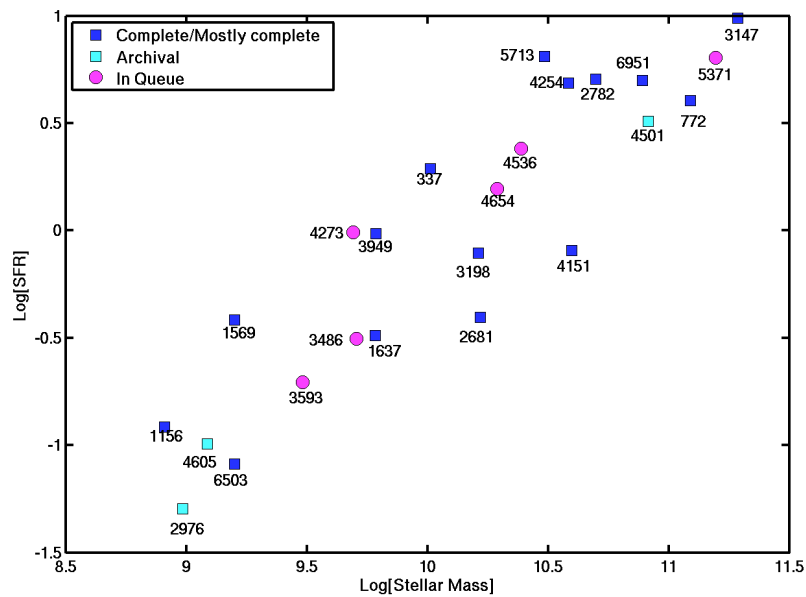


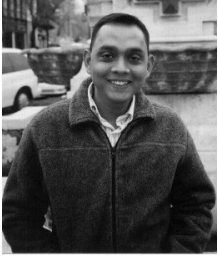
- Data are fit extremely well by OML10, after accounting for the effect of Z_d on UV propagation
- The core of OML10 is thermodynamic pressure equilibrium. It reproduces the pressure relation of BR06 very well

Elmegreen (1993); Blitz & Rosolowsky (2004,2006); KMT09; Ostriker, McKee, & Leroy 2010

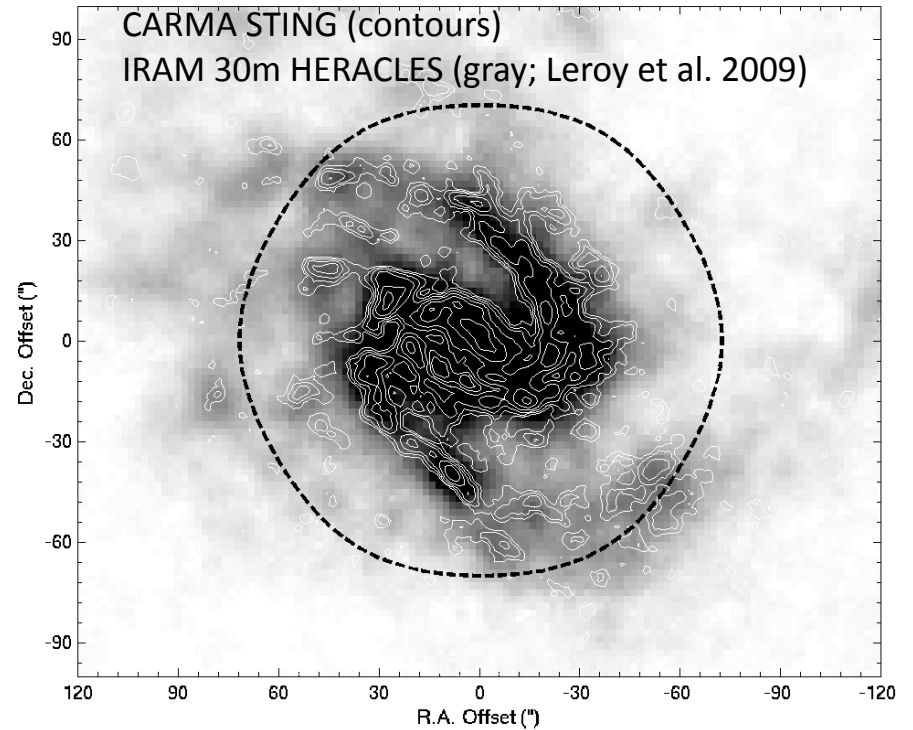
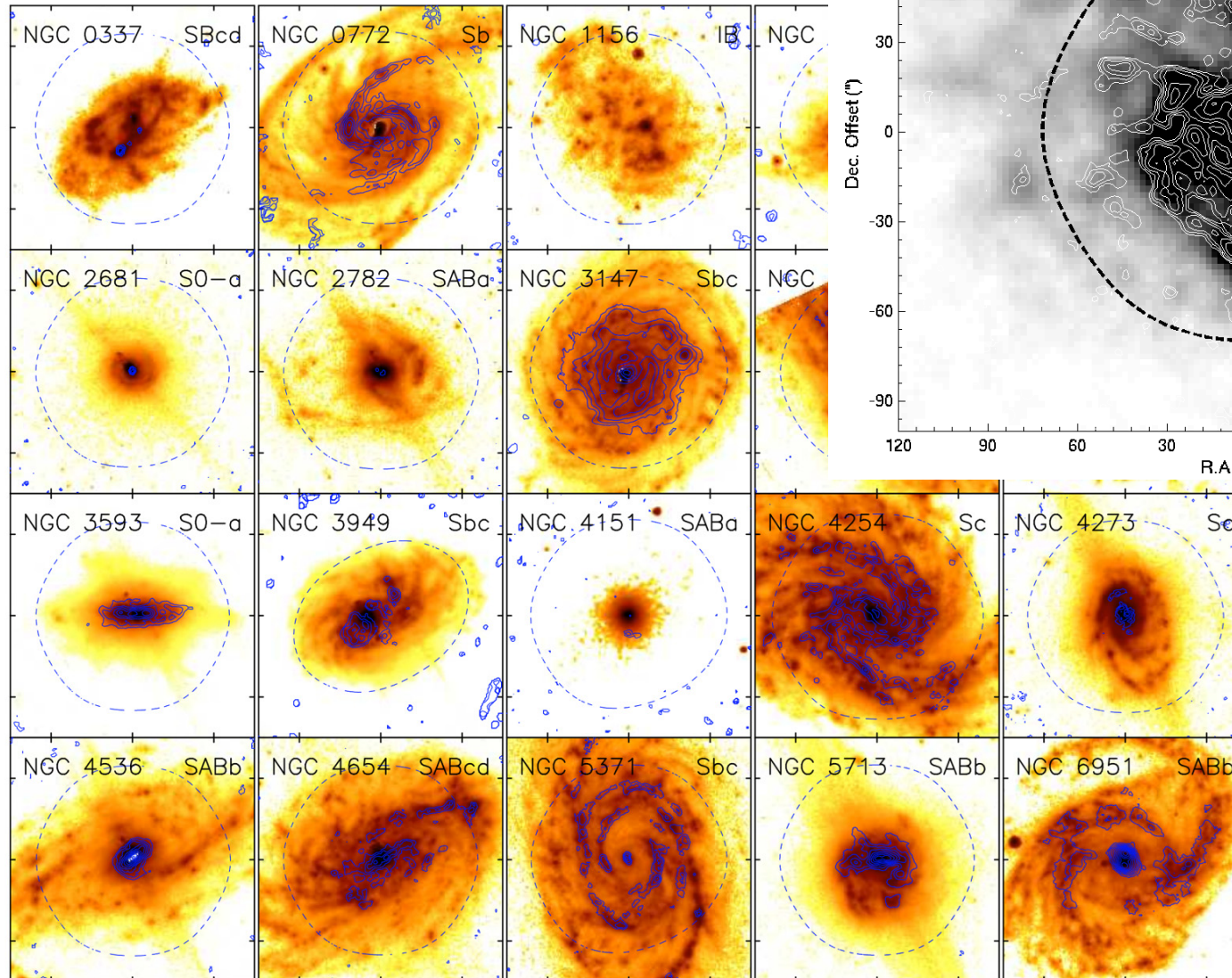


Molecular Gas in the Blue Sequence





STING



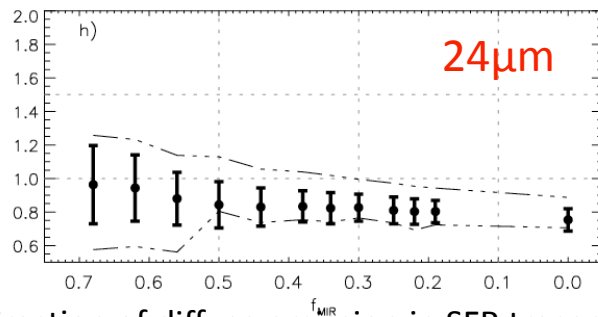
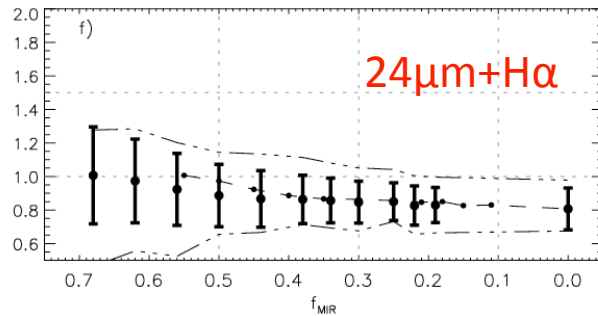
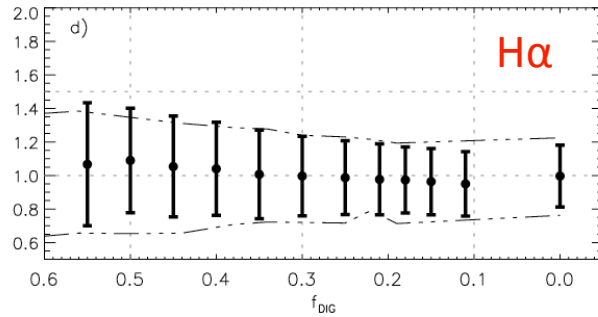
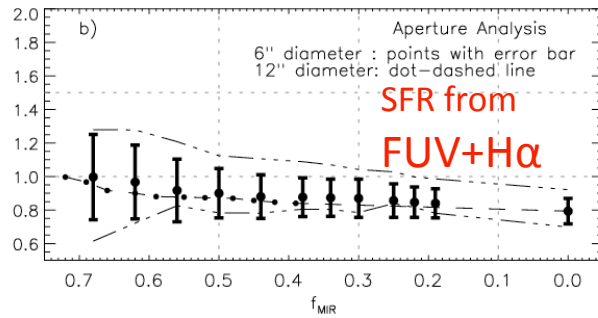
Is the relation
between gas and
SFR nonlinear or
linear?

(Rahman, Bolatto, Wong,
Leroy, Rosolowsky, Blitz,
et al. 2010)

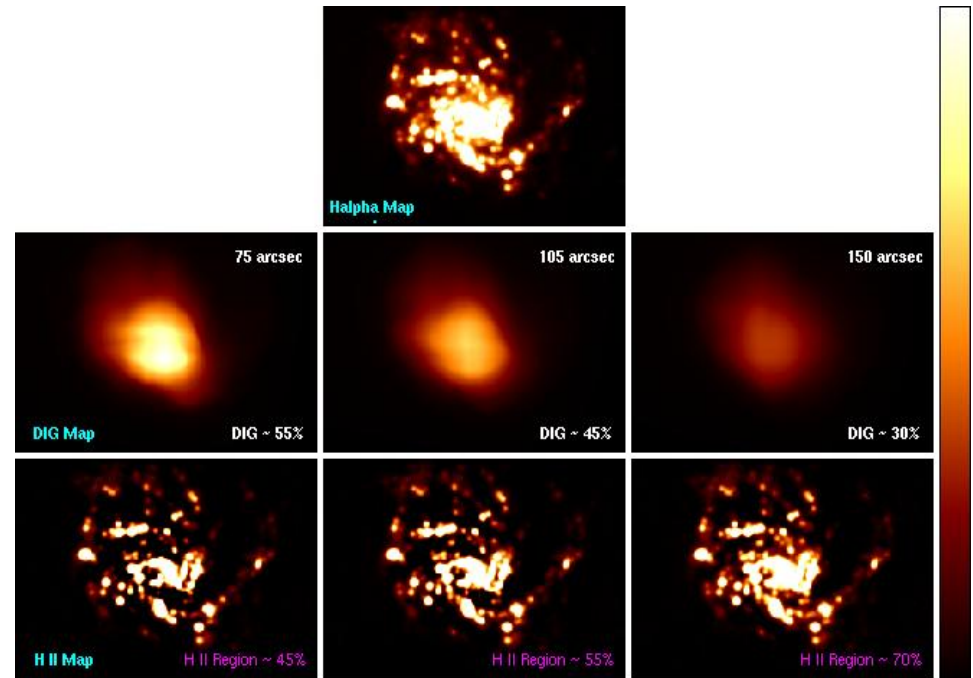
The Star Formation Law in NGC 4254

(Rahman et al., 2010)

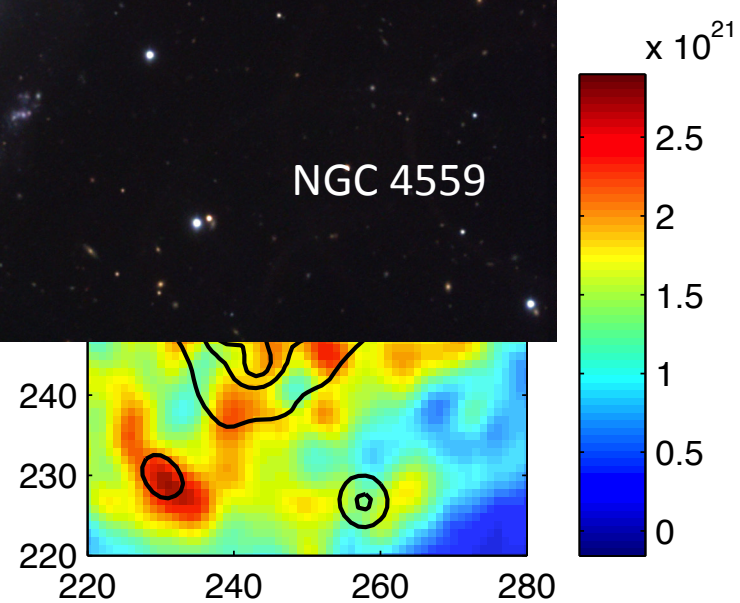
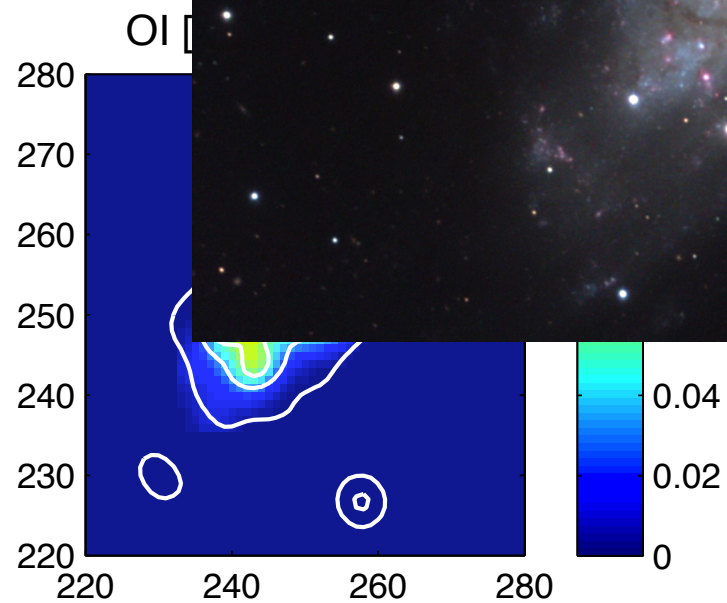
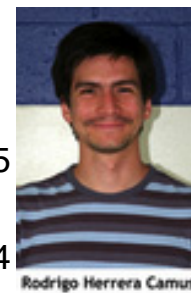
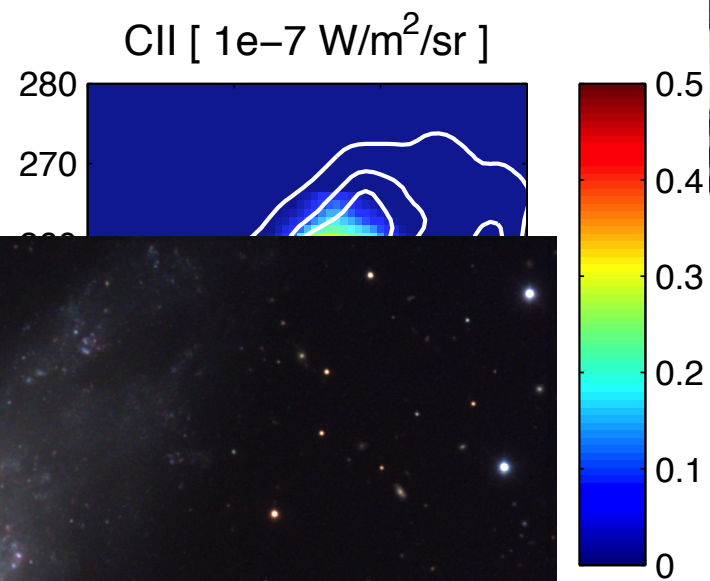
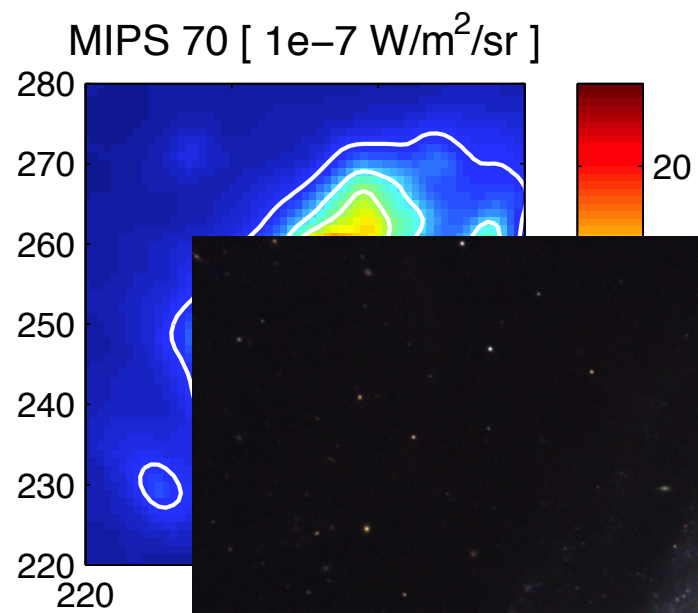
Power-law index of the Molecular Star Formation Law



Fraction of diffuse emission in SFR tracer



- We show how to robustly measure the index irrespective of SFR tracer, diffuse emission contamination, and fitting method.
- The molecular SFL is approximately linear in NGC4254



Opportunities for students

- CARMA STING (Nearby Galaxies)
 - Star formation law (PI)
- KINGFISH (Nearby Galaxies)
 - FIR spectroscopy and imaging
 - Submm spectroscopy just approved (col)
- HERITAGE (Magellanic Cloud science)
 - FIR spectroscopy just approved (PI)
- IRAM LP ($z \sim 1-2$ Medium Redshift)
 - EVLA program ongoing (PI)
 - $z \sim 3$ GRB massive host (PI)
- ALMA!
 - Early science call in 5 months
 - Partnership with PUC