

Eliminating Error in the Chemical Abundance Scale for Extragalactic H II Regions

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ABSTRACT

In an attempt to remove the systematic errors which have plagued the calibration of the H II region abundance sequence, we have theoretically modeled the extragalactic H II region sequence. We then used the theoretical spectra so generated in a double blind experiment to recover the chemical abundances using both the classical electron temperature + ionization correction factor technique, and the technique which depends on the use of strong emission lines (SELs) in the nebular spectrum to estimate the abundance of oxygen. We find a number of systematic trends, and we provide correction formulae which should remove systematic errors in the electron temperature + ionization correction factor technique. We also provide a critical evaluation of the various semi-empirical SEL techniques. Finally, we offer a scheme which should help to eliminate systematic errors in the SEL-derived chemical abundance scale for extragalactic H II regions.

Key words: galaxies: abundances, ISM: abundances, H II regions, methods: data analysis, techniques: spectroscopic

1 INTRODUCTION

To understand how galaxies in the early universe evolved into those that we see locally requires an understanding of both chemical and star formation history of galaxies over cosmic time. Indeed, observational data compiling hundreds of thousands galaxies at different redshifts are suggesting that there actually is a deep physical connection between metallicities, masses and star-formation rate (e.g. Tremonti et al. 2004; Brinchmann et al. 2004; Kewley & Ellison 2008; Lara-López et al. 2010a,b), although these results are mainly based on trends of the chemical abundances and not on their absolute metallicities.

Theoretical simulations can now predict the evolution of chemical abundances in galaxies (e.g. Mollá & Díaz 2005; Davé & Oppenheimer 2007). However, it is difficult to put these predictions to the test because the absolute calibration of the chemical abundance scale from H II regions is currently discrepant by more than a factor of two. See, for example the work of Kewley & Ellison (2008) where abundances were

calculated for 45,000 galaxies from the Sloan Digital Sky Survey (SDSS) (York et al. 2000; Abazajian et al. 2003) using 10 different calibrations given in the literature. Although all methods show some correlation with abundance, there is typically a scatter ranging over a factor of 2-3 in the absolute oxygen abundance derived from the various techniques. The same result was found by López-Sánchez & Esteban (2010), who also considered in their analysis the metallicities derived using a direct determination of the electron temperature of the ionized gas. Without a proper estimation of the absolute chemical abundances in galaxies, the comparison between observational data and theoretical models is not entirely satisfactory.

Thanks to its strong emission lines, the easiest element to measure in H II region emission spectra is oxygen. This is fortunate, since O is an α -process element made directly in short-lived massive stars. It comprises about 50% of all the heavy elements by mass throughout the universe, and is therefore representative of all the heavy elements. By contrast, Fe is made in lower mass stars, and is not released promptly into the interstellar medium of galaxies. The O/H abundance in H II regions –typically expressed in terms of

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$12+\log(\text{O}/\text{H})$, where O/H is the ratio of the number of oxygen to hydrogen atomic density. This has been determined in three semi-empirical ways, as well as through direct theoretical photoionization modeling:

1. By direct measurement of the electron temperature, from line ratios such as the $[\text{O III}] \lambda 4363/\lambda 5007$ (see Osterbrock & Ferland 2006), followed by an analysis of the various ionization fractions in the zones of the H II region which produce optical emission lines. We dub this the T_e **method**. This has been very extensively applied by observers over the past 30 years (e.g. Peimbert & Costero 1969; Stasińska 1978; Aller 1984; Díaz et al. 1987; Izotov, Thuan, & Lipovetski 1994; Esteban & Peimbert 1995; Vílchez & Esteban 1996; Stasińska 2005; Bresolin et al. 2005; Izotov et al. 2006; Esteban et al. 2004; López-Sánchez & Esteban 2009).

2. By direct measurement of recombination lines of ions of heavy elements such as O II and, in some cases, O I and comparing these with the recombination lines of hydrogen (e.g. Esteban et al. 1998, 2004, 2009; García-Rojas et al. 2004; García-Rojas & Esteban 2007; López-Sánchez et al. 2007; Peimbert et al. 2007). We refer to this technique as the **RL method**, although it is generally not useful in the determination of the cosmic abundance scale, since the recombination lines of heavy elements are generally too weak ($\sim 10^{4-6}$ times fainter than $\text{H}\beta$) to measure.

3. By measurement of the strong emission lines in the H II region spectrum calibrated using

- a. photoionization modeling (e.g. Pagel et al. 1979; McGaugh 1991; Kewley & Dopita 2002; Kobulnicky & Kewley 2004; Dopita et al. 2006a; Dors & Copetti 2006), or

- b. H II regions for which the oxygen abundance is known from the T_e method (e.g. Pilyugin 2001a,b; Denicoló, Terlevich & Terlevich 2002; Pettini & Pagel 2004; Pilyugin & Thuan 2005; Pérez-Montero & Díaz 2005; Pilyugin et al. 2010; Pilyugin & Mattsson 2011).

We will call this the **Strong Emission Line (SEL) method**, and it is nowadays extensively used to determine metallicities in large galaxy surveys.

Each of these methods has its particular problems, and range of application. More importantly, where all methods can be used, the derived O/H abundances often show systematic disagreements between the various methods. These may amount to factors of two or more. In this paper, we seek to resolve some of the causes of these discrepancies by re-deriving the chemical abundances of model H II regions using the T_e method and the standard SEL techniques. We use a double-blind approach, so that the analysis of the photoionization models was conducted without any *a priori* knowledge of the chemical abundances and physical parameters used in the photoionization models.

We will now briefly indicate how these alarming abundance discrepancies between the various methods may arise.

1.1 The T_e and RL methods

At normal nebular temperatures ($T_e \sim 10^4$ K), the auroral $[\text{O III}] \lambda 4363$ line is only a few percent as strong as the nebular $\lambda 5007$ line, so that high quality spectra are required to

give an adequate signal-to-noise ratio to measure the fainter line. When the S/N ratio is inadequate, the strength of the fainter line tends to be overestimated, leading to an underestimate in the derived abundance. This is also true for other commonly-used temperature sensitive line ratios.

A more insidious problem was first pointed out by Peimbert & Costero (1969). This still has not been resolved to everyone's satisfaction. Normal H II regions are not homogeneous but contain temperature gradients, dense inclusions in which collisional de-excitation of cooling lines leads to higher temperatures, as well as colliding supersonic flows in which both density and electron temperatures may be raised as a result of shocks. In all such regions of enhanced temperature, the $[\text{O III}] \lambda 4363/\lambda 5007$ emission line flux ratio is raised by a factor which depends only on the temperature. This would not be a problem except for the fact that the line emissivity is also raised in such regions by a factor proportional to the square of the electron density, n_e^2 . Thus, the temperature estimate provided by the forbidden line ratio is dominated by the line ratio characterizing these over-dense inclusions, rather than providing a measure of the electron temperature of the nebula as a whole. The electron temperature is therefore systematically overestimated. Hence, the presence of temperature gradients or fluctuations in the ionized gas leads to the abundances based on the T_e method being systematically underestimated (Peimbert 1967; Peimbert et al. 2007; Stasińska 2002, 2005).

Indeed, detailed studies that compared heavy elements abundances derived using both the T_e method and the RL method in Galactic (Esteban et al. 2004; García-Rojas et al. 2005, 2006; García-Rojas & Esteban 2007; Mesa-Delgado et al. 2009b; Mesa-Delgado & Esteban 2010; Mesa-Delgado et al. 2011) and extragalactic (Tsamis et al. 2003; Peimbert 2003; López-Sánchez et al. 2007; Esteban et al. 2009) H II regions find a very good agreement between both results when temperature fluctuations are considered. However, their origin and existence is still controversial because they are not well reproduced by standard photoionization models (Kingdon & Ferland 1995; Rodríguez & García-Rojas 2010), and hence additional mechanisms are proposed to explain the presence of temperature fluctuations in the ionized gas (see reviews by Esteban 2002; Peimbert & Peimbert 2006).

1.2 The SEL method

The strong emission line (SEL) method was first proposed by Pagel et al. (1979). It relies upon the ratio of the sum of the strong forbidden oxygen emission lines with respect to $\text{H}\beta$, the so-called R_{23} ratio; $R_{23} = I([\text{O II}] \lambda\lambda 3726, 3729 + [\text{O III}] \lambda\lambda 4959, 5007) / I(\text{H}\beta)$. Since then many other such ratios have been proposed. The most widely used of these are the calibrations of McGaugh (1991) and Kewley & Dopita (2002), based on detailed photoionization modelling, and the empirical relations provided by Pilyugin (2001a,b); Pilyugin & Thuan (2005) and Pilyugin, Vílchez & Thuan (2010). Both kinds of calibrations strive to improve the accuracy by making use of the $[\text{O III}]/[\text{O II}]$ ratio as ionization parameter, which accounts for the large scatter found in the R_{23} versus oxygen abundance calibration, which is larger than observational errors (Kobulnicky et al. 1999)

Hence, the SEL method is ideally suited to global abun-

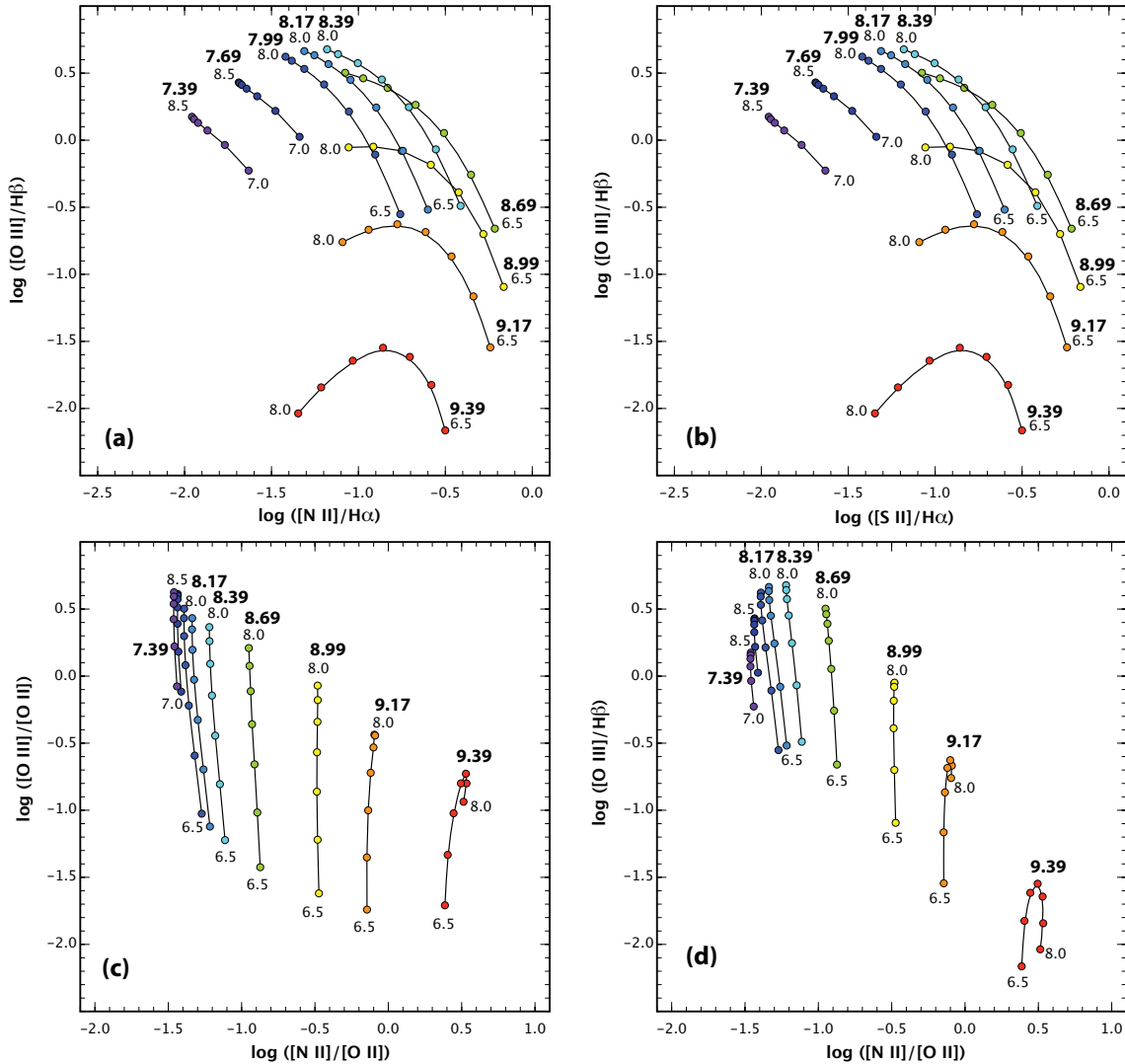


Figure 1. The complete model grid plotted on the Veilleux & Osterbrock (1987) diagnostics, $[\text{O III}]/\text{H}\beta$ vs. $[\text{N II}]/\text{H}\alpha$; (panel a) and $[\text{O III}]/\text{H}\beta$ vs. $[\text{S II}]/\text{H}\alpha$ (panel b). In the lower panels we show the Dopita et al. (2000) diagnostics $[\text{O III}]/[\text{O II}]$ vs. $[\text{N II}]/[\text{O II}]$ (panel c) and $[\text{O III}]/\text{H}\beta$ vs. $[\text{N II}]/[\text{O II}]$ (panel d). The models are identified by their input (gas + dust) $12 + \log(\text{O}/\text{H})$ (written in bold face and color coded) and their input ionization parameter, $\log q$. Models are stepped by 0.25 in $\log q$.

dance determinations in distant galaxies, since only the strong lines are visible against the stellar continuum. The abundance scale for ensembles of H II regions in galaxies has recently been calibrated by Dopita et al. (2006a). The basic problem with all ratios of optical forbidden lines to hydrogen recombination lines is that they are two-valued as a function of chemical abundance. This abundance ambiguity can only be raised by the simultaneous use of several such ratios. An additional problem is that not all SEL method give similar oxygen abundances: the SEL methods based on calibrations using photoionization models generally tend to overpredict (but not always, see Pérez-Montero et al. 2010; Dors et al. 2011) the observed oxygen abundances derived using the T_e method by 0.2–0.4 dex (Peimbert et al. 2007; Bresolin et al. 2009; López-Sánchez & Esteban 2010; Moustakas et al. 2010). As we should expect, that is not the case of the SEL techniques that are based on calibrations with objects for which the oxygen abundances are well known from the T_e method. Generally speaking,

the SEL method generally returns abundances which are higher than the T_e method, but in better agreement with those derived from recombination lines of heavy elements (López-Sánchez & Esteban 2010).

1.3 A double-blind test

In this paper, we seek to eliminate some of the systematic errors in abundance determinations by applying the T_e method and the SEL method to theoretically-generated H II region spectra, rather than using observed spectra. This has the advantage of eliminating observational errors but, more importantly, it also provides a direct test of whether the various methods can recover the chemical abundances and physical conditions which were fed into the theoretical models. In order to ensure objectivity, López-Sánchez (who applied the T_e and the SEL method) was provided a randomized selection of theoretical H II region models generated by Dopita. The data provided consisted of a set of line fluxes (relative

to $H\beta$) as delivered by the model. The models were identified only by letters, and no information was given about the input physical parameters which were input to generate the models. This is the classical double-blind technique often applied in medicine, but rarely in astrophysics.

In Section 2, we provide details of the techniques used to generate the theoretical spectra, in Section 3 we apply the T_e to the theoretical spectra and examine how well the input parameters have been recovered. Where possible, we provide correction formulae to apply. Section 4 presents the results provided by the most common SEL techniques. In this section, we also compare the oxygen abundances provided by the SEL methods with the derived T_e and RL abundances in high-quality Galactic and extragalactic H II spectra. Finally, Section 5 gives the conclusions and provides our recommended procedure needed to accurately derive the oxygen abundance in H II regions.

2 THE THEORETICAL MODELS

We have generated a sequence of photoionization models using the MAPPINGS III code, an updated version of the code originally described in Sutherland & Dopita (1993). These models are very similar to those produced by Dopita et al. (2000), and subsequently used by Kewley & Dopita (2002), but the chemical abundances have been specifically modified so as to provide a better fit to the Sloan Digital Sky Survey (SDSS) observed sequence of extragalactic H II regions (Kewley et al. 2006).

Our grid of photoionization models uses as stellar input the fluxes computed by the STARBURST 99 code (Leitherer et al. 1999) appropriate for continuous star formation. The input spectrum and the assumed stellar IMF input into the STARBURST 99 code is the same as used by Dopita et al. (2000). The later STARBURST 99 models (Vázquez & Leitherer 2005), which incorporate a fully self-consistent mass-loss formulation were not used as these generate a stellar radiation field that is somewhat too soft relative to these earlier models (see Dopita et al. 2006a, for a detailed discussion of the issue). To summarise, the most likely cause of this discrepancy is probably that the stellar wind is clumpy as a result of radiation pressure instabilities, so a multi-dimensional model atmosphere would need to be constructed. Effects of stellar rotation can also modify the emergent ionizing spectrum (Levesque et al. 2010). For consistency, we have interpolated the stellar fluxes to match the chemical abundances actually used in the model, allowing for the abundance scale shift caused by the re-calibration of the solar metallicity (Grevesse et al. 2010).

The extragalactic H II region sequence is clearly defined in the observational plane by the standard diagnostic line ratio plots of Baldwin, Phillips & Terlevich (1981) and Veilleux & Osterbrock (1987). These plot the ratios $[N II] \lambda 6584/H\alpha$, $[S II] \lambda \lambda 6717, 31/H\alpha$ and $[O I] \lambda 6300/H\alpha$ against the $[O II] \lambda 5007/H\beta$ emission line ratio. Kewley et al. (2006) gave these diagnostics for some 45,000 galaxies drawn from the SDSS. Our photoionisation models closely reproduce the observational sequence of extragalactic H II regions, bearing in mind that the ionization parameter tends to become higher as the abundance

decreases, thanks to the decreasing importance of the ram pressure of the stellar winds (Dopita et al. 2006a).

The solar metallicity for our photoionisation models was defined by the solar oxygen abundance given by Grevesse et al. (2010), $12+\log(O/H)=8.69\pm 0.05$. However, the abundance of helium and carbon were assumed to vary according the formula given by Dopita et al. (2006a). For nitrogen, we use a form similar to that given in Dopita et al. (2006b):

$$12 + \log(N/H) = 7.6 + \log[Z/Z_{\odot} + (Z/Z_{\odot})^2]. \quad (1)$$

Here the zero point of the nitrogen abundance has been adjusted to prove a better fit in the corresponding Veilleux & Osterbrock (1987) diagnostic diagram. Because nitrogen is an important coolant at high metallicity, this has the incidental effect of improving the fit of the models to the observed extragalactic H II region abundance sequence in the other two diagnostic diagrams. For all models we include dust, and dust physics according to the description in Dopita et al. (2006a), where the depletion factors of the various elements from the gaseous phase are also given.

Our models form a grid in chemical abundance or metallicity, defined by the quantity $12+\log(O/H)$, and the ionization parameter, $\log q$, the ratio of the number of photons passing per unit area and time to the number density of hydrogen atoms. The q parameter is related to the dimensionless ionization parameter U by $U = q/c$, where c is the speed of light.

Our photoionization models are spherical, with a target ionization parameter for the inner surface. This is unlike the Kewley & Dopita (2002) models, which were plane parallel. Our models are also isobaric with $\log(P/k) = 5.5 \text{ cm}^{-3} \text{ K}$, where k is the Boltzmann's Constant, so that the density varies throughout the model, but is typically $10 - 30 \text{ cm}^{-3}$. Unlike the Kewley & Dopita (2002) models, our models allow for photoelectric heating of the gas by dust grains, and include the effects of radiation pressure, which is becoming appreciable at the higher values of the ionization parameter.

2.1 The Double-Blind Test Models

From the theoretical model H II region grid we have taken, for a fixed metallicity, every second model in $\log q$, and applied to these the standard observational techniques for deriving abundances. In order that no bias was incurred in the analysis, the spectra were randomly selected from the input grid of models, and the raw spectra labelled simply A,B,C...Z, AA..AG were passed by Dopita to López-Sánchez for analysis. This analysis returned physical conditions, ionization parameters, and chemical abundances given by each of the semi-empirical techniques commonly used by observers, which included both the T_e and the SEL methods. We can then compare the results of this analysis to the known inputs into the original models to discover where there exist systematical errors, and to evaluate the magnitude of the random errors that can be generated by these techniques. The advantage of this approach is that it totally eliminates errors generated by the observations, and provides spectra of essentially infinite signal to noise for analysis.

In an ideal world, the parameters returned by the abundance analysis would be the same as those input into the

Table 1. The physical and chemical properties of the photoionization models used in the double-blind abundance fitting, organized according to chemical abundance $12+\log(\text{O}/\text{H})$, the nominal ionization parameter input into the model, the ionization parameter at the inner surface of the ionized gas, q_{in} , and mean ionization parameter in the ionized region of the model, $\langle q(\text{H}) \rangle$. For each model we give the emission line weighted mean temperatures (in K) for the [N II], [O III], [S II] and [O II] lines, and the mean recombination temperature for ionized hydrogen, $T(\text{rec})$.

Name	$12+\log(\text{O}/\text{H})$	$\log q$	q_{in}	$\langle q(\text{H}) \rangle$	$\langle T_{[\text{N II}]} \rangle$	$\langle T_{[\text{O III}]} \rangle$	$\langle T_{[\text{S II}]} \rangle$	$\langle T_{[\text{O II}]} \rangle$	$T(\text{rec})$
X	9.39	8.00	5.417E+07	2.598E+07	3363	2987	4277	4284	3439
J	9.39	7.50	2.510E+07	1.171E+07	3041	3620	4201	4194	3908
H	9.39	7.00	1.005E+07	4.426E+06	2881	4282	4352	4357	4307
C	9.39	6.50	3.393E+06	1.376E+06	2617	4568	4494	4558	4403
G	9.17	7.75	5.358E+07	2.432E+07	4478	4619	5544	5487	5004
AA	9.17	7.25	2.140E+07	9.721E+06	4111	5347	5633	5586	5511
O	9.17	6.75	7.393E+07	3.252E+06	3766	5765	5745	5782	5745
I	8.99	8.00	1.212E+08	4.573E+07	5988	5585	6780	6688	5928
R	8.99	7.50	4.210E+07	1.895E+07	6014	6162	6801	6688	6433
AC	8.99	7.00	1.494E+07	6.827E+06	5826	6736	6898	6872	6836
Z	8.99	6.50	4.883E+07	2.119E+06	5351	6936	6865	6935	6902
N	8.69	7.75	1.025E+08	3.850E+07	8604	8069	8782	8713	8294
A	8.69	7.25	3.265E+07	1.460E+07	8680	8448	8751	8729	8618
AG	8.69	6.75	1.076E+07	4.837E+06	8738	8764	8693	8791	8780
U	8.39	8.00	2.720E+08	6.599E+07	10490	10492	10173	10610	10520
F	8.39	7.50	7.609E+07	2.951E+07	10414	10492	10124	10517	10490
W	8.39	7.00	2.314E+07	1.030E+07	10290	10538	10000	10416	10420
AF	8.39	6.50	7.096E+06	3.055E+06	9960	10282	9663	10088	10040
T	8.17	7.75	1.720E+08	4.959E+07	11335	12146	10784	11676	11960
S	8.17	7.25	4.807E+07	1.997E+07	11207	11875	10784	11501	11630
L	8.17	6.75	1.410E+07	6.185E+06	10815	11379	10353	11060	11020
Q	7.99	8.00	3.853E+08	6.985E+07	11843	13584	11123	12474	13250
AE	7.99	7.50	1.005E+08	3.513E+07	11792	13180	11140	12366	12820
AB	7.99	7.00	2.787E+07	1.207E+07	11483	12420	10897	11912	11990
E	7.99	6.50	7.929E+06	3.320E+06	10793	11434	10284	11077	10980
B	7.69	8.25	7.843E+08	8.498E+07	12091	14699	11294	12997	14250
V	7.69	7.75	2.068E+08	5.365E+07	12094	14394	11320	12969	13930
M	7.69	7.25	5.561E+07	2.248E+07	11973	13558	11281	12676	13070
K	7.69	6.75	1.542E+07	6.641E+06	11428	12374	10783	11874	11840
Y	7.39	8.25	8.636E+08	8.379E+07	12220	15354	11398	13299	14830
AD	7.39	7.75	2.191E+08	5.455E+07	12227	14987	11421	13269	14450
P	7.39	7.25	5.769E+08	2.316E+07	12120	13985	11392	12952	13430
D	7.39	6.75	1.575E+08	6.754E+06	11559	12615	10866	12065	12030

photoionization models. However, in practice the different abundance techniques have different implicit assumptions about both the atomic physics relating to particular ionic species, and more generally, about the physics of the H II region itself. Therefore, we should not expect the results to agree exactly.

The physical and abundance parameters of the input models are given in Table 1. Note that we give three values of the ionization parameter, the nominal ionization parameter input into the model, q , the ionization parameter at the inner surface of the ionized gas, q_{in} , and mean ionization parameter in the model, $\langle q(\text{H}) \rangle$. These are related to each other, but not in any obvious way. For example, q is calculated on the basis of an assumption that the tem-

perature in the inner zone of the H II region is 10,000 K. Because of our isobaric assumption, q_{in} is lower than q for the high-abundance models, and higher than q_{in} for the low-abundance models. The ratio of q_{in} to $\langle q(\text{H}) \rangle$ is always greater than two due to the spherical divergence of the radiation field, and can reach values as high as ten where radiation pressure gradients start to play a role in determining the gas pressure.

As stated above, the input stellar spectrum, the physical conditions and the chemical abundance set were chosen so as to replicate, as far as possible, the observed sequence for extragalactic H II regions (Dopita et al. 2000; Kewley et al. 2006). In Figure 1 we show the two main diagnostic diagrams from Veilleux & Osterbrock (1987) and the diagrams which

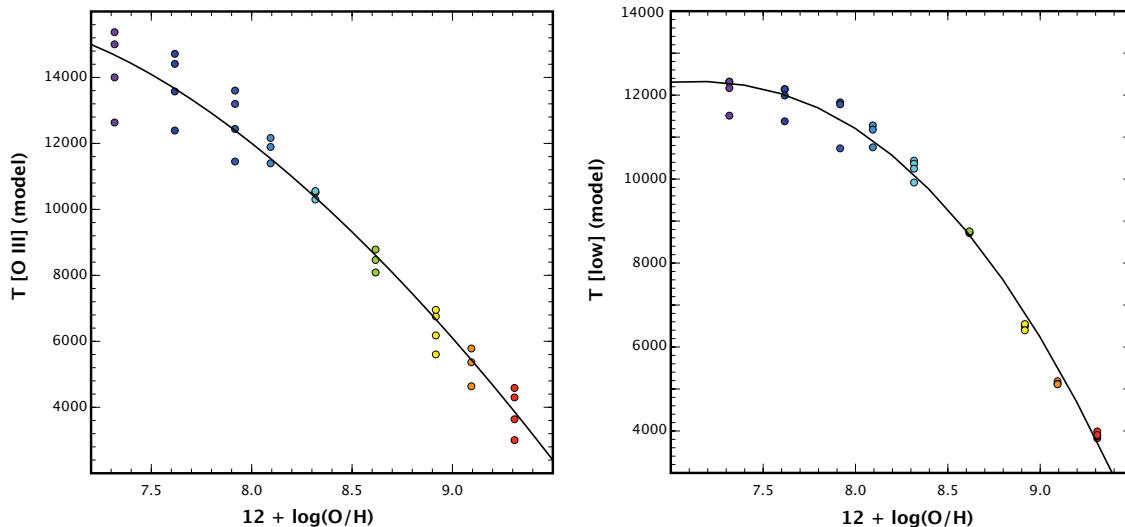


Figure 2. The relation between the oxygen abundance, $12 + \log(\text{O}/\text{H})$, and the electron temperature, T_e given by the models in the high-ionization zone containing $[\text{O III}]$ (left) and in the low-ionization zone containing the ions $[\text{S II}]$, $[\text{O II}]$, and $[\text{N II}]$ (right). The empirical fits we have derived for both high- and low-ionization zones as given by Eqs. 2 and 3 are shown as solid lines on these plots.

Dopita et al. (2000) used to separate the chemical abundance from the ionization parameter at the high-abundance end of the scale.

These models display quite a tight relationship between the oxygen abundance, $12 + \log(\text{O}/\text{H})$, and the electron temperature, T_e , given by the high-ionization zone from the $[\text{O III}]$ ($\lambda 4959 + \lambda 5007$)/ $\lambda 4363$ ratio, and in the low-ionization ($\text{H}^+ + \text{He}^0$) zone emitting the $[\text{O II}]$ ($\lambda 3726 + \lambda 3729$), $[\text{N II}]$ $\lambda 6584$ and $[\text{S II}]$ ($\lambda 6717 + \lambda 6731$) lines. This is shown in Fig. 2. Note that there is larger scatter in the plot for the $[\text{O III}]$ zone, because the temperature in this zone is more sensitive to the input ionization parameter. On the left hand plot, the left-hand upper points correspond to high ionization parameter, while the reverse is true for the points in the lower right-hand side of the diagram. There is much less variation in the local ionization parameter in the low-ionization zone of the model H II regions, so that the scatter is less for the right-hand panel of Fig. 2.

For each zone of the model H II region we have generated empirical fits to the oxygen abundance – electron temperature relationship. For $[\text{O III}]$ we find

$$T_e[\text{O III}] = 25200 - 1400x - 2000|x - 7.0|^{1.7}, \quad (2)$$

while for the low-ionization zone we have

$$T_e(\text{low}) = 9500 + 400x - 2000|x - 6.9|^{2.35}, \quad (3)$$

where $x = 12 + \log(\text{O}/\text{H})$ in both cases. These curves emphasize once again that the electron temperature is a key diagnostic for the chemical abundance in H II regions.

3 APPLICATION OF THE T_e METHOD

3.1 Physical conditions

We have taken care to analyze the model spectra in exactly the same way as an observer would do when presented with the spectrum of real H II regions. If there are no systematic effects in the analysis technique, this exercise should

yield the same electron temperatures and chemical abundances as was input into the models. One potential source of difference is the fact that the atomic data used in the nebular models is not necessarily the same as are used in the analysis of the spectra, although in many cases they are. The atomic data used for our models are detailed in Sutherland & Dopita (1993) (see their Section 3.5.2), which are similar to the atomic data included in IRAF task *temden* in the NEBULAR package by Shaw & Dufour (1995) (see their Table 1). The output of the various nebular modeling codes have been compared against each other for a number of standard models on more than one occasion through the so-called *Lexington Benchmarks* code-off exercise (Ferland 1995). Any egregious errors made by any one of them have subsequently been eliminated from the appropriate codes. We can therefore be fairly certain that other photoionization codes such as CLOUDY (e.g. Ferland et al. 1994) will produce very similar output to the MAPPINGS IIIs code used here.

The model spectra were created without considering any extinction or reddening, and not stellar absorption in the H I Balmer lines, to decrease the uncertainties when analyzing the data. Thus, for the analysis of the model spectra we have used a reddening coefficient, $c(\text{H}\beta)$, and a equivalent width of the stellar absorption underlying the H I Balmer lines, W_{abs} , equal to 0, following López-Sánchez & Esteban (2009).

We derived the electron temperature, T_e , and density, n_e , of the ionized gas of the models using several emission line ratios. The values obtained for each region are compiled in Table A1. All determinations were computed with the IRAF task *temden* in the NEBULAR package (Shaw & Dufour 1995). We used the updated atomic dataset for the O^+ , S^+ , and S^{++} ions as input to the NEBULAR routine. The references are indicated in Table 4 of García-Rojas et al. (2005).

In accord with Fig. 2, and with equations 2 and 3, we assumed a two-zone approximation for the H II region to describe the temperature structure of the nebulae. In the

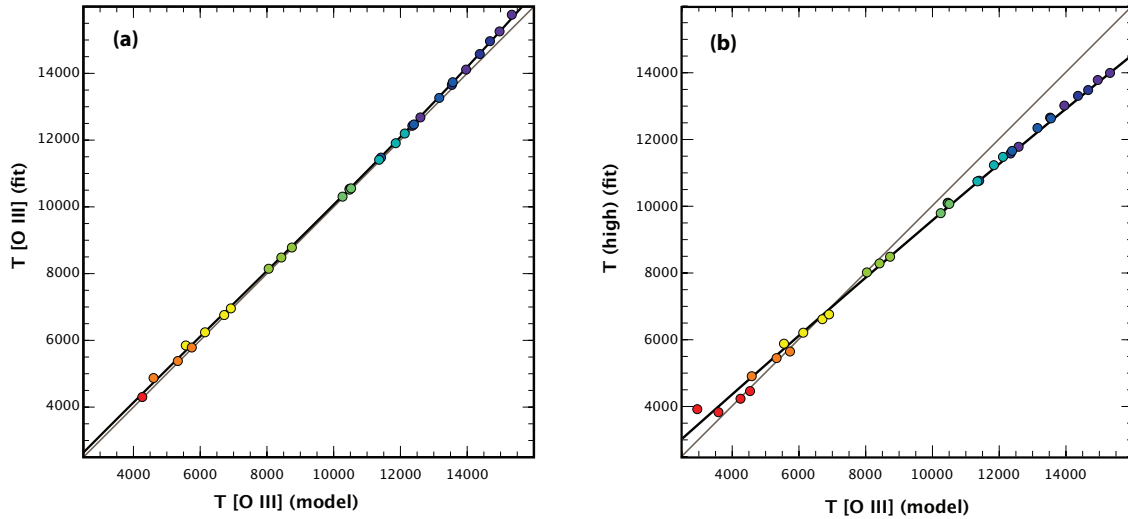


Figure 3. Comparison between the $T_e[\text{O III}]$ given by the model (x-axis) and the temperature derived by analysis of the model spectra. The left hand panel directly compares the $[\text{O III}]$ temperature from the model and the derived $[\text{O III}]$ temperature, while the right -hand panel compares the fitted temperature drawn from the mean of the $[\text{O III}]$, $[\text{S III}]$ and $[\text{Ar III}]$ temperatures with the $[\text{O III}]$ temperature from the model. The empirical fits we have derived are shown as a solid line for both panels. As in figure 2, the points are color coded according to the input abundance set.

high-ionization $\text{H}^+ + \text{He}^+$ zone, we used both the $[\text{O III}]$ temperature, $T_e[\text{O III}]$, and the mean of $T_e[\text{O III}]$, $T_e[\text{S III}]$ and $T_e[\text{Ar III}]$, $T_e(\text{high})$, as the representative temperature for high ionization potential ions. In the low-ionization zone containing $\text{H}^+ + \text{He}^0$ we adopted the mean of $T_e[\text{O II}]$, $T_e[\text{S II}]$ and $T_e[\text{N II}]$, $T_e(\text{low})$, for the low ionization potential ion temperatures.

The line ratios used to measure the temperatures are:

$$[\text{O III}] (\lambda 4959 + \lambda 5007) / \lambda 4363,$$

$$[\text{S III}] (\lambda 9069 + \lambda 9530) / \lambda 6311,$$

$$[\text{Ar III}] (\lambda 7135 + \lambda 7751) / \lambda 5192,$$

$$[\text{N II}] (\lambda 6548 + \lambda 6583) / \lambda 5755,$$

$$[\text{O II}] (\lambda 3726 + \lambda 3729) / (\lambda 7319 + \lambda 7330), \text{ and}$$

$$[\text{S II}] (\lambda 6717 + \lambda 6731) / (\lambda 4069 + \lambda 4076).$$

We note that some analyses of high-quality, deep spectrophotometric data of bright H II galaxies (e.g. Hägele et al. 2008, 2011; Pérez-Montero et al. 2010) consider the electron temperature derived for a particular ion only to obtain the ionic abundance of that particular ion. However, the two-zone approximation for the H II region is most widely used in extragalactic analyses.

Since, in the analysis of the spectra, no information about the pressures or density adopted in the models was given, we used the standard density-sensitive line ratios $[\text{O II}] \lambda 3726 / \lambda 3729$, $[\text{N I}] \lambda 5198 / \lambda 5200$, $[\text{Cl III}] \lambda 5517 / \lambda 5537$ and $[\text{S II}] \lambda 6717 / \lambda 6731$ to derive n_e . In practice, all models have a common pressure $P/k = 3 \times 10^5 \text{ cm}^{-3} \text{ K}$, all density-sensitive ion line ratios should be at their low-density limit. Table A1 compiles the value derived in all cases, and the average value adopted for each model. Once n_e was obtained we used it to derive T_e using the $[\text{O III}]$, $[\text{S III}]$, $[\text{Ar III}]$, $[\text{O II}]$, $[\text{S II}]$, and $[\text{N II}]$ line ratios, and we iterated until convergence was attained.

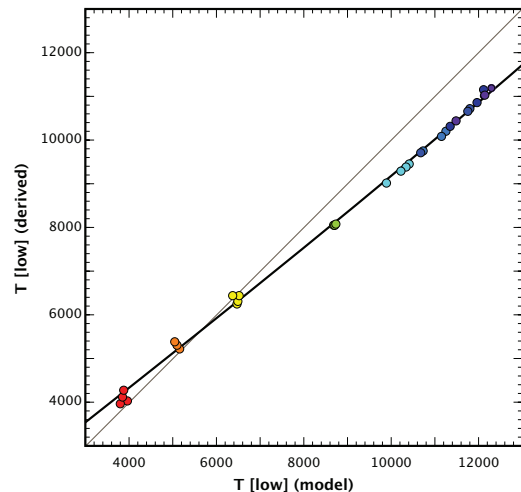


Figure 4. Comparison between the (emission weighted) T_e for the low-ionization species $[\text{N II}]$, $[\text{O II}]$ and $[\text{S II}]$, as returned by the model (x-axis) and the T_e derived for these same species from the temperature-sensitive line ratios (y-axis). The fit to the data (Eq. 6) is also plotted.

3.2 Results for T_e

The temperatures derived from the $[\text{O III}] (\lambda 4959 + \lambda 5007) / \lambda 4363$ ratio, and the mean temperature derived for the high-ionization species $[\text{O III}]$, $[\text{S III}]$ and $[\text{Ar III}]$, are compared with the $[\text{O III}]$ emission-line weighted temperature given by the model in Fig 3. The agreement for $[\text{O III}]$ is very good, while the mean of the high-ionization species tends to underestimate the true temperature.

We empirically derive the following third-order correction formula to transform the measured $[\text{O III}]$ line temper-

Table 2. Total abundances derived from the analysis of line intensity ratios of the model spectra using IRAF task *ionic* considering the electron temperature derived from all the high-ionization species, $T_e(\text{high})$, and the low ionization species, $T_e(\text{low})$, separately. All abundances are expressed in the form of $12 + \log(X/\text{H})$ or $\log(X/\text{O})$.

Model	O ⁺⁺ /O ⁺	O/H	N/H	N/O	S/H	S/O	Ne/H	Ne/O	Ar/H	Ar/O	Cl/H	Cl/O
A	-0.592	8.816	7.432	-1.384	7.147	-1.669	8.374	-0.443	6.883	-1.933	4.424	-4.393
B	0.326	7.777	6.129	-1.648	5.974	-1.803	7.109	-0.668	5.512	-2.264	3.426	-4.351
C ^a	-1.771	8.880	8.264	-0.616	7.601	-1.279	9.125	0.245	4.814	-4.066
D	-0.680	7.571	5.923	-1.649	5.843	-1.728	7.230	-0.342	5.618	-1.954	3.240	-4.332
E	-1.258	8.197	6.629	-1.567	6.466	-1.731	8.165	-0.032	6.632	-1.565	3.761	-4.436
F	-0.126	8.495	6.959	-1.536	6.810	-1.686	7.912	-0.583	6.340	-2.155	4.124	-4.371
G	-0.774	9.157	8.105	-1.051	7.595	-1.562	8.644	-0.512	7.301	-1.855	4.822	-4.335
H ^a	-0.889	8.763	8.205	-0.558	7.669	-1.094	8.600	-0.163	7.441	-1.323	4.884	-3.879
I	-0.654	9.050	7.796	-1.254	7.306	-1.744	8.495	-0.555	7.033	-2.017	4.568	-4.482
J ^a	-0.267	8.560	8.080	-0.480	7.723	-0.837	8.164	-0.396	7.011	-1.548	4.905	-3.654
K	-0.697	7.857	6.235	-1.623	6.142	-1.715	7.530	-0.327	5.931	-1.927	3.444	-4.413
L	-0.925	8.362	6.807	-1.555	6.627	-1.735	8.103	-0.259	6.573	-1.789	3.914	-4.448
M	-0.067	7.805	6.153	-1.653	6.094	-1.711	7.212	-0.594	5.597	-2.208	3.423	-4.383
N	-0.103	8.733	7.332	-1.401	7.116	-1.617	8.160	-0.573	6.630	-2.103	4.411	-4.323
O	-1.863	9.358	8.320	-1.038	7.611	-1.747	9.332	-0.026	7.996	-1.362	4.835	-4.523
P	-0.036	7.502	5.831	-1.670	5.786	-1.716	6.900	-0.602	5.283	-2.219	3.121	-4.381
Q	0.236	8.079	6.456	-1.622	6.308	-1.770	7.422	-0.656	5.829	-2.250	3.721	-4.358
R	-0.647	9.073	7.877	-1.196	7.439	-1.633	8.597	-0.476	7.174	-1.899	4.698	-4.374
S	-0.255	8.302	6.708	-1.594	6.598	-1.704	7.747	-0.555	6.163	-2.139	3.900	-4.402
T	0.112	8.256	6.657	-1.599	6.539	-1.717	7.621	-0.635	6.036	-2.221	3.896	-4.360
U	0.160	8.456	6.920	-1.536	6.760	-1.696	7.824	-0.632	6.252	-2.204	4.117	-4.338
V	0.230	7.780	6.129	-1.651	6.010	-1.770	7.125	-0.655	5.524	-2.256	3.423	-4.357
W	-0.677	8.564	7.055	-1.510	6.847	-1.717	8.167	-0.397	6.634	-1.930	4.132	-4.432
X ^a	-0.567	8.303	7.803	-0.501	7.492	-0.811	7.765	-0.538	6.544	-1.759	4.640	-3.663
Y	0.309	7.486	5.819	-1.667	5.664	-1.822	6.814	-0.672	5.211	-2.275	3.127	-4.359
Z	-2.018	9.183	8.028	-1.155	7.455	-1.728	9.371	0.188	7.939	-1.244	4.695	-4.488
AA	-1.189	9.320	8.267	-1.053	7.622	-1.698	8.962	-0.357	7.644	-1.676	4.854	-4.466
AB	-0.447	8.149	6.537	-1.612	6.429	-1.720	7.672	-0.477	6.077	-2.072	3.728	-4.422
AC	-1.232	9.163	7.988	-1.174	7.449	-1.714	8.923	-0.239	7.563	-1.600	4.699	-4.464
AD	0.243	7.481	5.813	-1.668	5.697	-1.784	6.822	-0.660	5.218	-2.264	3.123	-4.358
AE	0.048	8.092	6.467	-1.625	6.373	-1.719	7.466	-0.625	5.868	-2.224	3.719	-4.373
AF	-1.473	8.599	7.132	-1.467	6.870	-1.729	8.644	0.045	7.177	-1.423	4.151	-4.448
AG	-1.281	8.873	7.525	-1.348	7.161	-1.713	8.759	-0.114	7.347	-1.526	4.429	-4.444

^a For these models, there is not estimation of the high and low electron temperatures using the [O III] and [O II] ions because their auroral lines are not tabulated (see Table A1). Furthermore, *ionic* task does not allow electron temperatures lower than 5,000 K for the low ionization ions, so we adopted $T_e(\text{low}) = 5,000$ K in these four cases.

atures into the model [O III] line temperatures:

$$t_4^c = t_4(1 - 0.05(t_4 - 1)^2), \quad (4)$$

where $t_4 = T_e/10,000$ K is the temperature measured from the spectra, and t_4^c is the temperature corrected to that delivered by the corresponding model in units of 10^4 K. The corrections to $T_e[\text{O III}]$ are almost always smaller than the intrinsic uncertainty in the derived temperature – typically of order 500–1000 K.

For the mean of the high excitation species [O III], [S III] and [Ar III] the corresponding correction formula is:

$$t_4^c = t_4(1.12 + 0.02t_4^2) - 0.087, \quad (5)$$

with t_4 as the average value of the electron temperature derived using the [O III], [S III] and [Ar III] ratio.

For the low-ionization zones of the model H II regions we use the mean of the [O II], [S II] and [N II] temperatures given by the models, as was also done in the fitting exercise. The result of the comparison is shown in Fig 4. As we see, we find a clear linear correlation between model and fit, although the slope is not exactly unity. The empirical fit is given by the formula:

$$t_4^c = t_4(1.28 - 0.03t_4^2) - 0.15, \quad (6)$$

where t_4 and t_4^c are the temperature measured from the spectrum and the corrected to that delivered by the corresponding model in units of 10^4 K, respectively. The correlation

coefficients of the fits in Eqn. 4, 5 and 6 are all in excess of 0.98.

These differences between the temperatures delivered by the model are the result of the temperature structure within the ionized region. The temperatures delivered by the MAPPINGS III code are weighted by the emissivity of, for example, the [O III] $\lambda 5007$ line for [O III]. The [O III] $\lambda 4363$ line is more strongly affected by temperature stratification as pointed out by Peimbert & Costero (1969), so that the temperature delivered by measurement of the [O III] ($\lambda 4959 + \lambda 5007$)/ $\lambda 4363$ ratio is different to the strong-line weighted mean temperature delivered by the code.

For observers who are only able to measure the [O III] temperature in the high-ionization zone ($\text{H}^+ + \text{He}^+$), and need a correction formula to estimate the temperature in the low-ionization ($\text{H}^+ + \text{He}^0$) zone, the models provide a convenient relationship that fits to an accuracy of ± 300 K for temperatures above 10,000 K:

$$T_e(\text{low}) = T_e[\text{O III}] + 450 - 70 \exp[(T_e[\text{O III}]/5,000)^{1.22}]. \quad (7)$$

Note that the temperature sensitivity to the abundance is strong, and that in high-abundance H II regions, the low ionization species deliver a lower electron temperature. Note however, that the form and shape of Eq. 7 is rather different from the one proposed by Garnett (1992), $T_e(\text{low}) = 0.7 \times T_e[\text{O III}] + 3,000$. However, both the temperature offset and the slope are similar in the range

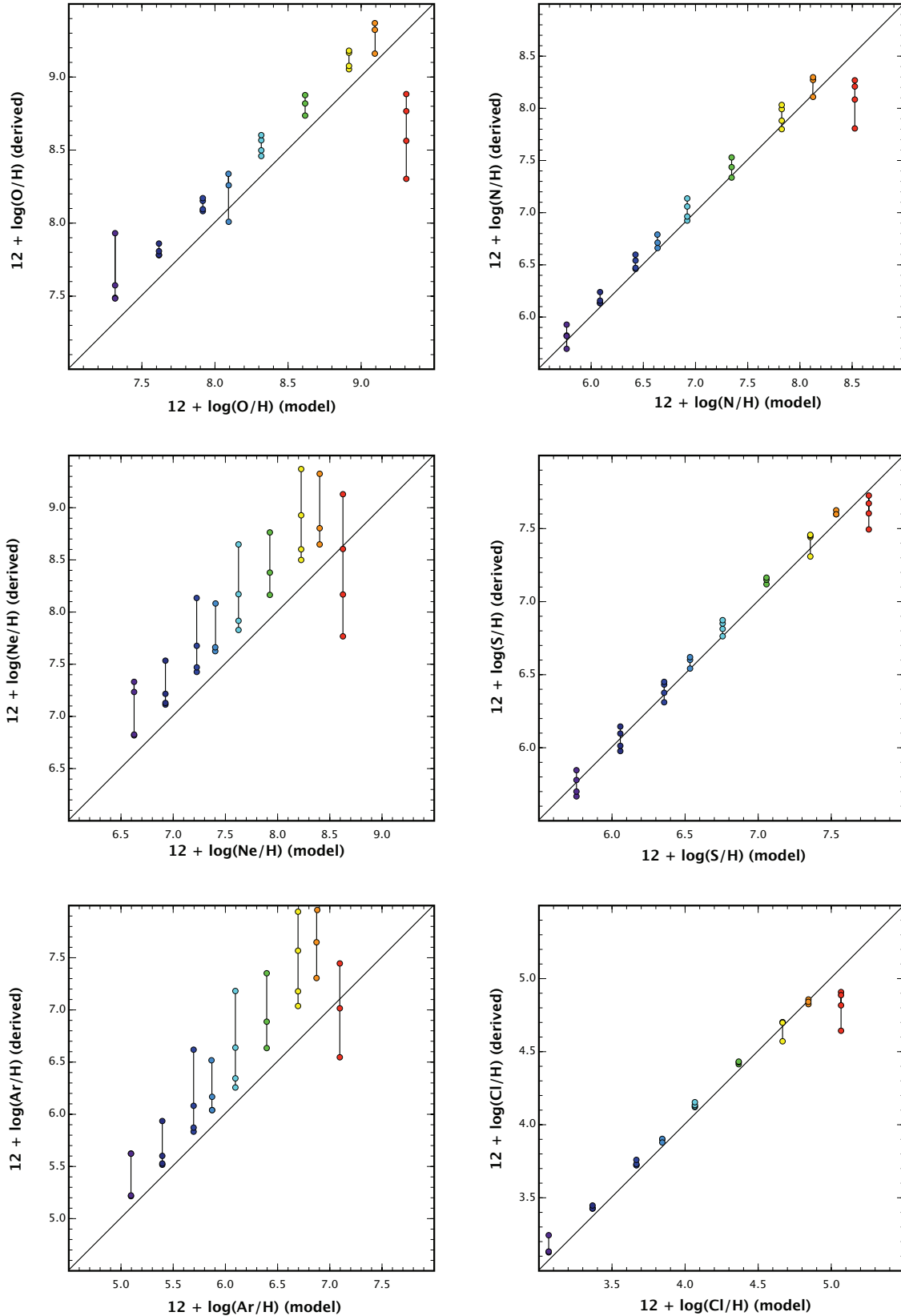


Figure 5. A comparison of the total abundances derived from the analysis of line intensity ratios of the model spectra and using T_e (high) and T_e (low) as representative for the electron temperature of the high- and low-ionization species, respectively (see Table 2), with the total abundances provided by the models. The fit is fairly good for the species for which the optical line emission arises principally in the low-ionization zone, N, S and Cl, but there is clearly a systematic error for the other species, O, Ne and Ar. For $12 + \log(\text{O}/\text{H}) = 9.39$, the T_e method fails because IRAF task *ionic* does not allow electron temperatures lower than 5,000 K for the low ionization ions, and hence $T_e(\text{low}) = 5,000$ K was adopted.

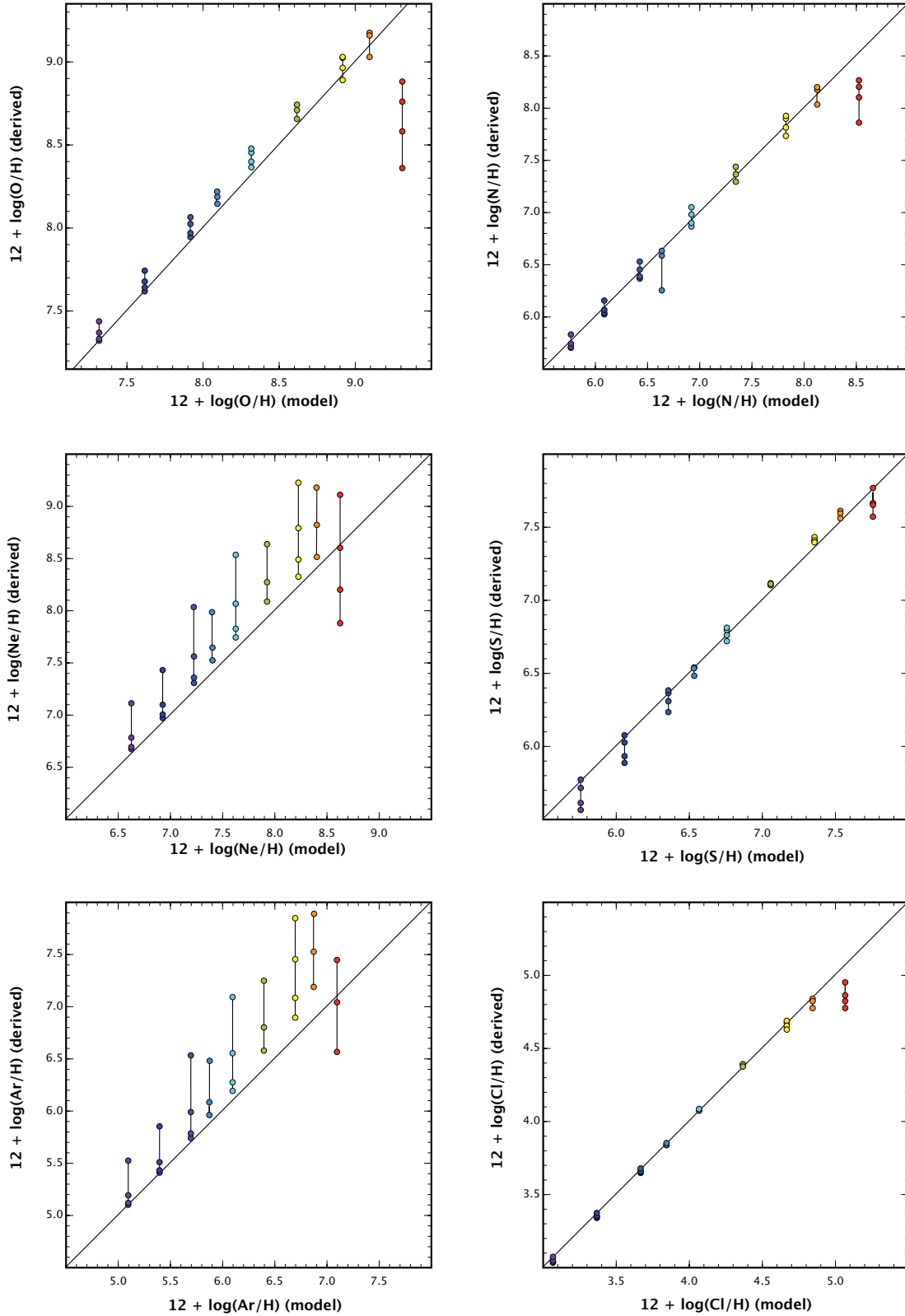


Figure 6. A comparison of the total abundances derived by using $T_e[\text{O III}]$ using the $[\text{O III}]$ ($\lambda 4959 + \lambda 5007$)/ $\lambda 4363$ ratio in the high-ionization zone and the $T_e[\text{O II}]$ using the $[\text{O II}]$ ($\lambda 3726 + \lambda 3729$)/($\lambda 7319 + \lambda 7330$) ratio in the low-ionization zone (see Table 3) with the total abundances provided by the models. As in Fig. 5, the fit is fairly good for the species for which the optical line emission arises principally in the low-ionization zone, N, S and Cl. Although it is slightly overpredicting the model abundances, the fit is now acceptable for O, except for $12 + \log(\text{O}/\text{H}) = 9.39$ because of the aforementioned problem when calculating $T_e[\text{O II}]$ for the low ionization ions. However, the fit still has a systematic error in the case of Ne (which is seen only in its high-ionization stage Ne^{++} and Ar (for which the emission is also dominated by its high ionization lines, and which is strongly affected by charge-exchange reactions). Random errors are also larger for these species.

Table 3. Total abundances derived from the analysis of line intensity ratios of the model spectra using IRAF task *ionic* considering only the electron temperature derived from the [O III] lines in the high-ionization zone, $T_e[\text{O III}]$, and that derived from the [O II] lines, $T_e[\text{O II}]$, in the low-ionization zone. All abundances are expressed in the form of $12 + \log(X/H)$ or $\log(X/O)$.

Model	O^{++}/O^+	O/H	N/H	N/O	S/H	S/O	Ne/H	Ne/O	Ar/H	Ar/O	Cl/H	Cl/O
A	-0.511	8.707	7.362	-1.344	7.114	-1.593	8.269	-0.438	6.797	-1.910	4.388	-4.319
B	0.344	7.616	6.021	-1.595	5.884	-1.732	6.968	-0.648	5.406	-2.210	3.337	-4.279
C ^a	-1.771	8.880	8.264	-0.616	7.601	-1.279	9.125	0.245	4.814	-4.066
D	-0.662	7.435	5.828	-1.607	5.770	-1.666	7.109	-0.327	5.521	-1.915	3.071	-4.365
E	-1.239	8.065	6.538	-1.527	6.397	-1.669	8.047	-0.018	6.538	-1.527	3.686	-4.380
F	-0.087	8.397	6.896	-1.501	6.759	-1.638	7.823	-0.574	6.271	-2.126	4.076	-4.321
G	-0.559	9.027	8.031	-0.996	7.609	-1.418	8.511	-0.515	7.185	-1.842	4.837	-4.190
H ^b	-0.934	8.758	8.200	-0.558	7.649	-1.109	8.598	-0.160	7.443	-1.315	4.860	-3.898
I	-0.140	8.888	7.730	-1.158	7.408	-1.480	8.322	-0.566	6.891	-1.997	4.671	-4.217
J ^a	-0.267	8.560	8.080	-0.480	7.723	-0.837	8.164	-0.396	7.011	-1.548	4.905	-3.654
K	-0.698	7.740	6.153	-1.587	6.073	-1.667	7.427	-0.313	5.849	-1.891	3.371	-4.369
L	-0.903	8.236	6.720	-1.516	6.562	-1.674	7.990	-0.247	6.482	-1.755	3.844	-4.392
M	-0.043	7.675	6.064	-1.610	6.024	-1.651	7.095	-0.579	5.506	-2.168	3.351	-4.323
N	-0.010	8.654	7.291	-1.363	7.100	-1.554	8.084	-0.570	6.575	-2.079	4.388	-4.266
O	-1.759	9.192	8.208	-0.984	7.549	-1.643	9.172	-0.020	7.887	-1.305	4.768	-4.424
P	-0.013	7.367	5.740	-1.627	5.713	-1.654	6.780	-0.586	5.189	-2.177	3.047	-4.320
Q	0.248	7.943	6.364	-1.579	6.232	-1.711	7.303	-0.640	5.737	-2.206	3.645	-4.298
R	-0.508	8.962	7.811	-1.151	7.429	-1.533	8.486	-0.476	7.080	-1.882	4.686	-4.276
S	-0.228	8.185	6.629	-1.556	6.538	-1.647	7.642	-0.543	6.080	-2.105	3.837	-4.348
T	0.139	8.143	6.583	-1.560	6.480	-1.663	7.520	-0.623	5.958	-2.185	3.835	-4.308
U	0.198	8.363	6.862	-1.501	6.717	-1.646	7.740	-0.623	6.189	-2.174	4.071	-4.292
V	0.243	7.639	6.033	-1.606	5.930	-1.709	7.001	-0.638	5.428	-2.210	3.344	-4.294
W	-0.645	8.451	6.977	-1.473	6.793	-1.658	8.063	-0.388	6.550	-1.901	4.074	-4.377
X ^a	-0.567	8.303	7.803	-0.501	7.492	-0.811	7.765	-0.538	6.544	-1.759	4.640	-3.663
Y	0.297	7.319	5.701	-1.619	5.562	-1.757	6.670	-0.650	5.098	-2.221	3.031	-4.289
Z	-1.929	9.027	7.922	-1.105	7.393	-1.634	9.222	0.195	7.844	-1.183	4.626	-4.400
AA	-1.054	9.173	8.172	-1.002	7.591	-1.583	8.818	-0.356	7.522	-1.651	4.821	-4.352
AB	-0.425	8.022	6.450	-1.572	6.362	-1.660	7.558	-0.464	5.986	-2.036	3.657	-4.365
AC	-1.134	9.022	7.895	-1.127	7.406	-1.616	8.787	-0.235	7.450	-1.571	4.652	-4.370
AD	0.251	7.330	5.709	-1.621	5.610	-1.720	6.690	-0.641	5.116	-2.215	3.037	-4.293
AE	0.071	7.967	6.383	-1.584	6.306	-1.661	7.355	-0.611	5.781	-2.185	3.650	-4.316
AF	-1.445	8.476	7.046	-1.429	6.809	-1.667	8.531	0.056	7.088	-1.388	4.083	-4.392
AG	-1.214	8.740	7.434	-1.305	7.109	-1.630	8.633	-0.106	7.244	-1.495	4.373	-4.367

^a For these models there is not estimation of the high and low electron temperatures using the [O III] and [O II] ions because their auroral lines are not tabulated (see Table A1). The results shown here are the same that in Table 2 because we used the $T_e(\text{high})$ and $T_e(\text{low})$. Furthermore, *ionic* task does not allow electron temperatures lower than 5,000 K for the low ionization ions, so we adopted $T_e(\text{low}) = 5,000$ K in these four cases.

$8,000 \leq T_e[\text{O III}] \leq 12,000$ where most of the high-quality nebula spectra have been obtained.

3.3 Ionic abundances

The IRAF package NEBULAR (Shaw & Dufour 1995) has been used to analyze the intensities of the forbidden lines using the temperatures and densities derived for the high- and low-ionization zones to derive ionic abundances of O^+ , O^{++} , N^+ , S^+ , S^{++} , Ne^{++} , Ar^{++} , Ar^{+3} , Cl^{++} , and Cl^{+3} . The ionic abundances so derived are listed in the Appendix in Table A2.

3.4 Total abundances

In all cases we adopted $\text{O}/\text{H} = \text{O}^+/\text{H}^+ + \text{O}^{++}/\text{H}^+$ to determine the total oxygen abundance. Although a very weak nebular He II $\lambda 4686$ line is produced in several models, the relative contribution of He^{++} to the total amount of helium is negligible, implying that O^{3+} has also a very low abundance in the nebula, thus we did not consider its contribution to the total O/H ratio. For the rest of the elements, we have to adopt a set of ionization correction factors (ICFs) to correct for the unseen ionization stages. The ICFs adopted here are basically the same used by López-Sánchez & Esteban (2009).

To derive the nitrogen abundance we assumed the standard ICF by Peimbert & Costero (1969): $\text{N}/\text{O} = \text{N}^+/\text{O}^+$, which is the typical assumption considered in the analysis of the ionized gas in extragalactic H II regions (e.g. López-Sánchez & Esteban 2010).

We measured two ionization stages of sulphur, S^+ and S^{++} , in all spectra. However, a significant contribution of S^{3+} is expected. We adopted the ICF given by Stasińska (1978), which is based on photoionization models of H II regions and is expressed as a function of the O^+/O ratio.

In the case of neon we also applied the classical ICF proposed by Peimbert & Costero (1969), that assumes that the ionization structure of Ne is similar to that of O. This is a good approximation for high ionization objects, where a small fraction of Ne^+ is expected.

For argon we have determinations of the Ar^{++} and Ar^{+3} abundances. However, some contribution of Ar^+ is expected. Hence, the total argon abundance was calculated by considering the ICFs proposed by Izotov et al. (1994).

We measured lines Cl^{++} and Cl^{+3} . As can be seen in Table A2, the dominant ionization stage is Cl^{++} and the contribution of Cl^+ to the total abundance is rather small. To take into account the Cl^+ fraction we have adopted the relation by Peimbert & Torres-Peimbert (1977).

The results we obtain depend on whether we use all of the high-ionization species to determine T_e in the high-ionization zone, $T_e(\text{high})$, and all the low-ionization

species to compute T_e in the low-ionization zone, $T_e(\text{low})$, or whether we use only the temperature derived from [O III] line ratio, $T_e[\text{O III}]$. The total abundances we obtain by these two approaches are shown in Tables 2 and 3, and are shown in graphical form in Fig 5 and 6, respectively.

When all of the high ionization species are used to derive $T_e(\text{high})$, the oxygen abundance is systematically overestimated by ~ 0.2 dex and the other species seen principally in their high-excitation stages, Ne and Ar, are overestimated by somewhat larger amounts. For $12+\log(\text{O}/\text{H})=9.39$, the T_e method fails because IRAF task *ionic* does not allow electron temperatures lower than 5,000 K for the low ionization ions. If we use only the temperature derived from the [O III] lines, $T_e[\text{O III}]$ using the [O III] ($\lambda 4959+\lambda 5007$)/ $\lambda 4363$ ratio in the high-ionization zone, and the temperature derived from the [O II] lines, $T_e[\text{O II}]$ using the [O II] ($\lambda 3726+\lambda 3729$)/($\lambda 7319+\lambda 7330$) ratio in the low-ionization zone, the systematic error is much less, specially for the total oxygen abundance, although it is still slightly overpredicted.

Although there are small systematic errors in this procedure, all are correctable. It is clear that the T_e method returns reliable abundances. Provided that the [O III] ($\lambda 4959+\lambda 5007$)/ $\lambda 4363$ ratio is used to determine the T_e in the high-ionization zone, the scatter in the derived abundances of O, N, S and Cl is typically ± 0.1 dex, provided that the O abundance is below $12+\log(\text{O}/\text{H}) < 9$. For high abundances, the T_e method returns systematically low abundances, where it can be applied. This is in agreement with the work of Stasińska (2002) and Stasińska (2005), who predicted that temperature gradients in these high-metallicity H II regions can cause the abundances to be underestimated by as much as ~ 0.4 dex.

Both for Ne and Ar, the systemic error and the scatter are distressingly large. We do not think that this is a consequence of using slightly different atomic data for models and line analysis using IRAF. However the cause is likely to be different in the two cases. In the case of Ne, the error is probably caused by the fact that only Ne^{++} is observable, and the [Ne III] $\lambda\lambda 3869, 3968$ doublet has sufficiently high excitation energy that the line emissivity is strongly biased towards the hottest part of the H II region. In the case of Ar, charge-exchange processes are very important in determining the ionic balance, and these strongly affect the ionization balance and therefore the ionization correction factors (Dopita et al. 1997). Furthermore, the assumptions of the ICFs used to derived the total abundances of Ne and Ar from their ionic abundances may be not valid in some cases (e.g. Pérez-Montero et al. 2007).

Given that Fig 2 shows that the electron temperature and the abundance are closely correlated, it is interesting to see just how well the oxygen abundance can be recovered from the measured [O III] temperature alone. To do this we used our empirical relation (Eq. 4) to correct the measured temperature and then applied the relationship between $T_e[\text{O III}]$ and $12+\log(\text{O}/\text{H})$ given by Eq. 2 to obtain the oxygen abundance. The result is shown in Fig. 7. For abundances $12+\log(\text{O}/\text{H}) > 8.0$ the abundances are recovered to an accuracy of ± 0.1 dex. At lower abundances, significant scatter is introduced by variations in the ionization parameter q , and it becomes essential to determine this parameter independently.

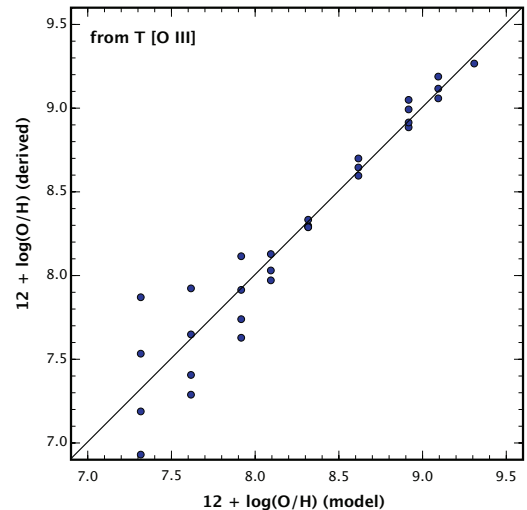


Figure 7. A comparison of the abundances derived by using only the $T_e[\text{O III}]$ measured from the [O III] ($\lambda 4959+\lambda 5007$)/ $\lambda 4363$ ratio. For abundances $12+\log(\text{O}/\text{H}) > 8.0$ the abundances are recovered to an accuracy of ± 0.1 dex. At lower abundances, the scatter is much increased because of the increased sensitivity of $T_e[\text{O III}]$ with the ionization parameter in this region.

4 APPLICATION OF THE SEL METHOD

The strong emission line (SEL) technique to derive nebular abundances is typically used when the spectra are of too low signal-to-noise ratio for the temperature sensitive emission lines to be detectable. This typically encompasses the majority of global determinations of metallicity in external galaxies, and all galaxies with redshifts large enough to be of cosmological interest. An extensive review of 10 metallicity calibrations, including theoretical and empirical methods, has been presented by Kewley & Ellison (2008), using data gathered from the SDSS. López-Sánchez & Esteban (2010) also review the most common empirical calibrations and compare their results with those derived using the T_e method. The majority of the empirical calibrations rely on ratios between bright emission lines to estimate the oxygen abundance. The most commonly used of these are known by their shorthand contractions; R_{23} , S_{23} , P , N_2 , O_3N_2 , N_2O_2 and y . Their definitions are:

$$R_3 = \frac{I([\text{O III}])\lambda 4959 + I([\text{O III}])\lambda 5007}{\text{H}\beta}, \quad (8)$$

$$R_2 = \frac{I([\text{O II}])\lambda\lambda 3726, 3729}{\text{H}\beta}, \quad (9)$$

$$R_{23} = R_3 + R_2, \quad (10)$$

$$P = \frac{R_3}{R_{23}}, \quad (11)$$

$$y = \log \frac{R_3}{R_2} = \log \frac{1}{P^{-1} - 1}, \quad (12)$$

$$N_2 = \frac{I([\text{N II}])\lambda 6584}{\text{H}\alpha}, \quad (13)$$

Table 4. Results for the oxygen abundance, in the form $12+\log(\text{O}/\text{H})$, according to the most-commonly used empirical SEL calibrations.

Model	T_e	BRANCH	M91	KD02o	KK04	KN2O2	P01	PT05	PVT10	PM11	D02	PP04a	PP04b	PP04c	S23	S23O23
A	8.71	HIGH	8.53	8.84	8.67	8.64	8.21	8.11	8.44	8.49	8.56	8.46	8.41	8.40	8.96	8.60
B	7.62	LOW	7.50	7.67	7.81	...	7.39	7.36	7.17	7.12	7.62	7.73	7.74	7.94	7.45	7.74
C	8.88	HIGH	9.13	9.31	9.35	9.24	8.86	8.43	9.11	8.94	8.72	8.59	8.58	9.25	8.53	9.65
D	7.44	LOW	7.23	7.57	7.53	...	7.34	7.01	6.73	6.31	7.71	7.80	7.84	8.28	7.28	8.31
E	8.01	LOW	7.96	7.85	8.10	...	8.77	6.94	7.18	6.74	8.21	8.19	8.18	8.51	8.17	8.50
F1 ^a	8.40	LOW	8.17	8.28	8.36	...	8.02	8.13	7.93	7.95	8.26	8.23	8.20	8.17	8.74	8.37
Fu ^a	8.40	HIGH	8.17	8.73	8.55	8.34	8.27	8.22	7.93	7.95	8.26	8.23	8.20	8.17	8.74	8.37
G	9.03	HIGH	9.07	9.22	9.16	9.09	8.84	8.70	8.73	8.76	8.54	8.45	8.39	8.69	8.46	9.30
H	8.76	HIGH	9.14	9.33	9.35	9.26	8.88	8.57	8.94	8.88	8.66	8.54	8.51	9.05	8.39	9.67
I	8.89	HIGH	8.95	9.02	9.06	8.92	8.76	8.69	8.65	8.67	8.44	8.37	8.31	8.45	8.55	8.94
J	8.56	HIGH	9.22	9.37	9.47	9.28	8.91	8.69	9.01	9.00	8.49	8.41	8.35	8.98	8.04	9.74
K	7.74	LOW	7.54	7.81	7.79	...	7.78	7.15	7.15	6.74	7.94	7.98	8.03	8.31	7.67	8.33
L	8.24	LOW	8.12	7.97	8.27	...	8.57	7.52	7.67	7.28	8.32	8.28	8.24	8.41	8.42	8.42
M	7.68	LOW	7.50	7.58	7.80	...	7.40	7.41	7.18	7.07	7.78	7.85	7.91	8.07	7.62	8.07
N	8.65	HIGH	8.61	8.77	8.74	8.63	8.41	8.36	8.52	8.55	8.38	8.32	8.27	8.26	8.74	8.51
O	9.19	HIGH	9.00	9.07	9.14	9.06	8.65	8.29	8.76	8.67	8.81	8.65	8.70	8.97	9.02	9.29
P	7.37	LOW	7.16	7.53	7.52	...	7.00	7.20	6.75	6.62	7.54	7.67	7.64	8.05	7.24	8.04
Q	7.94	LOW	7.91	7.98	8.15	...	7.73	7.74	7.53	7.52	7.86	7.92	7.98	7.98	7.90	7.86
R	8.96	HIGH	8.88	8.97	9.02	8.92	8.61	8.51	8.56	8.57	8.61	8.51	8.46	8.54	8.86	8.96
S	8.19	LOW	8.07	8.13	8.27	...	8.00	7.98	7.77	7.64	8.17	8.16	8.16	8.17	8.34	8.22
T	8.14	LOW	8.05	8.08	8.26	...	7.85	7.92	7.67	7.68	8.02	8.04	8.08	8.05	8.20	8.04
U1 ^a	8.36	LOW	8.13	8.22	8.32	...	7.90	8.02	7.88	7.89	8.13	8.12	8.13	8.08	8.40	8.13
Uu ^a	8.36	HIGH	8.13	8.58	8.54	8.33	8.34	8.29	7.88	7.89	8.13	8.12	8.13	8.08	8.40	8.13
V	7.64	LOW	7.50	7.65	7.81	...	7.38	7.36	7.16	7.11	7.65	7.75	7.78	7.97	7.51	7.83
W1 ^a	8.45	LOW	8.27	8.29	8.41	...	8.45	8.04	7.99	8.00	8.43	8.36	8.31	8.35	8.72	8.42
Wu ^a	8.45	HIGH	8.27	8.78	8.59	8.38	8.11	7.99	7.99	8.00	8.43	8.36	8.31	8.35	8.72	8.42
X	8.30	HIGH	9.32	9.40	...	9.28	8.92	8.62	9.12	9.18	8.36	8.31	8.26	9.05	7.71	9.76
Y	7.32	LOW	7.15	7.47	7.52	...	7.01	7.13	6.76	6.68	7.39	7.55	7.38	7.92	7.12	7.70
Z	9.03	HIGH	8.76	9.01	8.98	8.92	8.36	7.98	8.73	8.61	8.82	8.66	8.72	8.95	9.29	9.08
AA	9.17	HIGH	9.03	9.17	9.14	9.07	8.73	8.52	8.66	8.65	8.71	8.58	8.57	8.77	8.81	9.28
AB	8.02	LOW	7.90	7.84	8.12	...	7.96	7.69	7.61	7.36	8.11	8.11	8.13	8.22	8.12	8.25
AC	9.02	HIGH	8.79	9.02	8.97	8.92	8.42	8.18	8.53	8.53	8.77	8.63	8.65	8.70	9.12	9.00
AD	7.33	LOW	7.14	7.46	7.52	...	6.99	7.14	6.75	6.66	7.42	7.57	7.43	7.94	7.15	7.79
AE	7.97	LOW	7.90	7.89	8.14	...	7.75	7.76	7.55	7.51	7.94	7.98	8.03	8.05	8.02	8.02
AF1 ^a	8.48	LOW	8.38	7.85	8.44	...	9.43	7.29	7.80	7.93	8.51	8.42	8.37	8.62	8.84	8.61
AFu ^a	8.48	HIGH	8.38	8.80	8.72	8.47	8.07	7.75	7.80	7.93	8.51	8.42	8.37	8.62	8.84	8.61
AG	8.74	HIGH	8.45	8.76	8.65	8.67	8.03	7.74	8.48	8.50	8.69	8.56	8.54	8.62	9.15	8.71

^a For these models, for which we derived $12+\log(\text{O}/\text{H})\sim 8.4$ following the T_e method, we list the results of the empirical calibrations considering both the low and the high metallicity branches.

NOTE: The empirical calibrations and the parameters used for each of them are:

M91: McGaugh (1991) using R_{23} and y ;

KD02: Kewley & Dopita (2002) using R_{23} and the ionization parameter defined in that paper, q_{KD02o} ;

KK04: Kobulnicky & Kewley (2004) using R_{23} and the ionization parameter defined in that paper, q_{KK04} ;

KN2O2: Kewley & Dopita (2002) using the N_2O_2 parameter (calibration only valid for objects in the high metallicity branch);

P01: Pilyugin (2001a,b) using R_{23} and P ;

PT05: Pilyugin & Thuan (2005) using R_{23} and P ;

PVT10: Pilyugin et al. (2010) using R_{23} , P , N_2 and S_2 ;

PM11: Pilyugin & Mattsson (2011) using R_3 , N_2 and S_2 ;

D02: Denicoló, Terlevich & Terlevich (2002) using the N_2 parameter;

PP04a: Pettini & Pagel (2004), using a linear fit to the N_2 parameter;

PP04b: Pettini & Pagel (2004), using a cubic fit to the N_2 parameter;

PP04c: Pettini & Pagel (2004), using the O_3N_2 parameter;

S23: Pérez-Montero & Díaz (2005) using S_{23} ;

S23O23: Pérez-Montero & Díaz (2005) using S_{23} and R_{23} .

$$O_3N_2 = \frac{I([\text{O III}])\lambda 5007}{I([\text{N II}])\lambda 6584}, \quad (14)$$

$$N_2O_2 = \frac{I([\text{N II}])\lambda 6584}{I([\text{O II}])\lambda\lambda 3726, 3729}, \quad (15)$$

$$S_{23} = \frac{I([\text{S II}])\lambda\lambda 6717, 6731 + I([\text{S III}])\lambda\lambda 9076, 9532}{H\beta}, \quad (16)$$

$$S_{23}O_{23} = \frac{S_{23}}{R_{23}}. \quad (17)$$

Simple line ratios such as R_{23} (Jensen et al. 1976; Pagel et al. 1979), S_{23} (Vílchez & Esteban 1996; Christensen et al. 1997; Díaz & Pérez-Montero 2000) –or indeed any forbidden line to recombination line ratio– suffer

from being two-valued, and hence a separate calibration for the low and high metallicity regimes is usually needed. In the case of the R_{23} index, the calibrations are given for $12+\log(\text{O}/\text{H})\lesssim 8.1$ (low metallicity) and $12+\log(\text{O}/\text{H})\gtrsim 8.4$ (high metallicity). That means that a very large fraction of the star-forming regions lie in the ill-defined turning zone around $12+\log(\text{O}/\text{H})\sim 8.20\text{--}8.30$, where regions with the same R_{23} value have oxygen abundances that differ by almost an order of magnitude (see, for example, Figure A.1 in López-Sánchez & Esteban 2010). The reason of this behavior is that the intensity of the oxygen (or any other heavy element) lines do not monotonically increase with metallicity. At low abundance, the lines of the heavy element being considered are weak relative to a hydrogen recombination line due to their low abundance. However, at high abundance they are weak due to the low electron

temperature, which suppresses collisional excitations into the excited state responsible for the forbidden line. The turn-around in the line ratio can be extended to higher abundance by choosing lines with lower excitation potentials which make the ratio less temperature-sensitive. Hence the utility of S_{23} over R_{23} . The use of multiple line ratios also helps to remove the abundance ambiguities.

Ratios such as O_3N_2 (Alloin et al. 1979; Pettini & Pagel 2004) were introduced in an attempt to sidestep the abundance ambiguity entirely. This ratio is much more monotonic in abundance because of the strong secondary component of N enrichment alluded to above; see Eq. 1. This helps at high abundance, but the ratio once again becomes ambiguous when $12 + \log(O/H) \lesssim 8.0$. A ratio such as N_2O_2 (Dopita et al. 2000; Kewley & Dopita 2002) also provides a useful calibration for metallicity at high abundance. Its advantage is that both N^+ and O^+ co-exist in the same zone of the nebula, the $[O\text{ III}]$ lines become weaker at high metallicity due to the low electron temperature, while the $[N\text{ II}]$ lines become stronger due to the higher relative abundance of this element. However, its disadvantage is that the nitrogen and oxygen emission lines are widely separated in wavelength, making flux calibration and reddening corrections more critical.

The fundamental weakness in the use of all these strong-line ratios is that the overall spectrum of an H II region depends not only upon the chemical abundances, but also upon the ionization parameter q or U . That this is so is abundantly obvious from Fig. 1. The importance of this was recognized by Baldwin et al. (1981), who were amongst the first to use line ratio diagnostic diagrams. However, the first empirical calibration involving an ionization parameter was presented by McGaugh (1991), who developed models of H II regions using the photoionization code CLOUDY. McGaugh (1991) introduced the y parameter to derive (together with the R_{23} ratio) the oxygen abundance using only the bright oxygen lines. The analytical expressions for these models were given by Kobulnicky et al. (1999).

It is clear from Fig. 1c, and as emphasized by Dopita et al. (2000) and Kewley & Dopita (2002), that the $[O\text{ III}]/[O\text{ II}]$ ratio is probably the best one to use to determine the ionization parameter, particularly if the $[N\text{ II}]/[O\text{ II}]$ ratio is also measured. To a lesser extent R_3 can also be used. Kewley & Dopita (2002) provided a procedure to allow a simultaneous solution of both $\log q$ and $12 + \log(O/H)$, which involves the y and R_{23} parameters, although they also give a parametrization of the oxygen abundance and $\log q$ using other parameters such as $[S\text{ III}]/[S\text{ II}]$ (which is also a good estimator of the ionization parameter), N_2O_2 , S_{23} and O_3N_2 . Later, Kobulnicky & Kewley (2004) provided a parametrization of the Kewley & Dopita (2002) method using the y and R_{23} parameters with a form similar to that given by McGaugh (1991) calibration. Following this method, $\log q$ and the oxygen abundance are computed iteratively using only the bright oxygen lines.

Pilyugin and his co workers have attempted an empirical calibration of the R_{23} and P parameters using the bright $[O\text{ III}]$ $\lambda\lambda 4959, 5007$ and $[O\text{ II}]$ $\lambda\lambda 3726, 3729$ nebular lines. To get this calibration, they used a combination of photoionization models and empirical data from observed H II regions where the auroral $[O\text{ III}]$ $\lambda 4363$ line was available. Indeed, the so-called Pilyugin parameter, P , was in-

Table 5. Results of the comparison between the oxygen abundance given by several SEL methods and the oxygen abundance assumed by the theoretical models. We indicate the parameters, a and b , of the linear fit $x_{SEL} = a + bx_{model}$, the correlation coefficient, r , of this linear fit, the dispersion of the data, σ , and the average value of the difference between the abundance given by the SEL method and that assumed by the model. In the first row, we also include the comparison with the oxygen abundance derived following the T_e method. The description of the SEL methods is the same than that used in Table 4.

Method	a	b	r	σ	offset
Te	-0.02 ± 0.14	1.001 ± 0.017	0.9959	0.05	-0.01
M91	-0.20 ± 0.17	1.005 ± 0.020	0.9940	0.07	-0.15
KD02	0.16 ± 0.21	0.981 ± 0.025	0.9897	0.10	0.00
KN2O2	0.94 ± 0.18	0.887 ± 0.020	0.9960	0.05	-0.07
KK04	0.77 ± 0.10	0.916 ± 0.012	0.9975	0.07	0.07
KD	0.61 ± 0.16	0.922 ± 0.019	0.9932	0.09	-0.05
P01	1.30 ± 0.64	0.814 ± 0.075	0.8890	0.31	-0.25
PT05	1.58 ± 0.51	0.752 ± 0.061	0.9113	0.28	-0.50
PVT10	-1.58 ± 0.24	1.134 ± 0.029	0.9901	0.14	-0.44
PM11	-2.46 ± 0.44	1.229 ± 0.052	0.9736	0.25	-0.52
D02	2.11 ± 0.45	0.736 ± 0.055	0.9390	0.20	-0.04
PP04a	3.43 ± 0.35	0.575 ± 0.043	0.9392	0.25	-0.04
PP04b	3.39 ± 0.43	0.578 ± 0.052	0.9141	0.26	-0.05
PP04c	1.81 ± 1.04	0.767 ± 0.116	0.8548	0.19	-0.25
S23	-1.47 ± 0.66	1.186 ± 0.081	0.9481	0.24	0.04
S23O23	-1.76 ± 0.67	1.208 ± 0.076	0.9641	0.17	0.07

roduced by Pilyugin (2000), after confirming that the R_{23} parameter has a systematic error depending on the hardness of the ionizing radiation. Hence, the excitation parameter P also considers the ionization degree of the H II region. His first empirical calibrations involving both the R_{23} and the P parameter were presented in Pilyugin (2001a) and Pilyugin (2001b) for high- and low-metallicity H II regions. Subsequently, Pilyugin & Thuan (2005) and Pilyugin, Vílchez & Thuan (2010) tried to improved such calibrations including more spectroscopic measurements of H II regions in spiral and irregular galaxies derived using the T_e -method. Finally, Pilyugin & Mattsson (2011) provided another empirical calibration considering only the R_3 , N_2 and S_2 parameters (the so-called NS calibration), which was derived to estimate oxygen abundances in galaxies whose spectrum is lacking of the $[O\text{ II}]$ $\lambda\lambda 3726, 3729$ emission lines.

4.1 Comparison of the SEL Techniques

We have applied all of these commonly-used SEL techniques to our model spectra to derive oxygen abundances, again *via* a double-blind procedure. The different SEL techniques fall into various classes, depending on which combination of strong line ratios are used. Table A3 lists the values of all these parameters derived for each model. This table also includes the value derived for the q parameter obtained from the optimal calibration provided by Kewley & Dopita (2002) and Kobulnicky & Kewley (2004). Table 4 compiles the oxygen abundances derived for each model, indicating which branch (high or low metallicity) the model spectra would fall into when using the R_{23} parameter. We considered the values of the oxygen abundances derived from the T_e method to choose the expressions of the lower or the upper branch provided by the empirical calibrations. For models F, U, W and AF, for which we obtained oxygen abundances

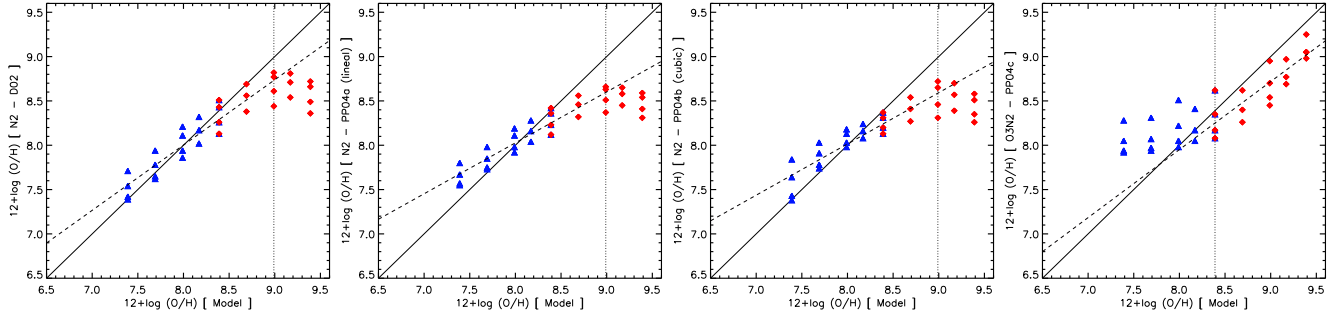


Figure 8. Comparison between the total abundances used for the models (x-axis) and those computed using the empirical calibrations considering the N_2 and the O_3N_2 parameters from Denicoló et al. (2002) and Pettini & Pagel (2004). The continuous line represents $y = x$. Red diamonds and blue triangles represent high- and low-metallicity models, respectively. The vertical dotted line indicates the lower (if using the N_2 parameter) or upper (when using the O_3N_2 parameter) limit of validity of the fits. The dashed line indicates the fit to the data within the appropriate metallicity range.

of $12+\log(\text{O}/\text{H}) \sim 8.4$ using the T_e method, the results provided by both the low- and high-metallicity branches are tabulated, and both sets of data are plotted in Figs. 8 to 11. Note that the calibrations that invoke the N_2 , O_3N_2 , S_{23} or $S_{23}O_{23}$ parameters give the same result for the low and high metallicity branches. The Kewley & Dopita (2002) calibration using the N_2O_2 parameter is only valid for objects in the high-metallicity branch.

To quantify the goodness of these SEL techniques, we have performed a linear fit between the oxygen abundances derived from these empirical methods and the oxygen abundances assumed by the models. In all cases, for models with $12+\log(\text{O}/\text{H})=8.39$ we have assumed the average value between the high and low metallicity branches. We note that not all models have been considered in this fit, but only those within the validity range of each SEL calibration. Table 5 compiles the results for these linear fits, including the correlation coefficient, the dispersion of the data and the average value of the difference between the abundance given by the SEL method and that assumed by the models (offset). First row of Table 5 lists the results of a linear fit between the oxygen abundances derived following the T_e method and those imposed by the models. This fit does not consider the models with $12+\log(\text{O}/\text{H})=9.39$ (see Fig. 6).

4.1.1 N_2 and O_3N_2 methods

Let us first consider the techniques which rely upon the N_2 parameter; that of Denicoló, Terlevich & Terlevich (2002) and those drawn from the Pettini & Pagel (2004) which involve a linear interpolation of the N_2 parameter, a cubic fit to this same parameter, and their fit using the O_3N_2 parameter. These are shown in Fig. 8. All these calibrations have similar scatter $\sim 0.20 - 0.25$ dex. However, they are well-behaved (if rather curved) in their abundance sequence. This type of curvature was also observed by Kewley & Ellison (2008), if we can interpret the mass sequence of galaxies as equivalent to an abundance sequence.

Yin et al (2007) indicated that the N_2 and the O_3N_2 indices are only useful for calibrating metallicities of galaxies with $12+\log(\text{O}/\text{H}) < 8.5$, while Pérez-Montero & Contini (2009) established that empirical calibrations using the O_3N_2 parameter are not valid for objects with $12+\log(\text{O}/\text{H}) \lesssim 8.0$. Following Fig. 8, we suggest that the

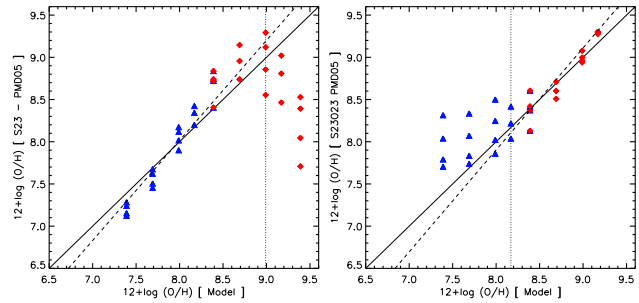


Figure 9. Comparison between the total abundances used for the models (x-axis) and those computed using the empirical calibrations considering the S_{23} and the $S_{23}O_{23}$ parameters from Pérez-Montero & Díaz (2005). The continuous line represents $y = x$. Red diamonds and blue triangles represent high- and low-metallicity models, respectively. The vertical dotted line indicates the upper (using the S_{23} parameter) or lower (using the $S_{23}O_{23}$ parameter) limit of validity of the fits. The dashed line indicates the fit to the data within the appropriate metallicity range.

N_2 method should not be applied for $12+\log(\text{O}/\text{H}) \gtrsim 8.7-9.0$, while the O_3N_2 calibration provided by Pettini & Pagel (2004) is only valid for $12+\log(\text{O}/\text{H}) \gtrsim 8.7$. Importantly, in all cases the result has a high uncertainty, ~ 0.25 dex, which is a consequence of lacking of any parameter which considers the ionization degree of the gas. Furthermore, in the case of galaxies showing an overabundance of nitrogen (e.g. Pustilnik et al. 2004; López-Sánchez et al. 2007; López-Sánchez & Esteban 2010; Monreal-Ibero et al. 2010), the application of any N_2 or O_3N_2 calibration will provide misleading oxygen abundances.

4.1.2 S_{23} and $S_{23}O_{23}$ methods

Figure 9 compares the oxygen abundances assumed by our models with the results provided using the Pérez-Montero & Díaz (2005) calibrations, that consider the S_{23} and $S_{23}O_{23}$ parameters. In the case of using the S_{23} index, the calibration seems to be valid up to $12+\log(\text{O}/\text{H}) \sim 9.0$, having a dispersion of 0.24 dex. However, this method cannot be applied for $12+\log(\text{O}/\text{H}) \gtrsim 9.0$. On the other hand, the $S_{23}O_{23}$ calibration works quite well at high metallicities, $12+\log(\text{O}/\text{H}) \gtrsim 8.5$. But it is not valid for $12+\log(\text{O}/\text{H}) \lesssim 8.2$ as the result basically does not de-

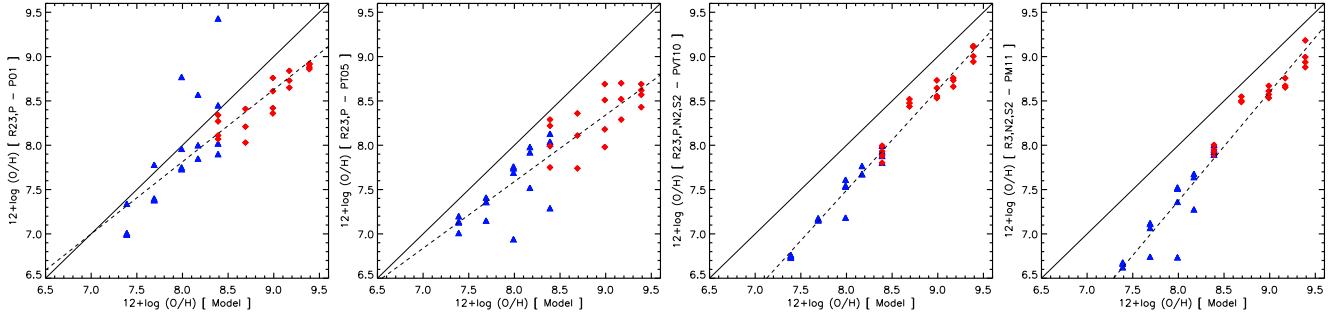


Figure 10. Comparison between the abundances used for the models (x-axis) and those computed via the Pilyugin method using the R_{23} and P parameters. From left to right, panels show the results for the empirical calibrations provided by Pilyugin (2001a,b), Pilyugin & Thuan (2005), Pilyugin et al. (2010) and Pilyugin & Mattsson (2011). The continuous line represents $y = x$. Red diamonds and blue triangles represent high- and low-metallicity models, respectively. The dashed line indicates the fit to the data. Note that, in the left panel, the three data points with SEL abundances very much larger to those of the models are in the intermediate-metallicity regime and also have a relatively large $\log q$. Hence, following the Pilyugin (2001a,b) method the high-metallicity calibration should be considered and averaged to the value derived using the low-metallicity calibration to get a more appropriate result of the oxygen abundance.

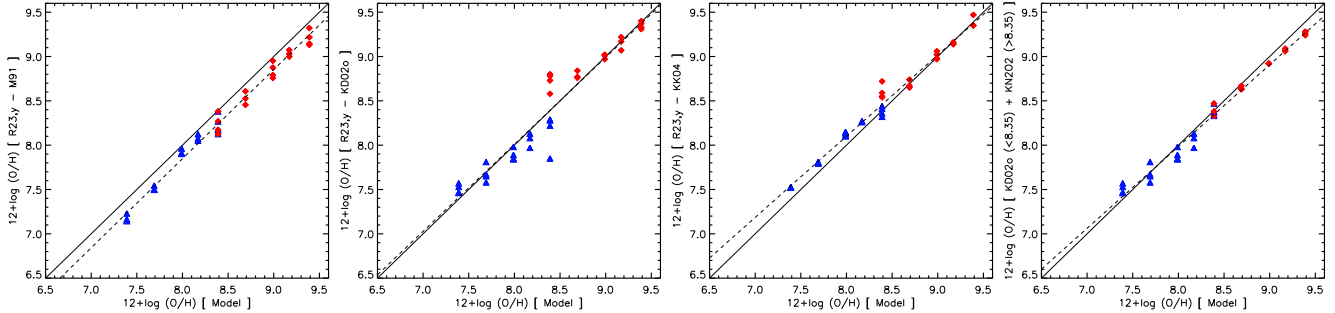


Figure 11. Comparison between the total abundances used for the models (x-axis) and those computed using the McGaugh (1991) method (left panel) and the Kewley & Dopita (2002) and Kobulnicky & Kewley (2004) techniques, all of them based on photoionization models. The continuous line represents $y = x$. Red diamonds and blue triangles represent high- and low-metallicity models, respectively. The dashed line indicates the fit to the data. The far right panel shows the comparison with the oxygen abundance considering the N_2O_2 method for the high-metallicity branch, $12+\log(\text{O}/\text{H}) \geq 8.35$, and the R_{23} and q_{KD020} for the low-metallicity branch, following Kewley & Dopita (2002).

pend on metallicity, having a constant value which depends on q , and a high dispersion, >0.6 dex.

Much of the scatter in the Denicoló, Terlevich & Terlevich (2002), Pettini & Pagel (2004) and Pérez-Montero & Díaz (2005) calibrations is due to a failure to include the effects of the ionization parameter on the line ratios. That is specially important in the analysis of 2D spectroscopic data, as the oxygen abundance maps derived using these calibrations may directly reflect the ionization structure of the H II region and not actual changes in metallicity (López-Sánchez et al. 2011). It is clear that the scatter could be reduced to ~ 0.1 dex if the ionization parameter is taken into account, since the vertical separation of the points on the figures is almost entirely the result of the different ionization parameters used in the models. However, this would not solve the curvature issue.

4.1.3 Pilyugin method

We next consider the Pilyugin et al. empirical calibrations of R_{23} and P (Pilyugin 2001a,b; Pilyugin & Thuan 2005; Pilyugin et al. 2010; Pilyugin & Mattsson 2011). As we discussed before, P is an empirical parameter that traces

the ionization degree of the H II region. The results we obtain are shown in Fig. 10. The main characteristic of these SEL calibrations is that there is a systematic offset in the sense that the Pilyugin method tends to underestimate the abundance by up to ~ 0.25 dex (Pilyugin 2001a,b), ~ 0.4 dex (Pilyugin et al. 2010) and ~ 0.5 dex (Pilyugin & Thuan 2005; Pilyugin & Mattsson 2011). It is also clear that the scatter for the Pilyugin (2001a,b) and Pilyugin & Thuan (2005) empirical calibrations is typically 0.3 dex. In both calibrations, the scatter becomes considerably larger in the $12+\log(\text{O}/\text{H}) \sim 8.0-8.4$ range, as a result of the ambiguity in determining the appropriate branch on which the R_{23} is located. We note that for this intermediate-metallicity regime it is common to give an average value between the oxygen abundances derived for the high and low metallicity branches if using the Pilyugin (2001a,b) and Pilyugin & Thuan (2005) method (e.g. Moustakas et al. 2010; López-Sánchez & Esteban 2010).

In the case of the Pilyugin et al. (2010) and Pilyugin & Mattsson (2011) methods (that also considers the N_2 and S_2 indices), three calibrations are provided for low, intermediate and high metallicity. In these cases, the scatter is slightly lower, ~ 0.14 dex and ~ 0.25 dex, respectively, than the scatter found using the previous calibrations.

The offset is also large, 0.4-0.5 dex, when comparing with that provided by the models.

The very interesting point here is that Pilyugin method is giving the best results when it is applied for real objects for which the oxygen abundance has been computed using the T_e method (e.g. Bresolin et al. 2009; López-Sánchez & Esteban 2010; Moustakas et al. 2010). However, this method is clearly failing when predicting the metallicities of the theoretical models using the bright emission lines. We will further discuss this issue in Sect. 4.2.

4.1.4 Kewley and Dopita method

Although Pilyugin method considers a parameter which is related to ionization degree of the gas, the scatter provided by this calibration is still relatively high. However, the scatter is much reduced when using a SEL method which is based on photoionization models, as those provided by the McGaugh (1991), Kewley & Dopita (2002) and Kobulnicky & Kewley (2004) methods. All these models explicitly consider an ionization parameter, y . Fig. 11 shows their results. As we see, in all these cases the agreement is excellent. We remind the reader that we are plotting both the high and low metallicity results for the models with $12+\log(\text{O}/\text{H})=8.39$.

Left panel of Fig. 11 shows the case of the McGaugh (1991) calibration using the Kobulnicky et al. (1999) parametrization. This calibration systematically underestimates the abundances given by the models by ~ 0.15 dex, but their dispersion is only 0.07 dex. In the case of using the Kewley & Dopita (2002) or Kobulnicky & Kewley (2004) methods considering the R_{23} and y parameter, the results are much better. Their method allows to iteratively determine the ionization parameter and the oxygen abundance. Both, the techniques by Kewley & Dopita and by Kobulnicky & Kewley, eliminate the scatter in the SEL technique to a very large extent. For the Kewley & Dopita (2002) technique it is ~ 0.1 dex, while for the Kobulnicky & Kewley (2004) technique it is ~ 0.07 dex excepting the ambiguous region around $12+\log(\text{O}/\text{H})=8.4$.

Far right panel of Fig. 11 shows a fit using the $N_2\text{O}_2$ calibration provided by Kewley & Dopita (2002) for the high-metallicity branch, $12+\log(\text{O}/\text{H})\geq 8.35$, and relied on the R_{23} and the technique they use for obtaining the ionization parameter for the lower metallicities. This assumption provides very good results, as this method has a very low offset, 0.05 dex, and dispersion, 0.09 dex.

The slight offset in the abundances derived following the Kewley & Dopita (2002) and Kobulnicky & Kewley (2004) methods is probably due to the fact that they used a different grid of models. The models we use in this paper are more closely related to the earlier Kewley & Dopita (2002) grid, except in respect of the much more sophisticated treatment of dust physics employed here.

However, the recent analysis of the ionized gas within a sample of ~ 40 strong star-forming galaxies performed by López-Sánchez & Esteban (2010) shows that the oxygen abundances provided by photoionization models (McGaugh 1991; Kewley & Dopita 2002; Kobulnicky & Kewley 2004) are, systematically, 0.2–0.3 dex higher than the oxygen abundances provided by the T_e method. This observational result has been also found by Yin et al (2007), Bresolin et al.

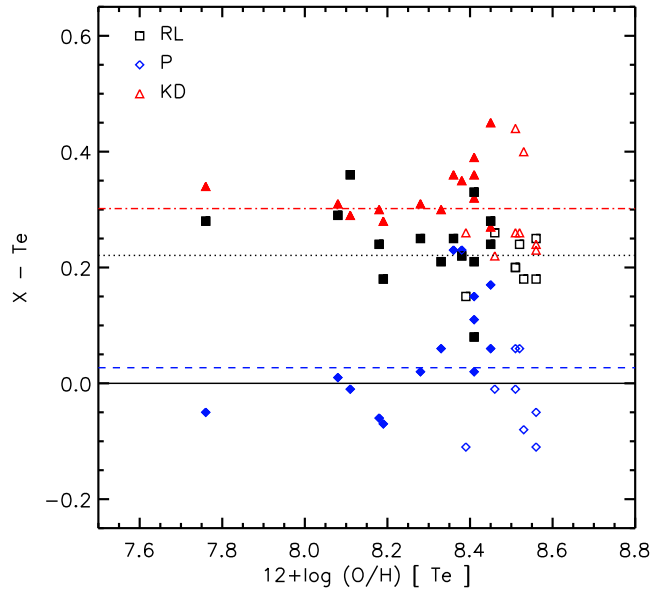


Figure 12. Comparison between the oxygen abundance derived following the T_e , RL, P and KD methods for deep, high-resolution observations of Galactic (open symbols) and extragalactic (filled symbols) H II regions for which oxygen recombination lines have been detected. The x -axis plots the T_e abundance, where the y -axis represents the difference between RL (black squares), P (blue diamonds) or KD (red triangles) abundances and the T_e abundance. The black dotted line plots the average difference between the T_e and the RL methods. The blue dashed line represents the average difference between the T_e and the P methods. The red dotted-dashed line represents the average difference between the T_e and the KD methods.

(2009), Moustakas et al. (2010) and López-Sánchez et al. (2011); in some objects the differences reach up to 0.6 dex. Interesting, however, the photoionization models developed by Dors et al. (2011) provided O/H values close to those obtained using the T_e method. Detailed tailor-made photoionization models of individual galaxies (e.g. Pérez-Montero et al. 2010) also provided results similar to those derived from the observed electron temperatures of the ionized gas.

4.2 Comparison of the SEL techniques with real data

Hence, we have further investigated the validity of the SEL techniques using real data. For this, we aim to use the deepest, highest quality observational data of Galactic and extragalactic H II regions nowadays available. Particularly, we also want to compare the results provided by the T_e and the most common SEL techniques (the Pilyugin and the Kewley & Dopita methods) with the oxygen abundances derived using oxygen recombination lines (RL). As it is well known (e.g. Peimbert 2003; Esteban et al. 2002, 2004, 2009; García-Rojas et al. 2004, 2005; García-Rojas & Esteban 2007; Peimbert et al. 2007), in all cases the oxygen abundances determined using RL are 0.2-0.3 dex higher than the oxygen abundances derived using the T_e method, which is based on collisionally excited lines (CEL). The effect does not depend on metallicity, electron density, electron and Balmer temperatures or ionization

Table 6. Results of the oxygen abundances given by the Te, RL, P and KD methods for Galactic (G) and extragalactic (E) H II regions for which oxygen recombination lines have been detected.

Object	R_2	R_3	$\log N_2$	T_e	RL	P	KD	Type	Ref
M 42	2.981	0.375	-1.247	8.51±0.03	8.71±0.03	8.57	8.77	G	E04
NGC 3576	4.169	2.447	-0.767	8.56±0.03	8.74±0.06	8.51	8.80	G	GR04
S 311	1.318	5.591	-1.720	8.39±0.05	8.54±0.10	8.28	8.65	G	GR05
M 16	0.668	3.547	-2.055	8.56±0.07	8.81±0.07	8.45	8.79	G	GR06
M 20	0.350	0.009	-0.549	8.53±0.06	8.71±0.07	8.45	8.93	G	GR06
NGC 3603	0.828	0.389	-1.932	8.46±0.05	8.72±0.05	8.45	8.68	G	GR06
M 8	3.038	5.002	-1.184	8.51±0.05	8.71±0.04	8.50	8.95	G	GR07
M 17	3.038	5.002	-1.184	8.52±0.04	8.76±0.04	8.58	8.78	G	GR07
30 Dor	1.823	1.109	-0.692	8.33±0.02	8.54±0.05	8.39	8.63	E	Pe03
N 11 B	1.925	5.728	-1.502	8.41±0.00	8.74±0.00	8.43	8.80	E	T03
N 66	2.052	6.349	-1.360	8.11±0.00	8.47±0.00	8.10	8.40	E	T03
NGC 6822V	3.000	3.764	-1.296	8.08±0.03	8.37±0.09	8.09	8.39	E	PPR05
NGC 5253A	2.429	3.860	-1.020	8.18±0.04	8.42±0.13	8.12	8.48	E	LS07
NGC 5253B	1.538	0.091	-0.430	8.19±0.04	8.37±0.10	8.12	8.47	E	LS07
NGC 5253C	0.554	1.229	-2.160	8.28±0.04	8.53±0.09	8.30	8.59	E	LS07
NGC 595	0.596	0.287	-0.792	8.45±0.03	8.69±0.05	8.51	8.72	E	E09
NGC 604	0.255	0.032	-0.627	8.38±0.02	8.60±0.03	8.61	8.73	E	E09
H 1013	1.038	1.178	-0.938	8.45±0.06	8.73±0.09	8.62	8.90	E	E09
NGC 5461	0.121	0.030	-0.865	8.41±0.03	8.49±0.06	8.56	8.73	E	E09
VS 44	1.603	0.690	-1.618	8.36±0.02	8.61±0.04	8.59	8.72	E	E09
NGC 2363	4.134	1.111	-1.096	7.76±0.02	8.04±0.05	7.71	8.10	E	E09
K 932	1.085	2.203	-1.842	8.41±0.02	8.62±0.03	8.52	8.77	E	E09

REFERENCES: E04: Esteban et al. (2004), E09: Esteban et al. (2009), GR04: García-Rojas et al. (2004), GR05: García-Rojas et al. (2005), GR06: García-Rojas et al. (2006), GR07: García-Rojas & Esteban (2007), LS07: López-Sánchez et al. (2007), Pe03: Peimbert (2003), PPR05: Peimbert, Peimbert & Ruiz (2005), T03: Tsamis et al. (2003).

degree (García-Rojas & Esteban 2007; Mesa-Delgado et al. 2009a) and it may be attributed to the presence of temperature fluctuations in H II regions.

Moreover, recent analyses (Peimbert, Peimbert & Ruiz 2005; Przybilla, Nieva & Butler 2008; Simón-Díaz & Stasińska 2011) have found that the oxygen nebular abundance based on RL agrees better with the abundances of the stars associated with the nebulae than the one derived from the T_e method, suggesting that RL abundances are more reliable than the abundances derived from CEL.

Table 6 lists all the 22 Galactic and extragalactic H II regions found in the literature for which oxygen abundances are available using both the RL and the T_e methods. Indeed, the average offset between T_e and RL abundances is 0.22 dex. Table 6 also compiles the derived R_{23} , P , y and N_2 indexes derived from the reddening-corrected data. We must emphasize that, in almost all the cases, the error in the [O II] $\lambda\lambda 3726, 3729$ lines is less than 5%. Using these observational parameters, we then applied the Pilyugin et al. method and the Kewley & Dopita method to determine the oxygen abundances using these SEL techniques. Table 6 compiles the average value derived for the Pilyugin (2001a,b); Pilyugin & Thuan (2005); Pilyugin et al. (2010) techniques (P) and the average value obtained for the Kewley & Dopita (2002) (using the R_{23} , y and N_2O_2 parameters) and the Kobulnicky & Kewley (2004) method (KD).

As we see in Table 6, we find again that the results obtained using the P technique agree quite well with the oxygen abundance determined using the T_e method (offset of 0.03 dex), but the results provided by the KD method are systematically ~ 0.30 dex larger. That is exactly the opposite behavior we discussed in the previous section. This result is graphically shown in Fig. 12, which compares the T_e abundance (x -axis) with the oxygen abundances derived

from RL (black squares), the P method (blue diamonds) and the KD technique (red triangles).

Interestingly, the results provided using the KD technique have a better agreement with the values determined from RL. These are systematically ~ 0.08 dex lower than the oxygen abundances given by the KD method, but both agree within the errors (the dispersion of the RL data is ~ 0.06 dex, while the dispersion of the KD data is ~ 0.07 dex). This observational result, which was previously noted by Peimbert et al. (2007) and López-Sánchez & Esteban (2010), is actually suggesting that the real metallicities of the H II regions may be not those given by the standard T_e method based on line ratios of collisionally excited lines (CELs), but the values derived using recombination lines (RL). If this is correct, then the KD method is the only SEL technique that is reproducing those metallicities.

As we have seen along this paper, the oxygen abundances of our model spectra are well reproduced using the T_e method and SEL techniques based on photoionization models (McGaugh 1991; Kewley & Dopita 2002; Kobulnicky & Kewley 2004) but 0.3–0.5 dex higher than those estimated using the Pilyugin method. However, the situation is inverted in other works based on the comparison of observations with the results of the SEL methods (e.g. Peimbert et al. 2007; Yin et al 2007; Bresolin et al. 2009; López-Sánchez & Esteban 2010; Moustakas et al. 2010; López-Sánchez et al. 2011; Rosales-Ortega et al. 2011), i.e., the oxygen abundance in real nebulae is better reproduced by those SEL methods which rely in an empirical calibration derived using H II regions where a direct estimation of the electron temperature exists.

A possible explanation of this puzzling situation is that current photoionization models are not properly reproducing the behavior of collisionally excited lines (CELs) in real nebulae. As we said before, several authors have invoked the presence of temperature fluctuations in the neb-

ulae (Peimbert 1967) in order to address the differences between the abundances determined from CELs and RLs. Due to the strong dependence on T_e of the intensity of CELs, such temperature fluctuations cause the metallicities derived from CELs to be underestimated. Although our models consider global temperature gradients due to the hardening of the ionizing radiation field along the nebula, they did not reproduce the values of the temperature fluctuations parameter (t_2) that are necessary to make the O/H ratios determined from RLs and CELs to converge (e.g. García-Rojas & Esteban 2007). These hypothetical fluctuations must be of small spatial scale and may be produced by, for example, density inhomogeneities or localized deposition of mechanical energy that heat the gas.

Summarizing, assuming the validity of the scenario outlined above, the oxygen abundances determined from RLs and using SEL techniques based on photoionization models (KD method) would be closer to the real ones, while values determined from CELs and empirical calibrations (P method) would provide abundances about 0.2–0.3 dex lower. This should have a profound impact in current galaxy metallicity estimations that have been obtained using CELs or any empirical calibration—such as those provided by Pilyugin et al. or Pettini & Pagel (2004) – that should be revised upwards.

Finally, we want to note that the small offset of ~ 0.08 dex between the RL and the KD may be a consequence of the depletion of the oxygen atoms in dust grains. Our models are taking into account this effect, but observations suggest that oxygen depletion in dust grains may be decreasing the actual gas-phase metallicity (i.e., that derived from RL) of H II regions by 0.08–0.12 dex (Mesa-Delgado et al. 2009b; Peimbert & Peimbert 2010; Simón-Díaz & Stasińska 2011).

5 CONCLUSIONS

In this paper, we have subjected the commonly-used techniques for deriving chemical abundances to a critical “double-blind” analysis. We find that the originally-used T_e method is capable of delivering reliable abundances for O, N, S, and Cl. However, the abundances inferred for the noble gases Ne and Ar are considerably less secure, although probably for different physical reasons.

The electron temperature itself is a reasonably good indicator of the oxygen abundance, as the one is the dominant controller of the other. Error arises because there is also a weaker dependence of electron temperature on the ionization parameter. For the [O III] lines, this renders an abundance obtained from the electron temperature alone unreliable below $12 + \log(\text{O}/\text{H}) \lesssim 8.0$.

The temperature derived from the temperature-sensitive collisionally excited line ratios is always somewhat different from emission-weighted temperature defined for the principal emission lines – $\lambda 4959 + \lambda 5007$ in the case of [O III]–. This is due to the variability of electron temperature through the zone containing the ion of interest due to hardening of the radiation field and spherical divergence of the radiation field (which changes the photoelectric heating rate), and change in ionization state (which influences the cooling rate of the plasma). In this sense, the models naturally pro-

duce electron temperature gradients inside H II regions, such as those discussed in Stasińska (1978), Garnett (1992) and Stasińska (2005).

Although our models can reproduce these “global” temperature fluctuations, they cannot simulate the small-scale temperature fluctuations, such as those first discussed by Peimbert (1967) and Peimbert & Costero (1969), produced by turbulent or photo-evaporative flows, shocks, or radiation shielding by dense inclusions or globules of un-ionized material. Thus we can be reasonably sure that the true temperature fluctuations in real H II regions should be greater than the models predict. The effect of this would be to raise the temperature as measured by the [O III] ($\lambda 4959 + \lambda 5007$)/ $\lambda 4363$, which would then feed into a systematic underestimate in the oxygen abundance as delivered by the T_e method. Hence, accepting the existence of small-scale temperature fluctuations within the ionized gas and assuming that the RLs give the true metallicities, those oxygen abundances determined from CELs and empirical calibrations (P method) would provide abundances which are 0.2–0.3 dex lower than the real ones. That is because of the strong dependence of the intensity of CELs on T_e : the existence of such small-scale temperature fluctuations will cause an underestimation of the oxygen abundances using CELs lines or strong-line methods which are based on the bright nebular lines. However, oxygen abundances derived using SEL techniques based on photoionization models (such as the KD method, for which the T_e is fixed a priori) would be closer to the real values.

We have demonstrated that those SEL techniques which do not explicitly solve for the ionization parameter are quite unreliable, resulting in large scatter, systemic error, and abundance-dependent errors. To obtain the ionization parameter, we need a measurement of the [O III] $\lambda 5007$ /[O II] $\lambda \lambda 3726, 3729$ ratio. For abundances $12 + \log(\text{O}/\text{H}) \lesssim 8.0$ the oxygen abundance could then be estimated from the [O III] $\lambda 5007$ /H β line ratio. This is verified in Fig 13, which provides a new diagnostic plot for this low abundance regime.

At high abundance, the most sensitive diagnostic plot is that of [O III] $\lambda 5007$ /[O II] $\lambda 3726, 3729$ *vs.* [N II] $\lambda 6584$ /[O II] $\lambda \lambda 3726, 3729$, shown in panel c of Fig. 1. The main error here is the degree to which the N/O abundance has been accurately calibrated against the O abundance. We believe that the relationship given in Eq. 1 is probably good, since the model grids reproduce the Veilleux & Osterbrock (1987) diagnostic plots shown in Fig. 1, panels (a) and (b).

Figures 1, panel (c) and 13 thus provide a SEL technique that is applicable to the determination of both metallicity and ionization parameter over the full abundance range, with the caveat that the sensitivity of this technique is rather poor in the range $8.0 \lesssim 12 + \log(\text{O}/\text{H}) \lesssim 8.4$. In this range, common to all SEL techniques, rather accurately-measured line ratios are required.

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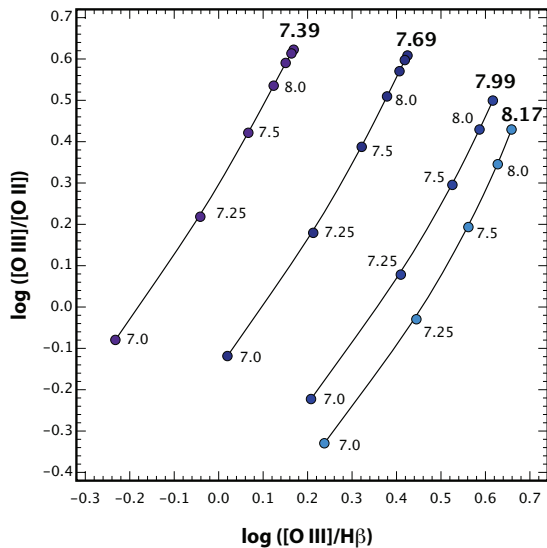


Figure 13. The model grid of $[\text{O III}] \lambda 5007/\text{H}\beta$ vs. $[\text{O III}] \lambda 5007/[\text{O II}] \lambda \lambda 3726, 3729$ at the low-abundance end. This confirms that both $12 + \log(\text{O}/\text{H})$ and $\log q$ can be derived from this diagnostic diagram for $12 + \log(\text{O}/\text{H}) \lesssim 8.0$.

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REFERENCES

- Abazajian, K., Adelman-McCarthy, J. K., Agüeros, M. A., et al. 2003, *AJ*, 126, 2081
- Aller, L. H., ed. 1984, *emphPhysics of Thermal Gaseous Nebulae* (Dordrecht: Reidel)
- Alloin D., Collin-Souffrin S., Joly M., Vigroux L., 1979, *A&A*, 78, 200
- Baldwin J., Phillips M. M., Terlevich R. J., 1981, *PASP*, 93, 5
- Bresolin, F.; Schaerer, D.; González Delgado, R. M. & Stasińska, G. 2005, *A&A*, 441, 981
- Bresolin, F., Gieren, W., Kudritzki, R.-P., Pietrzyński, G., Urbaneja, M.A. & Carraro, G. 2009, *ApJ*, 700, 309
- Brinchmann, J., Charlot, S., White, S. D. M., et al. 2004, *MNRAS*, 351, 1151
- Christensen T., Petersen L., Gammelgaard P., 1997, *A&A*, 322, 41
- Davé, R. & Oppenheimer, B.D. 2007, *MNRAS*, 374, 427
- Denicoló, G., Terlevich, R. & Terlevich, E. 2002, *MNRAS*, 330, 69
- Díaz, A.I., Terlevich, E., Pagel, B.E.J., Vílchez, J.M. & Edmunds, M. G. 1987, *MNRAS*, 226, 19
- Díaz A. I., Pérez-Montero E., 2000, *MNRAS*, 312, 130
- Dopita, M. A., Vassiliadis, E., Wood, P. R., Meatheringham, S. J., Harrington, J. P., Bohlin, R. C., Ford, H. C., Stecher, T. P., & Maran, S. P. 1997, *ApJ*, 474, 188
- Dopita, M.A., Kewley, L. J., Heisler, C.A. & Sutherland, R.S. 2000, *ApJ*, 542, 224
- Dopita, M. A., Fischera, J., Sutherland, R. S., Kewley, L. J., Leitherer, C., Tuffs, R. J., Popescu, C. C., van Breugel, W., & Groves, B. A. 2006a, *ApJS*, 167, 177
- Dopita, M. A., Fischera, J., Sutherland, R. S., Kewley, L. J., Tuffs, R. J., Popescu, C. C., van Breugel, W., Groves, B. A. & Leitherer, C. 2006b, *ApJ*, 647, 244
- Dors, O. L., Jr. & Copetti, M.V.F. 2006, *A&A*, 452, 473
- Dors, O. L., Jr., Krabbe, A., Hägele, G. F. & Pérez-Montero, E. 2011, *MNRAS*, 415, 3616
- Esteban, C. 2002, *RMxAC*, 12, 56
- Esteban, C. & Peimbert, M. 1995, *A&A*, 300, 78
- Esteban, C., Peimbert, M., Torres-Peimbert, S. & Escalante, V. 1998, *MNRAS*, 295, 401
- Esteban, C., Peimbert, M., Torres-Peimbert, S., & Rodríguez, M. 2002, *ApJ*, 581, 241
- Esteban, C., Peimbert, M., García-Rojas, J., Ruiz, M. T., Peimbert, A. & Rodríguez, M., 2004, *MNRAS*, 355, 229
- Esteban, C., Bresolin, F., Peimbert, M., García-Rojas, J., Peimbert, A. & Mesa-Delgado, A. 2009, *ApJ*, 700, 654
- Ferland, G., Fabian, A. C. & Johnston, R. 1994, *MNRAS*, 266, 399
- Ferland, G. 1995, *The Analysis of Emission Lines: A Meeting in Honor of the 70th Birthdays of D. E. Osterbrock and M. J. Seaton*, 83
- García-Rojas, J., Esteban, C., Peimbert, M., Rodríguez, M., Ruiz, M. T., & Peimbert, A. 2004, *APJS*, 153, 501
- García-Rojas, J., Esteban, C., Peimbert, A., Peimbert, M., Rodríguez, & M., Ruiz, 2005, *MNRAS*, 362, 301
- García-Rojas, J., Esteban, C., Peimbert, M., Costado, M.T., Rodríguez, M., Peimbert, A. & Ruiz, M. T. 2006, *MNRAS*, 368, 253
- García-Rojas, J. & Esteban, C., 2007, *ApJ*, 670, 457
- Garnett, D. R. 1992, *AJ*, 103, 1330
- Grevesse, N.; Asplund, M.; Sauval, A. J.; Scott, P. 2010, *Ap&SS*, 328, 179
- Hägele, G. F., Díaz, Á. I., Terlevich, E., Terlevich, R.; Pérez-Montero, E. & Cardaci, M. V. 2008, *MNRAS*, 383, 209
- Hägele, G. F., García-Benito, R., Pérez-Montero, E., Díaz, Á. I., Cardaci, M. V., Firpo, V., Terlevich, E. & Terlevich, R. 2011, *MNRAS*, 414, 272
- Izotov, Y.I., Thuan, T.X., & Lipovetski, 1994, *ApJ*, 435, 647
- Izotov, Y.I., Stasińska, G., Meynet, G., Guseva, N.G. & Thuan, T.X. 2006, *A&A*, 448, 955
- Jensen E. B., Strom S. E., 1976, *ApJ*, 209, 748
- Kewley, L.J. & Dopita, M.A. 2002, *ApJS*, 142, 35
- Kewley, L.J., & Ellison, S.E. 2008, *ApJ*, 681, 1183
- Kewley, L. J., Groves, B., Kauffmann, G., & Heckman, T. 2006, *MNRAS*, 372, 961
- Kingdon J. & Ferland G. J., 1995, *ApJ*, 442, 714
- Kobulnicky, H. A., Kennicutt, R. C., Jr., & Pizagno, J. L. 1999, *ApJ*, 514, 544
- Kobulnicky H. A. & Kewley L. J. 2004, *ApJ*, 617, 240
- Lara-López, M. A. and Bongiovanni, A. and Cepa, J. and Pérez García, A. M. and Sánchez-Portal, M. and Castañeda, H. O. and Fernández Lorenzo, M. and Pović, M. 2010a, *A&A*, 519, 31
- Lara-López, M. A. and Cepa, J. and Bongiovanni, A. and Pérez García, A. M. and Ederoclite, A. and Castañeda, H. and Fernández Lorenzo, M. and Pović, M. and Sánchez-Portal, M. 2010b, *A&A*, 521, 53
- Leitherer, C., Schaerer, D., Goldader, J.D., Delgado, R.M.G., Robert, C., Kune, D.F., de Mello, D.F., Devost, D & Heckman, T.M. et al. 1999, *ApJS*, 123, 3
- Levesque, E. M., Kewley, L. J., & Larson, K. L. 2010, *AJ*, 139, 712
- López-Sánchez, Á.R., Esteban, C., García-Rojas, J., Peimbert, M. & Rodríguez, M. 2007, *ApJ*, 656, 168
- López-Sánchez, Á.R. & Esteban, C. 2009, *A&A*, 508, 615
- López-Sánchez, Á.R. & Esteban, C. 2010, *A&A*, 517, 85
- López-Sánchez, Á.R., Mesa-Delgado, A., López-Martín, L & Esteban, C. 2011, *MNRAS*, 411, 2076
- McGaugh, S.S. 1991, *ApJ*, 380, 140
- Mesa-Delgado, A., López-Martín, L., Esteban, C., García-Rojas & Luridiana, V. 2009a, *MNRAS*, 394, 693
- Mesa-Delgado, A., Esteban, C., García-Rojas, Luridiana, V., Bautista, M., Rodríguez, M., López-Martín, L. & Peimbert, M. 2009b, *MNRAS*, 395, 855
- Mesa-Delgado, A., & Esteban, C. 2010, *MNRAS*, 405, 2651
- Mesa-Delgado, A., Núñez-Díaz, M., Esteban, C., López-Martín, L., & García-Rojas, J. 2011, *MNRAS*, 417, 420
- Mollá M., Díaz A. I., 2005, *MNRAS*, 358, 521
- Monreal-Ibero, A., Vílchez, J.M.; Walsh, J.R. & Muñoz-Tuñón, C. 2010, *A&A*, 517, 27
- Moustakas, J., Kennicutt, R.C., Jr., Tremonti, C. A., Dale, D. A., Smith, J.-D. T. & Calzetti, D. 2010, *ApJS*, 190, 233
- Osterbrock, D. E., & Ferland, G. J. 2006, *Astrophysics of gaseous nebulae and active galactic nuclei*, 2nd. ed. by D.E. Osterbrock and G.J. Ferland. Sausalito, CA: University Science Books, 2006
- Pagel B. E. J., Edmunds M. G., Blackwell D. E., Chun M. S., Smith G., 1979, *MNRAS*, 189, 95
- Peimbert, A. 2003, *ApJ*, 584, 735
- Peimbert, A., Peimbert, M. & Ruiz, M.T. 2005, *ApJ*, 634, 1056
- Peimbert A. & Peimbert M., 2010, *ApJ*, 724, 791
- Peimbert, M. 1967, *ApJ*, 150, 825
- Peimbert, M. & Costero, R. 1969, *Boletín de los Observatorios Tonantzintla y Tacubaya*, 5, 3
- Peimbert, M. & Torres-Peimbert, S. 1977, *MNRAS*, 179, 217
- Peimbert M. & Peimbert A., 2006, in Barlow M. J., Méndez R. H., eds, *Planetary Nebulae in our Galaxy and Beyond* Vol. 234 of IAU Symposium, Temperature Variations and Chemical Abundances in Planetary Nebulae. pp 227234
- Peimbert, M., Peimbert, A., Esteban, C.; García-Rojas, J., Bresolin, F., Carigi, L., Ruiz, M.T. & López-Sánchez, Á.R. 2007, *RMxAC*, 29, 72
- Pérez-Montero, E. & Díaz, Á.I. 2005, *MNRAS*, 361, 1063
- Pérez-Montero, E.; Hägele, G.F., Contini, T. & Díaz, Á.I. 2007, *MNRAS* 381, 125

- Pérez-Montero, E. & Contini, T. 2009, MNRAS, 398, 949
 Pérez-Montero, E.; García-Benito, R., Hägele, G.F. & Díaz, Á.I. 2010, MNRAS, 404, 2037
 Pettini, M. & Pagel, B.E.J. 2004, MNRAS, 348, 59
 Pilyugin, L.S. 2000, A&A, 362, 325
 Pilyugin, L.S. 2001a, A&A, 369, 594
 Pilyugin, L.S. 2001b, A&A, 374, 412
 Pilyugin, L.S. & Thuan, T.X. 2005, ApJ, 631, 231
 Pilyugin, L.S., Víchez, J.M. & Thuan, T.X. 2010, ApJ, 720, 1738
 Pilyugin, L.S. & Mattsson, L. 2011, MNRAS, 412, 1145
 Przybilla, N., Nieva, M.-F. & Butler, K. 2008, ApJ, 688, 103
 Pustilnik, S., Kniazev, A., Pramskij, A., Izotov, Y., Foltz, C., Brosch, N., Martin, J.-M. & Ugryumov, A. 2004, A&A 419, 469
 Rodríguez, M. & García-Rojas, G. 2010, ApJ, 708, 1551
 Rosales-Ortega, F.F., Díaz, A.I., Kennicutt, R.C. & Sánchez, S.F. 2011, MNRAS, 415, 2439
 Shaw, R.A. & Dufour, R.J. 1995, PASP, 107, 896
 Simón-Díaz, S. & Stasińska, G. 2011, A&A, 526, 48
 Stasińska, G. 1978, A&A, 66, 257
 Stasińska, G. 2002, Rev. Mex. Conf. Ser., 12, 62
 Stasińska, G. 2005, A&A, 434, 507
 Sutherland, R. S. & Dopita, M. A 1993, ApJS, 88, 253
 Tremonti, C.A., et al. 2004, ApJ, 613, 898
 Tsamis, Y.G., Balrow, M.J., Liu, X.-W., Danziger, I.J. & Storey, P.J. 2003, MNRAS, 338, 687
 Vázquez, G.A. & Leitherer, C. 2005, ApJ, 621, 695
 Veilleux, S., & Osterbrock, D. E. 1987, ApJS, 63, 295
 Víchez J. M., Esteban C., 1996, MNRAS, 280, 720
 Yin, S.Y., Liang, Y.C., Hammer, F., Brinchmann, J., Zhang, B., Deng, L.C. & Flores, H., 2007, A&A, 462, 535
 York, D. G., Adelman, J., Anderson, John, E., Jr., et al. 2000, AJ, 120, 1579

APPENDIX A: ELECTRON TEMPERATURES, ELECTRON DENSITIES AND IONIC ABUNDANCES FOR THE MODEL H II REGIONS.

Here we compile the derived parameters for the model H II regions. The first table, Table A1, gives the derived densities and temperatures in the various zones of the H II region. Table A2 gives the derived ionic abundances. Finally, Table A3 gives the line ratio values used to compute the oxygen abundance using the empirical calibrations as well as the ionization parameters derived for the models using the Kewley & Dopita (2002) and Kobulnicky & Kewley (2004) techniques.

Table A1. Electron density and temperature obtained from the analysis of emission line ratios of model spectra using IRAF task *temden*.

Model	n_e [cm ⁻²]					T_e High [K]				T_e Low [K]			
	[S II]	[O II]	[N I]	[Cl III]	Adopted	[O III]	[S III]	[Ar III]	Adopted	[O II]	[S II]	[N II]	Adopted
A	6	213	12	113	86	8465	8037	8389	8297	8419	6419	9274	8037
B	0	243	7	136	96	14947	12222	13307	13492	12141	7621	13658	11140
C	30	0	0	68	49	0	4425	4519	4472	0	3843	4680	4261^a
D	0	236	7	135	126	12663	11009	11706	11792	11122	7337	12820	10426
E	0	228	8	130	122	11459	10134	10732	10775	10311	7086	11802	9733
F	0	226	9	130	121	10498	9625	10159	10094	9834	7023	11254	9370
G	25	182	23	76	76	4855	4993	4908	4918	5423	4535	5660	5206
H	33	0	0	68	50	4285	4210	4236	4243	0	3737	4474	4105^a
I	19	196	21	87	80	5830	5873	5973	5892	6424	5313	6986	6241
J	35	0	0	71	53	0	3910	3768	3839	0	3633	4269	3951^a
K	0	234	7	135	125	12413	10843	11519	11591	10968	7304	12633	10301
L	0	229	8	130	122	11394	10139	10737	10756	10297	7122	11796	9738
M	0	240	7	138	128	13637	11749	12613	12666	11614	7514	13408	10845
N	6	214	12	113	86	8129	7811	8144	8028	8409	6435	9280	8041
O	21	182	21	77	75	5769	5542	5665	5658	5681	4660	5996	5445
P	0	241	7	138	128	14096	12034	12953	13027	11833	7556	13658	11015
Q	0	239	7	135	127	13718	11640	12568	12642	11462	7454	13202	10706
R	16	194	19	86	78	6225	6158	6283	6222	6549	5322	6996	6289
S	0	232	8	135	125	11890	10586	11248	11241	10661	7284	12268	10071
T	0	234	8	136	126	12181	10781	11506	11489	10804	7314	12451	10189
U	0	228	9	131	122	10525	9641	10182	10116	9913	7044	11364	9440
V	0	241	7	139	129	14562	12172	13223	13319	11840	7525	13640	11001
W	0	225	9	129	121	10537	9589	10116	10080	9748	6962	11119	9276
X	40	0	0	78	59	0	3930	0	3930	0	3695	4338	4016^a
Y	0	243	7	136	128	15740	12547	13728	14005	12101	7554	13892	11182
Z	13	194	17	90	78	6941	6580	6790	6770	6768	5356	7252	6458
AA	22	181	23	75	75	5364	5295	5338	5332	5498	4588	5788	5291
AB	0	234	8	135	125	12452	10927	11628	11669	11001	7359	12678	10346
AC	15	194	17	89	78	6741	6483	6654	6626	6712	5378	7189	6426
AD	0	243	6	138	129	15240	12497	13644	13793	12078	7563	13888	11176
AE	0	238	7	139	128	13248	11484	12330	12354	11363	7459	13115	10645
AF	0	222	10	125	119	10289	9308	9814	9803	9470	6799	10740	9003
AG	5	212	12	113	85	8765	8182	8545	8497	8477	6389	9318	8061

^a *ionic* task does not allow electron temperatures lower than 5,000 K for the low ionization ions, so we adopted $T_e(\text{low}) = 5,000$ K for these cases.

Table A2. Ionic abundances derived for the results of the photoionization models using IRAF task *ionic* considering the electron temperature derived from all the high-ionization species, $T_e(\text{high})$, and the low ionization species, $T_e(\text{low})$, separately. All numbers are expressed in the form of $12 + \log(X/H)$. The label *disp.* indicates the dispersion of the data when considering the values given by different emission lines.

Model	O ⁺ /H ⁺	O ⁺⁺ /H ⁺	N ⁺ /H ⁺	S ⁺ /H ⁺	S ⁺⁺ /H ⁺	Ne ⁺⁺ /H ⁺	Ar ⁺ /H ⁺	Ar ⁺⁺ /H ⁺	Cl ⁺⁺ /H ⁺	Cl ³⁺ /H ⁺
A	8.717	8.126	7.333	6.327	7.074	7.683	6.458	4.364	4.272	3.578
<i>disp.</i>	0.031	0.016	0.085	0.095	0.018	0.071	0.007	0.001	0.001	0.090
B	7.283	7.609	5.635	4.928	5.873	6.941	5.326	4.915	2.810	3.251
<i>disp.</i>	0.036	0.049	0.084	0.122	0.036	0.071	0.006	0.000	0.001	0.090
C	8.873	7.101	8.257	7.224	7.365	7.347	6.586	0.000	4.575	2.487
<i>disp.</i>	0.115	4.100	0.070	0.189	0.011	0.071	0.013	0.000	0.001	0.090
D	7.489	6.809	5.840	5.208	5.727	6.468	5.132	2.969	2.882	2.502
<i>disp.</i>	0.029	0.039	0.091	0.119	0.028	0.071	0.004	0.001	0.000	0.090
E	8.173	6.915	6.606	6.003	6.282	6.884	5.703	2.871	3.517	2.695
<i>disp.</i>	0.029	0.037	0.092	0.113	0.028	0.071	0.003	0.001	0.000	0.090
F	8.252	8.127	6.717	5.796	6.752	7.544	6.157	4.787	3.868	3.666
<i>disp.</i>	0.026	0.026	0.091	0.106	0.625	0.071	0.003	0.001	0.000	0.090
G	9.089	8.315	8.038	6.764	7.525	7.803	6.747	4.544	4.725	3.535
<i>disp.</i>	0.038	0.018	0.079	0.088	0.015	0.071	0.003	0.000	0.000	0.090
H	8.710	7.821	8.152	7.003	7.564	7.658	6.802	3.624	4.771	3.019
<i>disp.</i>	0.144	0.017	0.122	0.219	0.008	0.071	0.002	0.000	0.001	0.090
I	8.963	8.309	7.709	6.447	7.240	7.754	6.559	4.909	4.416	3.763
<i>disp.</i>	0.061	0.076	0.137	0.058	0.048	0.071	0.041	0.000	0.000	0.090
J	8.372	8.105	7.892	6.667	7.676	7.710	6.781	4.095	4.855	3.231
<i>disp.</i>	0.170	4.680	0.178	0.238	0.022	0.071	0.028	0.002	0.001	0.090
K	7.778	7.081	6.155	5.502	6.028	6.754	5.433	3.223	3.194	2.765
<i>disp.</i>	0.026	0.039	0.087	0.120	0.028	0.071	0.004	0.001	0.000	0.090
L	8.313	7.388	6.758	6.015	6.505	7.129	5.905	3.456	3.711	3.024
<i>disp.</i>	0.028	0.035	0.091	0.112	0.027	0.071	0.002	0.001	0.000	0.090
M	7.536	7.470	5.884	5.167	6.022	6.876	5.431	4.191	3.036	3.096
<i>disp.</i>	0.029	0.038	0.089	0.120	0.029	0.071	0.003	0.001	0.000	0.090
N	8.481	8.378	7.079	6.029	7.065	7.804	6.458	5.041	4.216	3.858
<i>disp.</i>	0.031	0.010	0.085	0.094	0.016	0.071	0.010	0.001	0.001	0.090
O	9.185	7.426	8.201	7.180	7.307	7.405	6.597	3.204	4.517	2.846
<i>disp.</i>	0.025	0.000	0.047	0.119	0.034	0.071	0.018	0.000	0.000	0.090
P	7.218	7.183	5.548	4.862	5.712	6.581	5.124	3.946	2.705	2.818
<i>disp.</i>	0.030	0.039	0.089	0.122	0.030	0.071	0.003	0.001	0.000	0.090
Q	7.644	7.880	6.022	5.232	6.224	7.223	5.667	5.094	3.210	3.505
<i>disp.</i>	0.029	0.042	0.090	0.120	0.032	0.071	0.003	0.001	0.000	0.090
R	8.984	8.338	7.788	6.615	7.368	7.862	6.711	4.579	4.579	3.652
<i>disp.</i>	0.034	0.001	0.082	0.089	0.008	0.071	0.009	0.000	0.000	0.090
S	8.110	7.855	6.516	5.674	6.535	7.300	5.927	4.379	3.643	3.414
<i>disp.</i>	0.027	0.033	0.090	0.112	0.026	0.071	0.001	0.001	0.000	0.090
T	7.896	8.008	6.297	5.444	6.472	7.373	5.896	5.017	3.510	3.603
<i>disp.</i>	0.027	0.033	0.090	0.114	0.027	0.071	0.001	0.001	0.000	0.090
U	8.067	8.227	6.531	5.607	6.692	7.596	6.117	5.290	3.755	3.812
<i>disp.</i>	0.026	0.026	0.091	0.107	0.023	0.071	0.003	0.002	0.000	0.090
V	7.349	7.579	5.698	4.983	5.921	6.924	5.369	4.758	2.868	3.221
<i>disp.</i>	0.030	0.043	0.089	0.123	0.033	0.071	0.004	0.001	0.000	0.090
W	8.481	7.804	6.972	6.106	6.759	7.406	6.150	3.992	3.946	3.351
<i>disp.</i>	0.026	0.029	0.091	0.107	0.024	0.071	0.002	0.001	0.000	0.090
X	8.199	7.632	7.698	6.368	7.457	7.095	6.138	3.674	4.599	2.834
<i>disp.</i>	0.152	4.407	0.159	0.227	0.000	0.071	0.607	0.003	0.000	0.090
Y	7.004	7.312	5.337	4.657	5.561	6.641	5.015	4.628	2.484	2.956
<i>disp.</i>	0.031	0.053	0.088	0.126	0.039	0.071	0.009	0.001	0.000	0.090
Z	9.179	7.161	8.024	7.151	7.158	7.350	6.502	2.898	4.387	2.758
<i>disp.</i>	0.036	0.025	0.087	0.098	0.020	0.071	0.002	0.001	0.001	0.090
AA	9.292	8.103	8.240	7.056	7.484	7.746	6.769	4.105	4.698	3.336
<i>disp.</i>	0.037	0.008	0.083	0.090	0.006	0.071	0.001	0.000	0.000	0.090
AB	8.016	7.570	6.405	5.639	6.349	7.093	5.742	3.905	3.478	3.168
<i>disp.</i>	0.028	0.036	0.090	0.116	0.027	0.071	0.002	0.001	0.000	0.090
AC	9.138	7.906	7.964	6.903	7.303	7.666	6.654	3.869	4.529	3.294
<i>disp.</i>	0.035	0.017	0.084	0.093	0.016	0.071	0.003	0.001	0.001	0.090
AD	7.042	7.285	5.374	4.690	5.604	6.625	5.055	4.492	2.535	2.931
<i>disp.</i>	0.031	0.047	0.089	0.125	0.035	0.071	0.005	0.001	0.000	0.090
AE	7.766	7.814	6.141	5.348	6.304	7.189	5.722	4.717	3.322	3.422
<i>disp.</i>	0.028	0.037	0.090	0.118	0.029	0.071	0.002	0.001	0.000	0.090
AF	8.585	7.112	7.118	6.439	6.669	7.157	6.084	3.024	3.914	2.857
<i>disp.</i>	0.027	0.033	0.091	0.107	0.026	0.071	0.001	0.001	0.000	0.090
AG	8.851	7.570	7.503	6.639	7.005	7.456	6.400	3.515	4.239	3.153
<i>disp.</i>	0.033	0.024	0.085	0.098	0.021	0.071	0.004	0.001	0.001	0.090

Table A3. Parameters used to compute the oxygen abundance using empirical calibrations.

Model	R_2	R_3	R_{23}	P	y	$\log N_2$	$\log O_3N_2$	$\log O_2N_2$	S_2	S_{23}	$\log q_{KD02}^a$	$\log q_{KK04}^b$
A	4.169	2.447	6.616	0.370	-0.231	-0.640	1.026	1.387	0.623	2.445	7.72	7.32
B	0.668	3.547	4.215	0.842	0.725	-1.928	2.481	1.880	0.060	0.367	8.40	7.87
C	0.350	0.009	0.359	0.025	-1.583	-0.422	-1.619	0.093	1.001	1.560	6.69	6.47
D	0.828	0.389	1.216	0.319	-0.329	-1.805	1.397	1.850	0.098	0.269	7.16	7.13
E	2.981	0.375	3.356	0.112	-0.901	-1.120	0.691	1.721	0.510	1.028	7.30	6.87
F ^c	3.038	5.002	8.040	0.622	0.217	-1.057	1.754	1.667	0.284	1.960	7.81	7.65 , 7.74
G	0.596	0.287	0.883	0.325	-0.318	-0.665	0.120	0.567	0.352	1.452	7.47	7.83
H	0.255	0.032	0.288	0.112	-0.898	-0.500	-0.993	0.034	0.620	1.337	6.78	7.52
I	1.038	1.178	2.215	0.532	0.055	-0.811	0.879	0.954	0.266	1.604	7.78	8.09
J	0.121	0.030	0.152	0.200	-0.603	-0.738	-0.783	-0.051	0.294	0.875	6.94	8.58
K	1.603	0.690	2.293	0.301	-0.366	-1.491	1.328	1.823	0.191	0.522	7.08	7.14
L	4.134	1.111	5.245	0.212	-0.571	-0.969	1.013	1.712	0.524	1.384	7.61	7.09
M	1.085	2.203	3.288	0.670	0.308	-1.715	2.056	1.877	0.098	0.481	8.00	7.56
N	2.429	3.860	6.289	0.614	0.201	-0.893	1.477	1.405	0.313	1.959	7.90	7.85
O	1.538	0.091	1.629	0.056	-1.227	-0.303	-0.739	0.617	1.397	2.610	6.87	6.99
P	0.554	1.229	1.782	0.689	0.346	-2.033	2.120	1.903	0.050	0.247	7.54	7.54
Q	1.318	5.591	6.909	0.809	0.627	-1.593	2.338	1.840	0.110	0.721	8.31	7.90
R	1.823	1.109	2.932	0.378	-0.216	-0.565	0.608	0.953	0.532	2.210	7.55	7.73
S	3.000	3.764	6.764	0.556	0.099	-1.169	1.742	1.773	0.260	1.262	7.94	7.53
T	1.925	5.728	7.653	0.748	0.473	-1.375	2.131	1.787	0.158	1.061	8.34	7.81
U ^c	2.052	6.349	8.401	0.756	0.490	-1.233	2.034	1.672	0.188	1.356	8.37	7.85 , 7.97
V	0.744	3.232	3.976	0.813	0.638	-1.883	2.390	1.882	0.066	0.400	8.32	7.80
W ^c	4.906	2.354	7.260	0.324	-0.319	-0.814	1.183	1.632	0.565	1.919	7.62	7.28 , 7.34
X	0.079	0.012	0.092	0.134	-0.812	-0.961	-1.002	-0.012	0.145	0.550	7.12	...
Y	0.357	1.968	2.325	0.846	0.741	-2.242	2.533	1.922	0.033	0.192	8.42	7.82
Z	3.360	0.107	3.467	0.031	-1.495	-0.288	-0.683	0.941	2.049	3.382	6.90	6.76
AA	1.085	0.275	1.361	0.202	-0.595	-0.434	-0.128	0.597	0.742	2.104	7.01	7.49
AB	2.707	2.180	4.887	0.446	-0.094	-1.257	1.593	1.816	0.257	0.959	7.91	7.36
AC	2.975	0.545	3.520	0.155	-0.737	-0.354	0.088	0.954	1.127	2.874	7.44	7.22
AD	0.389	1.796	2.185	0.822	0.664	-2.206	2.458	1.923	0.035	0.206	8.34	7.76
AE	1.710	4.535	6.245	0.726	0.424	-1.486	2.140	1.846	0.141	0.844	8.11	7.73
AF ^c	5.403	0.431	5.834	0.074	-1.098	-0.709	0.342	1.569	1.125	2.169	8.18	6.82 , 6.88
AG	5.738	0.737	6.475	0.114	-0.891	-0.465	0.330	1.351	1.296	2.945	7.15	6.98

^a Value derived for the q parameter (in units of cm s^{-1}) obtained using the optimal calibration given by Kewley & Dopita (2002).^b Value derived for the q parameter (in units of cm s^{-1}) obtained using the iterative procedure described in Kobulnicky & Kewley (2004).^c For these models, for which we derived $12+\log(\text{O}/\text{H})\sim 8.4$ following the T_e method, we list the results of the empirical calibrations considering both the low and the high metallicity branches. The q_{KK04} parameters listed here for these models are for the low and high metallicity branches, respectively.